

Stellar Astronomy and Astrophysics (SS07):

Exercise 3 (for June 14, 2007)

1. Period-Luminosity relation for Cepheids:

In the lecture we derived

1. the period-density relation $P \propto \rho^{-1/2}$,
 2. a Mass-Luminosity relation $L \propto M^\alpha$,
 3. and have shown that the instability strip for pulsation is restricted to a narrow strip of almost constant T_{eff} in the HRD.
- a) If we assume that for Cepheids $\alpha = 4$ (a somewhat higher value than for main-sequence stars), show that a period-luminosity relation ($\log L$ vs. $\log P$) exists. Hint: Use Stefan-Boltzmann's law.
- b) Assume that the instability strip is not at $T_{\text{eff}} = \text{const.}$ but more accurately given by $\log L = \beta \log T_{\text{eff}} + \delta$. How does the period-luminosity function look like in this case?
- c) Is this relation sufficient to determine the distance of a Cepheid from measuring its period and an apparent magnitude?

2. Hertzsprung-Russell Diagram for Simple Model Stars

Consider a family of chemically homogeneous stars that are similar in every respect except for their masses M and radii R . Using a dimensionless radius variable, $x = r/R$, we can define similarity functions $F(x)$ such that

$$\rho(r) = \frac{M}{R^3} F_\rho(x) \quad \text{and} \quad m(r) = M F_m(x).$$

Assume an ideal equation of state for the stellar material, i.e. $P = \rho k_B T / \mu_m$ with Boltzmann's constant k_B and mean molecular mass $\mu_m = \mu m_p$ (where μ = mean molecular weight and m_p = proton mass) and with P and T being pressure and temperature. Assume furthermore that energy is transported radiatively with an opacity obeying Kramer's law ($\kappa \propto \rho T^{-7/2}$) and that nuclear energy is generated by the PP chain where the energy production scales as $\varepsilon_{\text{PP}} \propto \rho^2 T^4$.

- a) Use the fundamental equations of stellar structure as outlined in the lecture to derive the following scaling relations for pressure P , temperature T , and energy flux due to radiative transport L_{rad} as well as due to nuclear fusion L_{nuc} :

$$\begin{aligned} P(r) &= \frac{M^2}{R^4} F_P(x), \\ T(r) &= \frac{M}{R} F_T(x), \\ L_{\text{rad}}(r) &= \frac{M^{5.5}}{R^{0.5}} F_{\text{rad}}(x), \\ L_{\text{nuc}}(r) &= \frac{M^6}{R^7} F_{\text{nuc}}(x), \end{aligned}$$

where again the F -functions are common to all family members.

- b) Note that the energy flux transported by radiative diffusion increases slowly while the flux generated by nuclear fusion rises rapidly as the star contracts. Sketch L_{rad} and L_{nuc} as function of radius. Find the radius and luminosity as function of total mass at which the PP chain can produce enough energy to compensate the radiative losses at the surface. This is when the star reaches a quasi-equilibrium state: the stellar main sequence.
- c) Demonstrate that all stars of the homologous family in this phase lie on a line in the Hertzsprung-Russell diagram with

$$L \propto T_{\text{E}}^{4.12}$$

with T_{E} being the effective temperature at the surface.