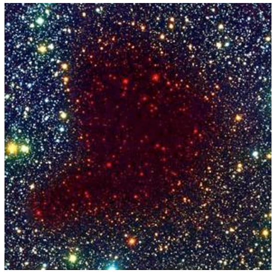
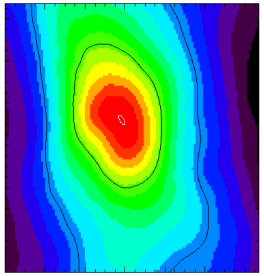
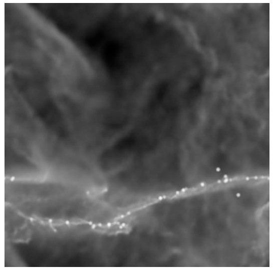
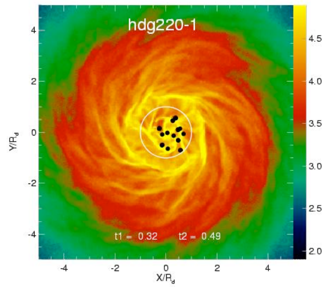
Star Formation





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thanks to ...

- many thanks to
  - Robi Banerjee (ITA)
  - Paul Clark (ITA)
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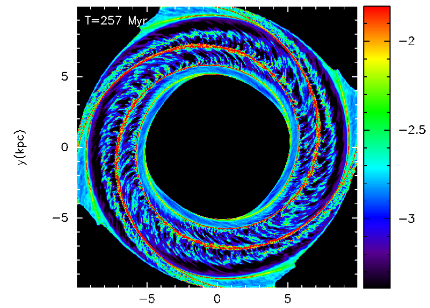
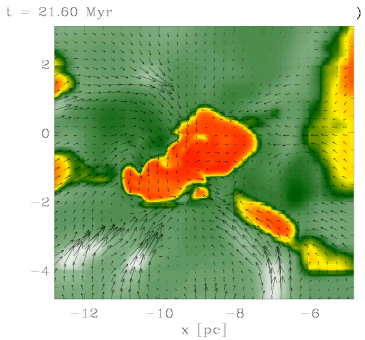
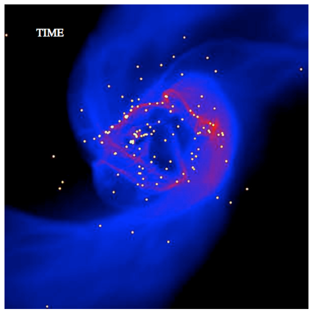
Agenda

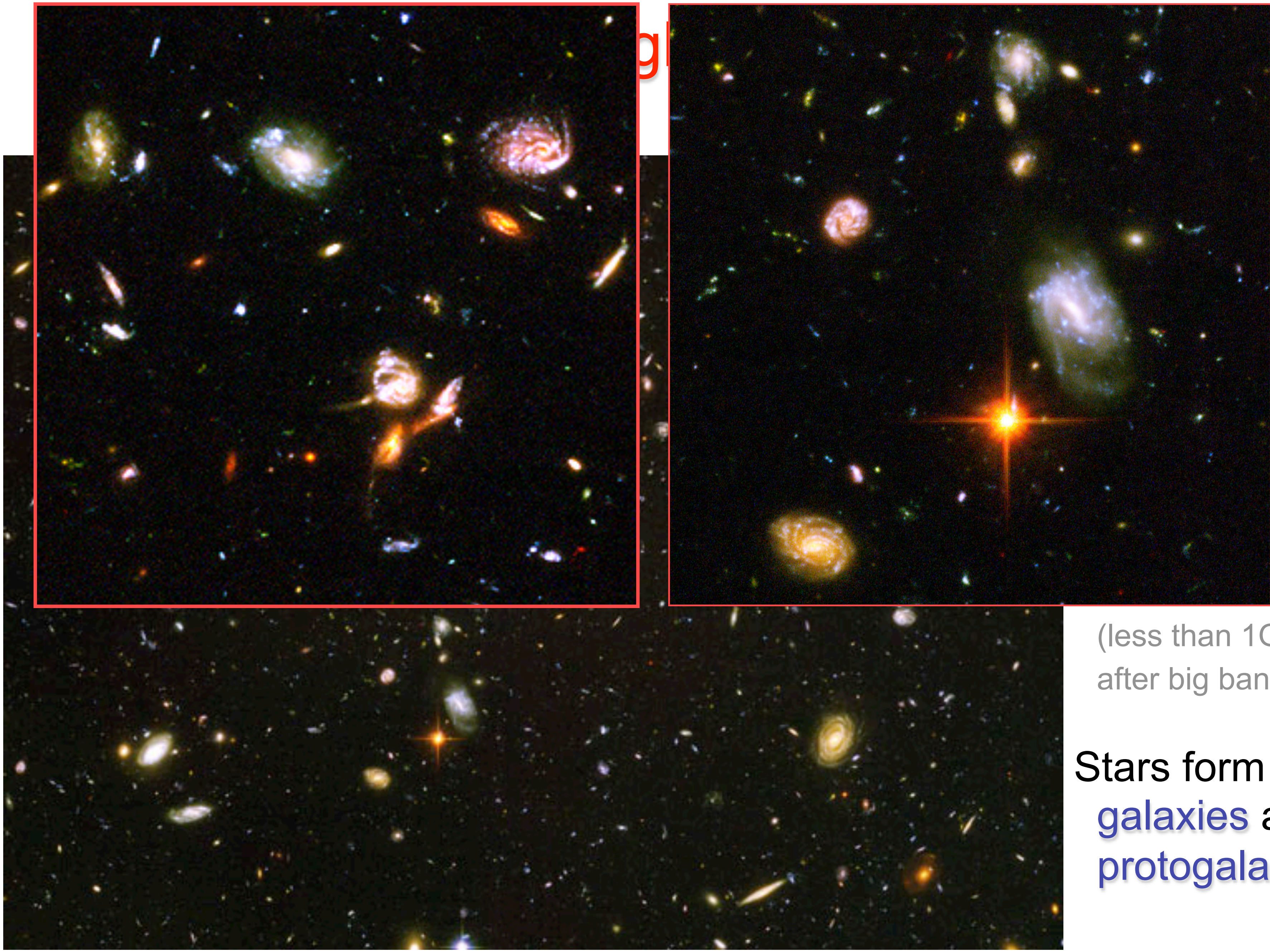
- phenomenology
  - what we need to explain
- dynamic star formation theory
  - gravity vs. turbulence (and all the rest)
- examples and predictions
  - formation of molecular clouds in galactic disks (H_2 & CO chemistry)
  - universal IMF: importance of turbulence and thermodynamics







phenomenology



(Hubble Ultra-Deep Field, from HST Web site)

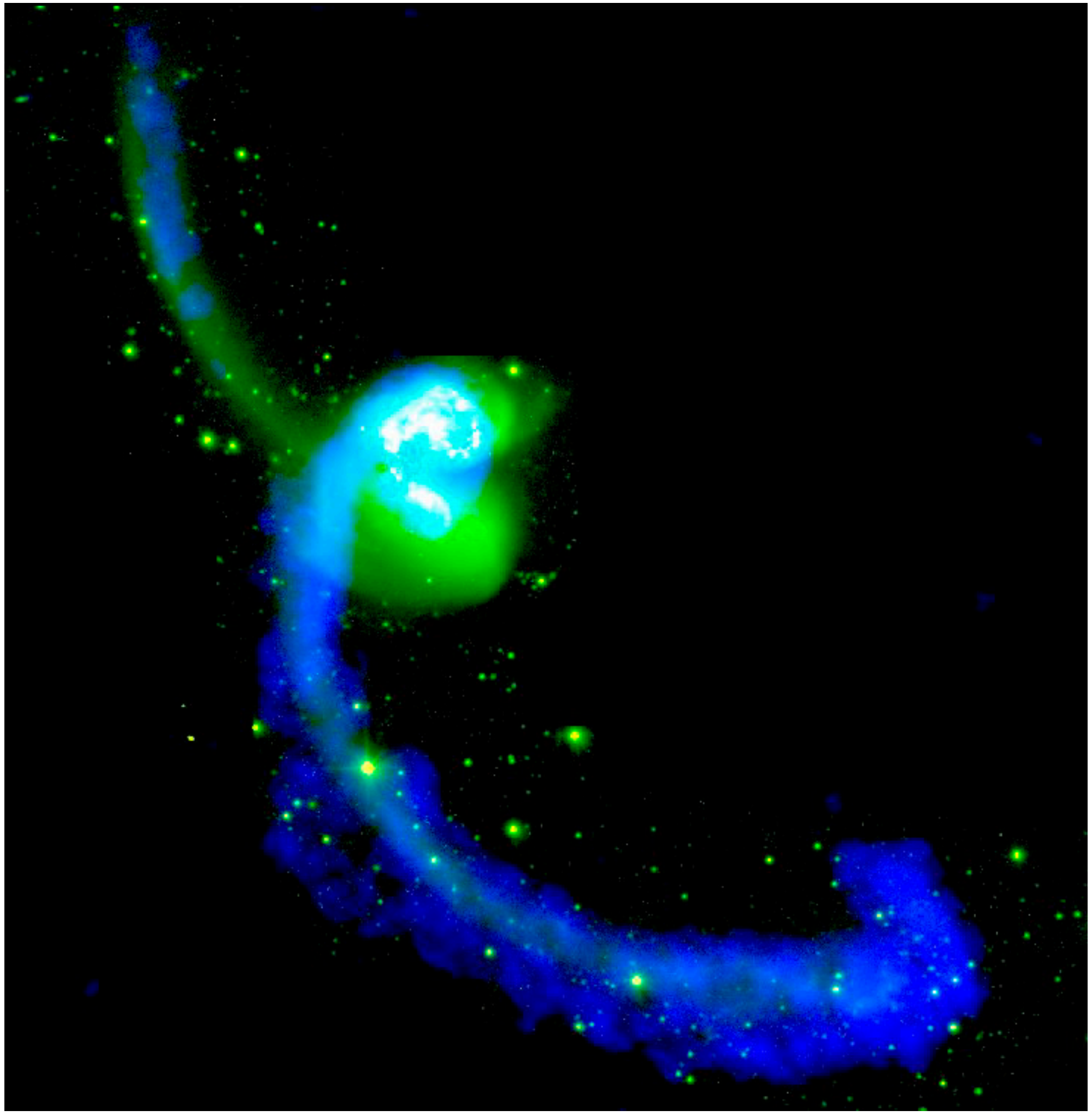
tion

rly

(less than 1Gyr after big bangl)

Stars form in galaxies and protogalaxies

Star formation in interacting galaxies:

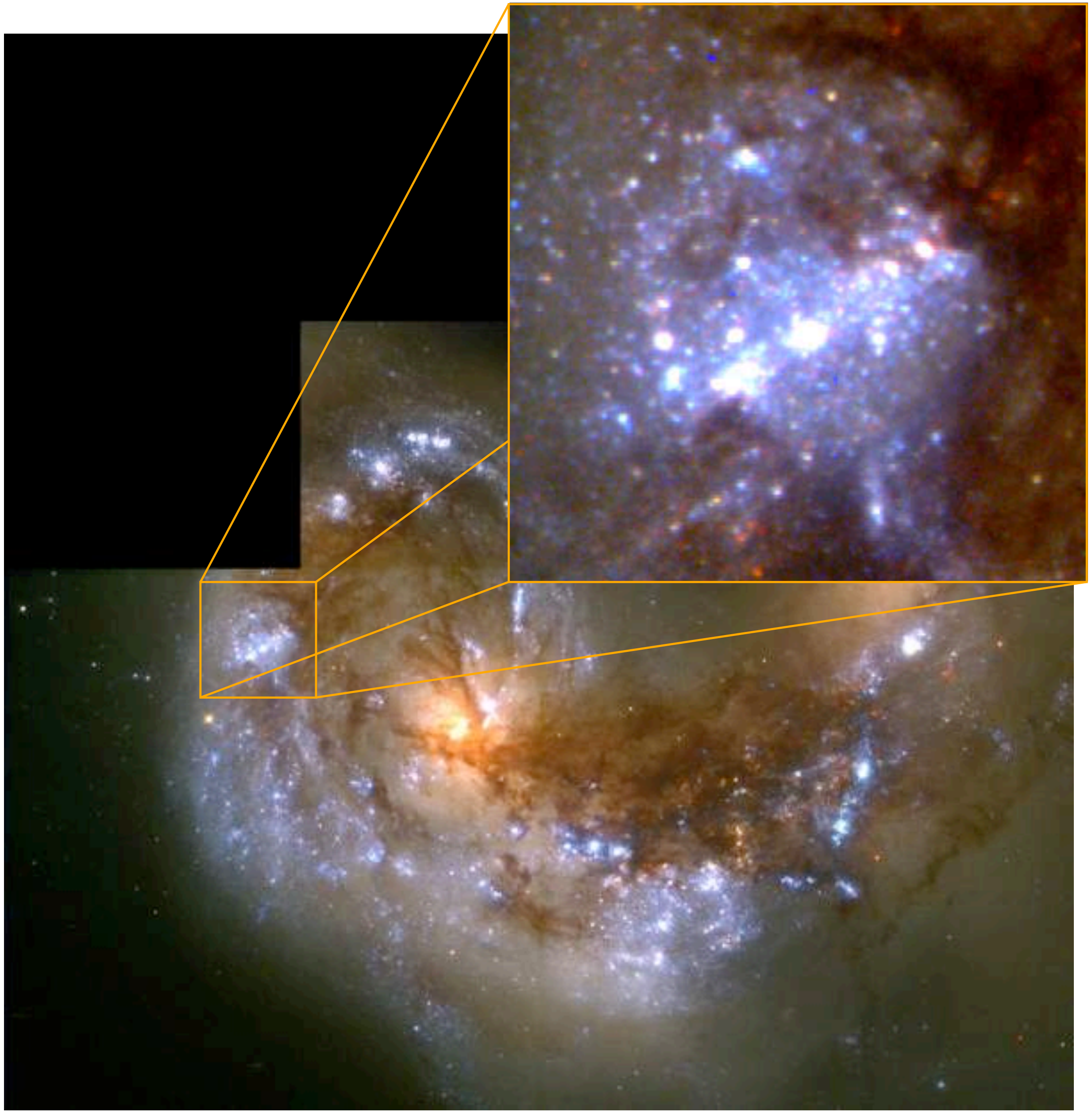


(from the Chandra Webpage)

Antennae galaxy

- *NGC4038/39*
- *distance: 19.2Mpc*
- *vis. Magn: 11.2*
- *optical: white, green*
- *radio: blue*

Star formation in interacting galaxies:



(HST: Whitmore & Schweizer 1997)

Antennae galaxy

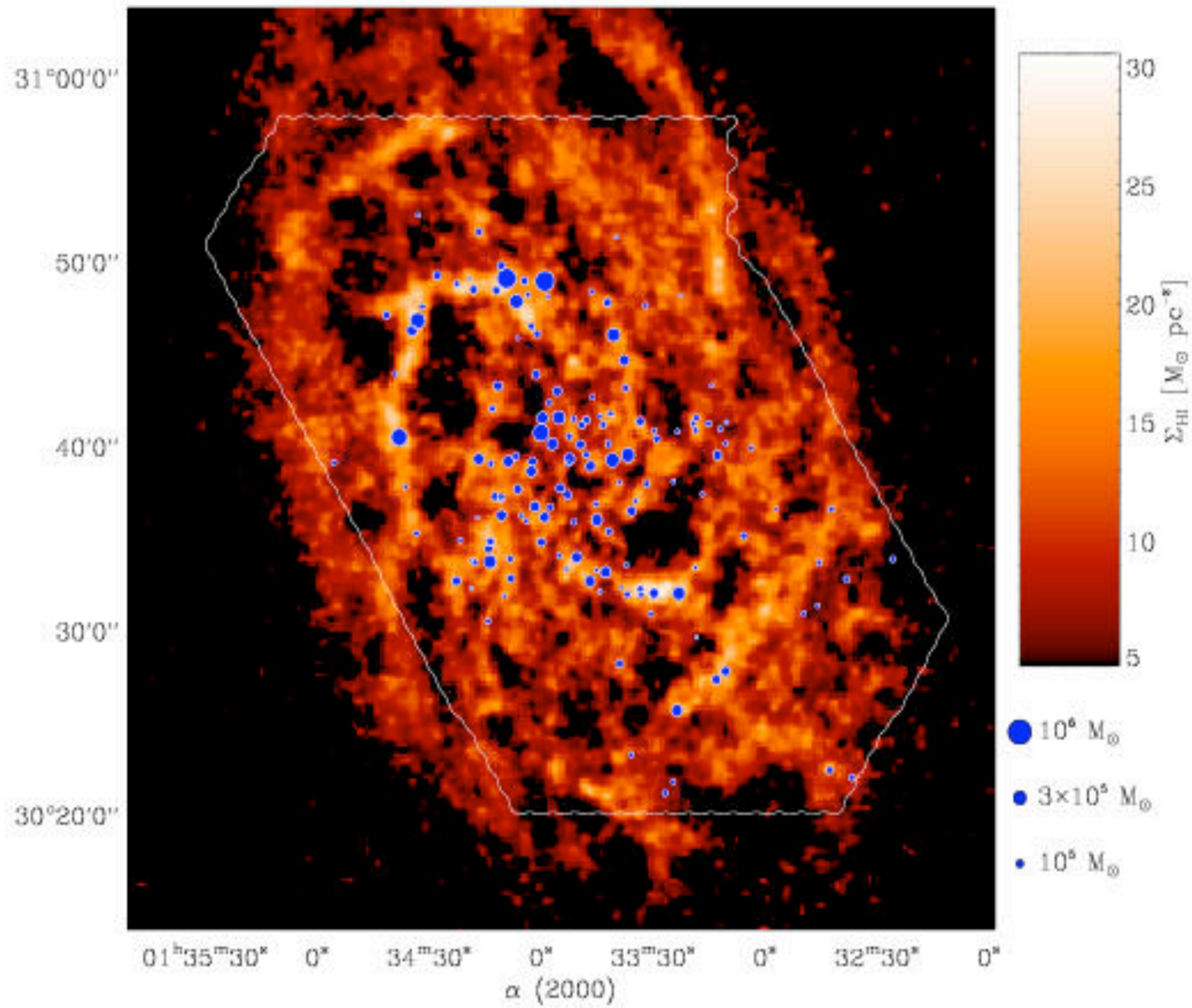
- Star formation burst in interacting (merging) galaxies
- Strong perturbation SF in tidal “tales”
- Large-scale gravitational motion determines SF
- Stars form in “knobs” (i.e. superclusters)

young stars in spiral galaxies

(NGC 4622 from the Hubble Heritage Team)

- Star formation *always* is associated with *clouds of gas and dust*.
- Star formation is essentially a *local phenomenon* (on ~pc scale)
- HOW is star formation is *influenced* by *global* properties of the galaxy?

correlation between H_2 and HI

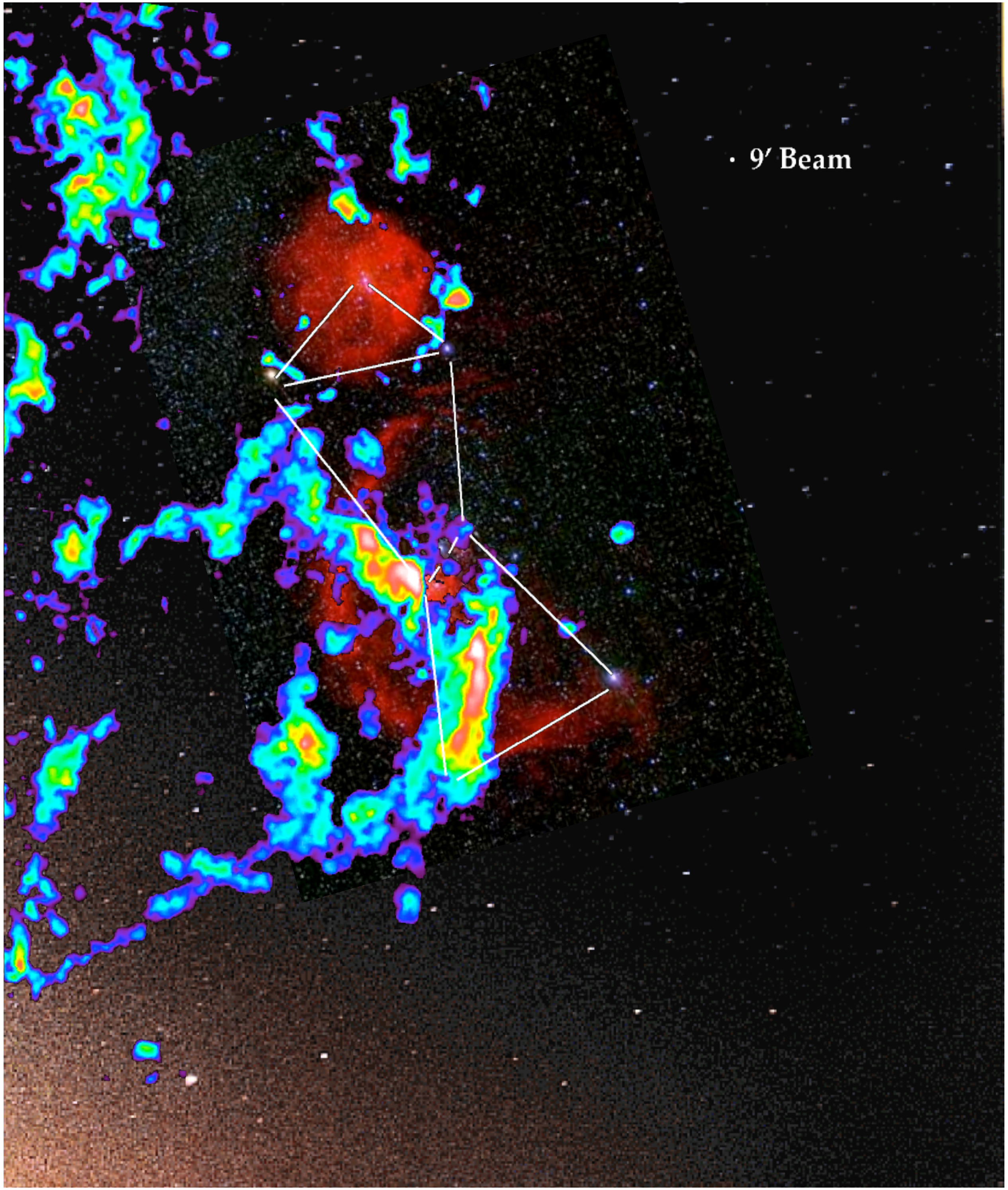


Compare H_2 - HI in M33:

- H_2: BIMA-SONG Survey, see Blitz et al.
- HI: Observations with Westerbork Radio T.

H_2 clouds are seen in regions of high HI density
(in spiral arms and filaments)

(Deul & van der Hulst 1987, Blitz et al. 2004)

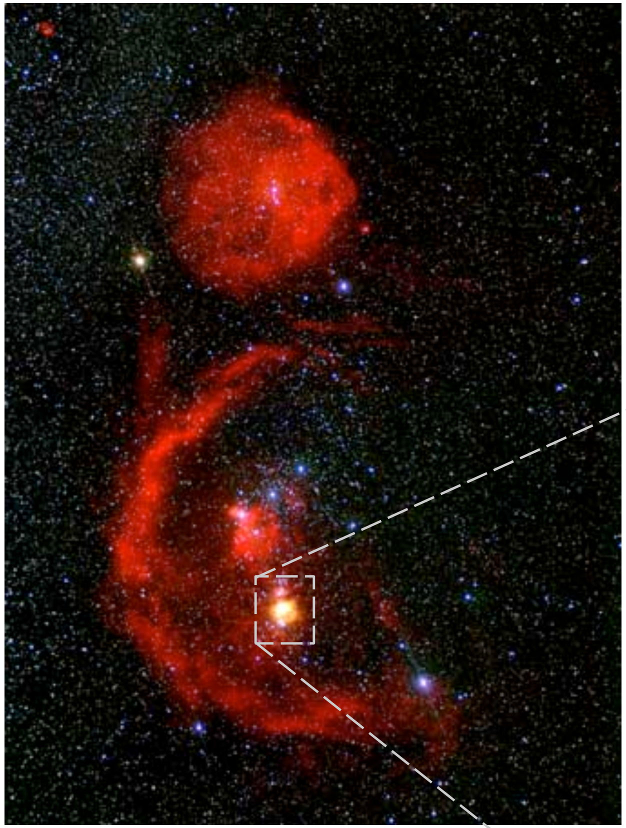


Star formation in Orion

We see

- *Stars* (in visible light)
- Atomic hydrogen (in Hα -- red)
- Molecular hydrogen H_2 (radio emission -- color coded)

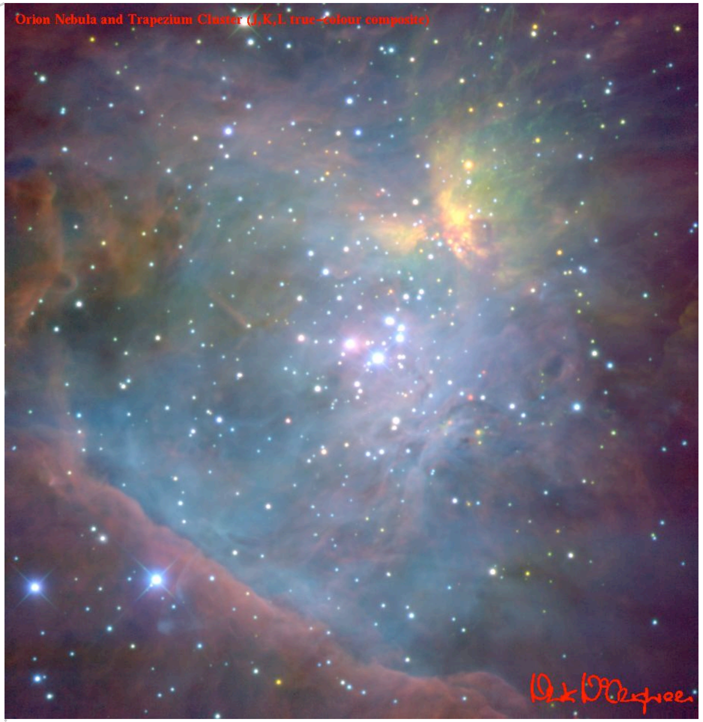
Local star forming region: The Trapezium Cluster in Orion



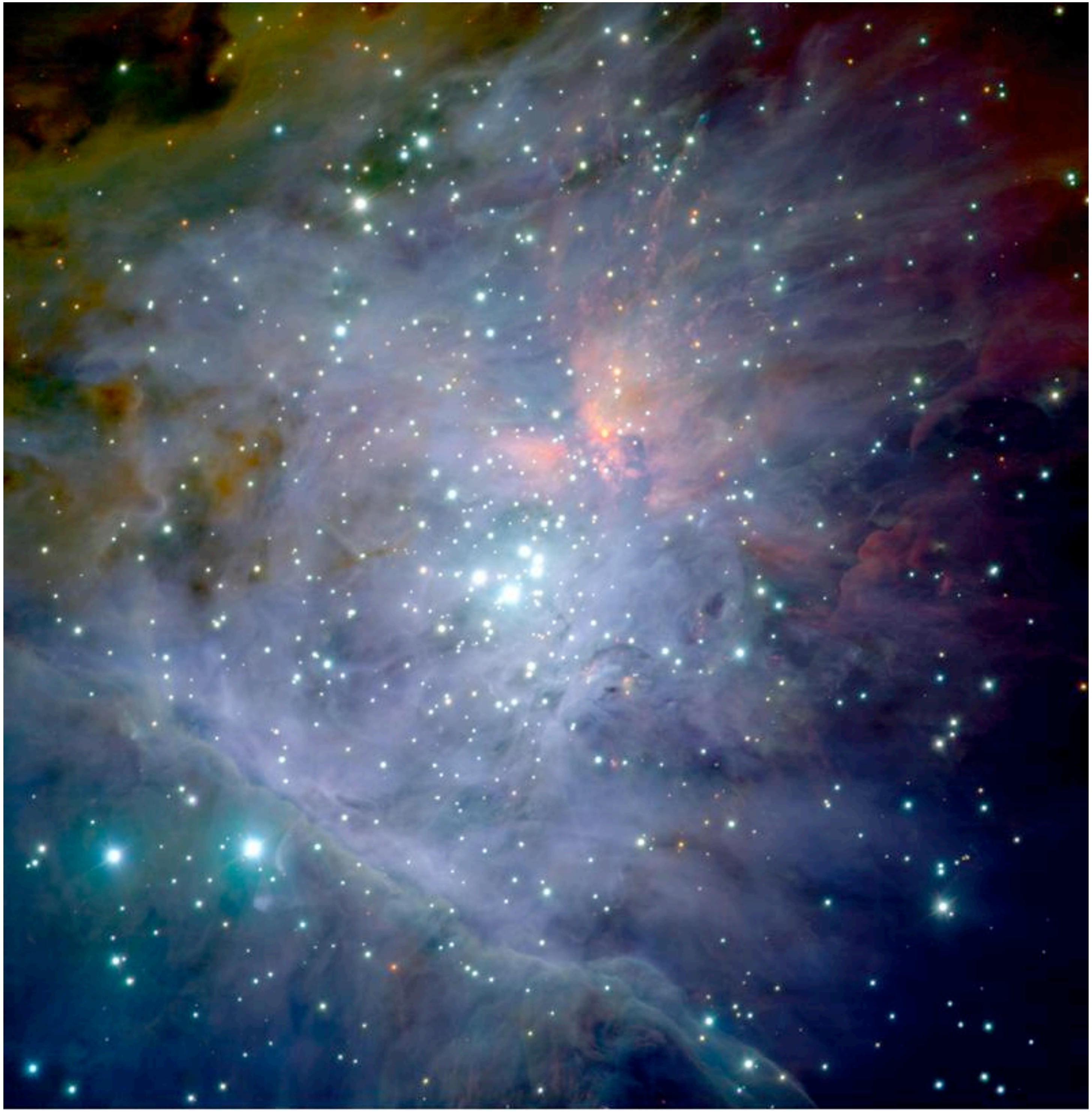
Orion molecular cloud

The Orion molecular cloud is the birth- place of several young embedded star clusters.

The Trapezium cluster is only visible in the IR and contains about 2000 newly born stars.



Trapezium cluster



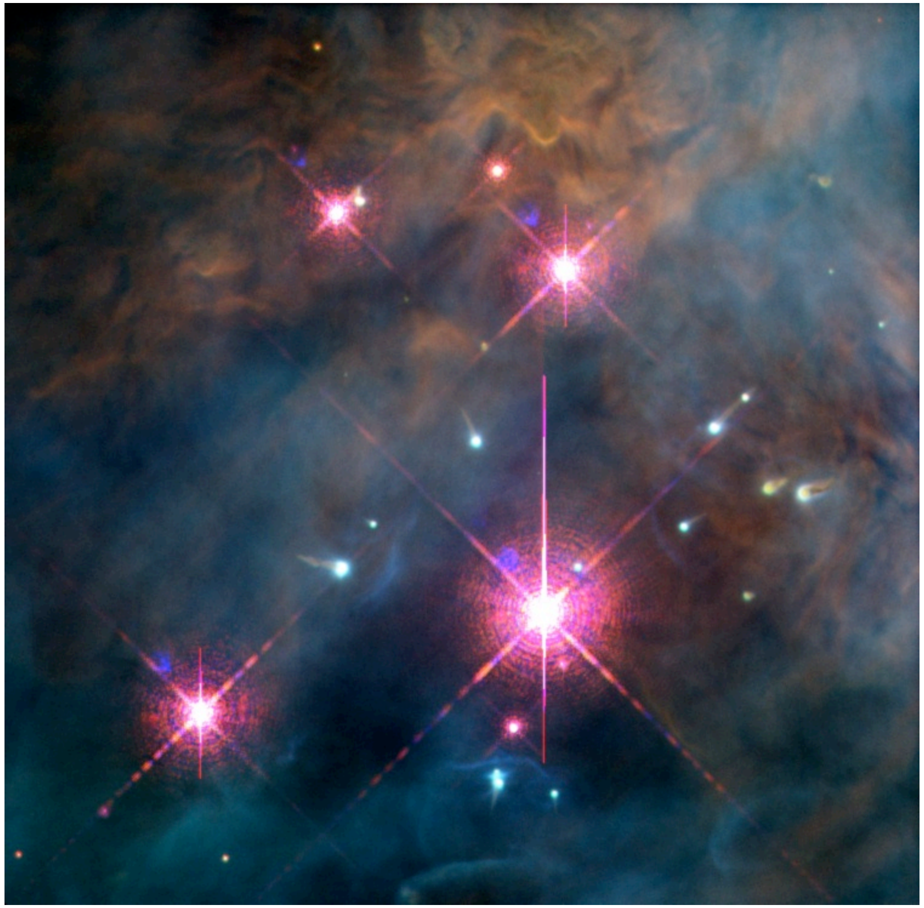
Trapezium Cluster

(detail)

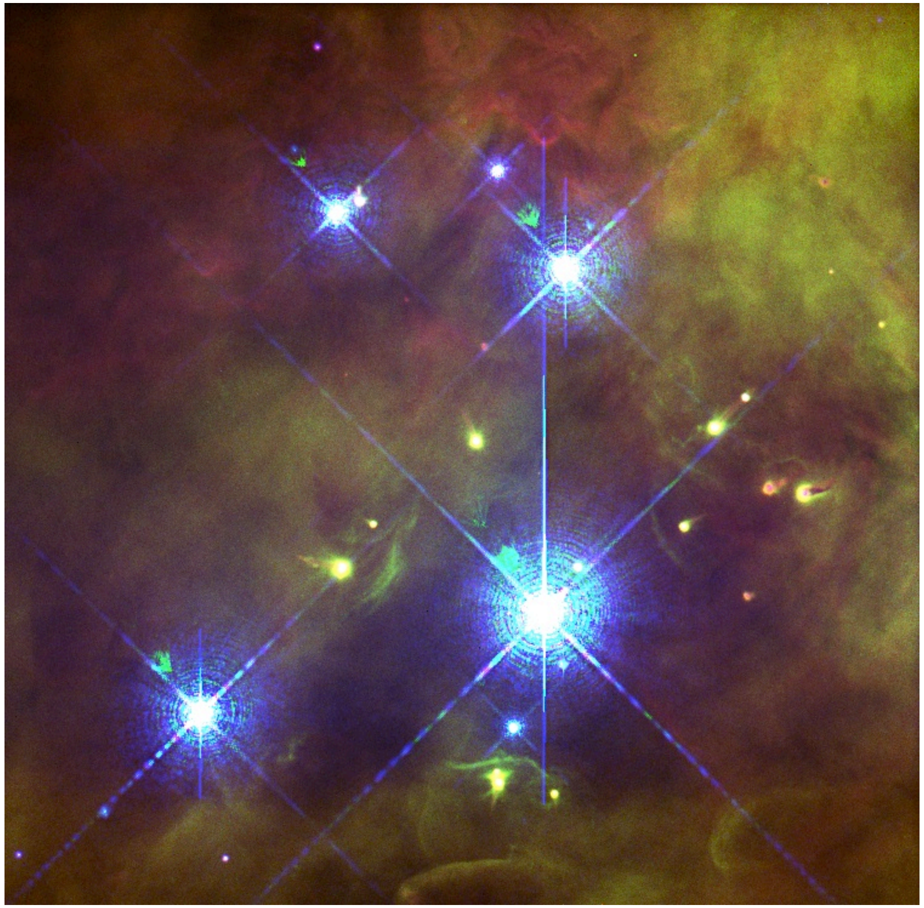
- stars form in clusters
- stars form in molecular clouds
- (proto)stellar feedback is important

(color composite J,H,K by M. McCaughrean, VLT, Paranal, Chile)

Trapezium Cluster: Central Region



Ionizing radiation from central star
Θ1C Orionis

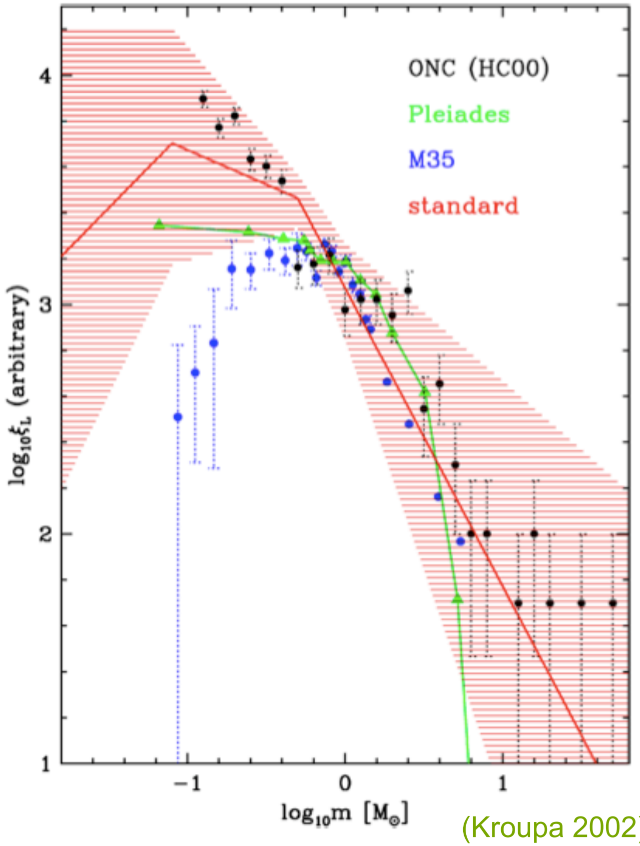


Proplyds: Evaporating ``protoplanetary´´ disks around young low-mass protostars

(images: Doug Johnstone et al.)

stellar mass fuction

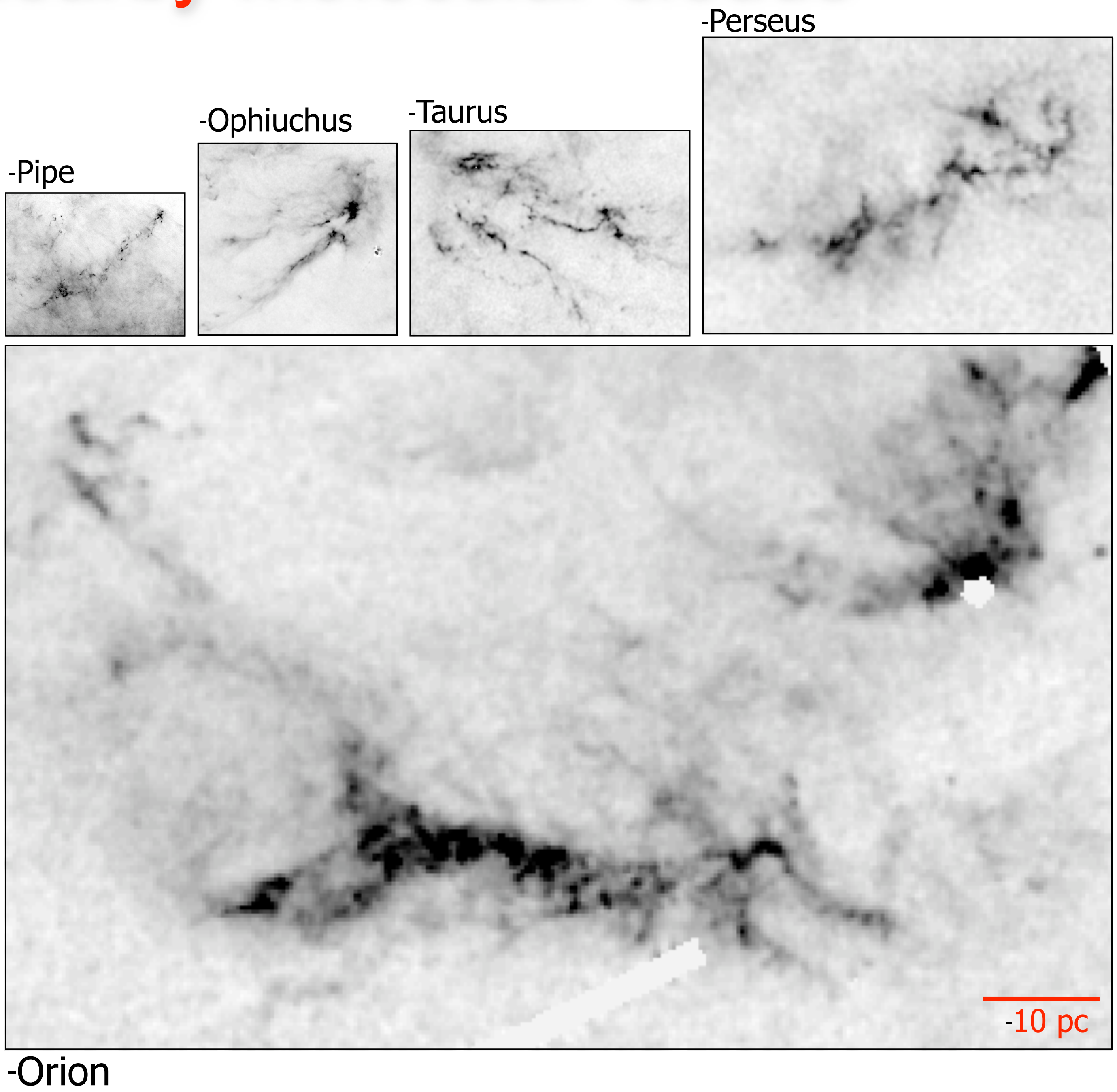
- stars seem to follow a universal mass function at birth --> IMF



(Kroupa 2002)

Orion, NGC 3603, 30 Doradus
(Zinnecker & Yorke 2007)

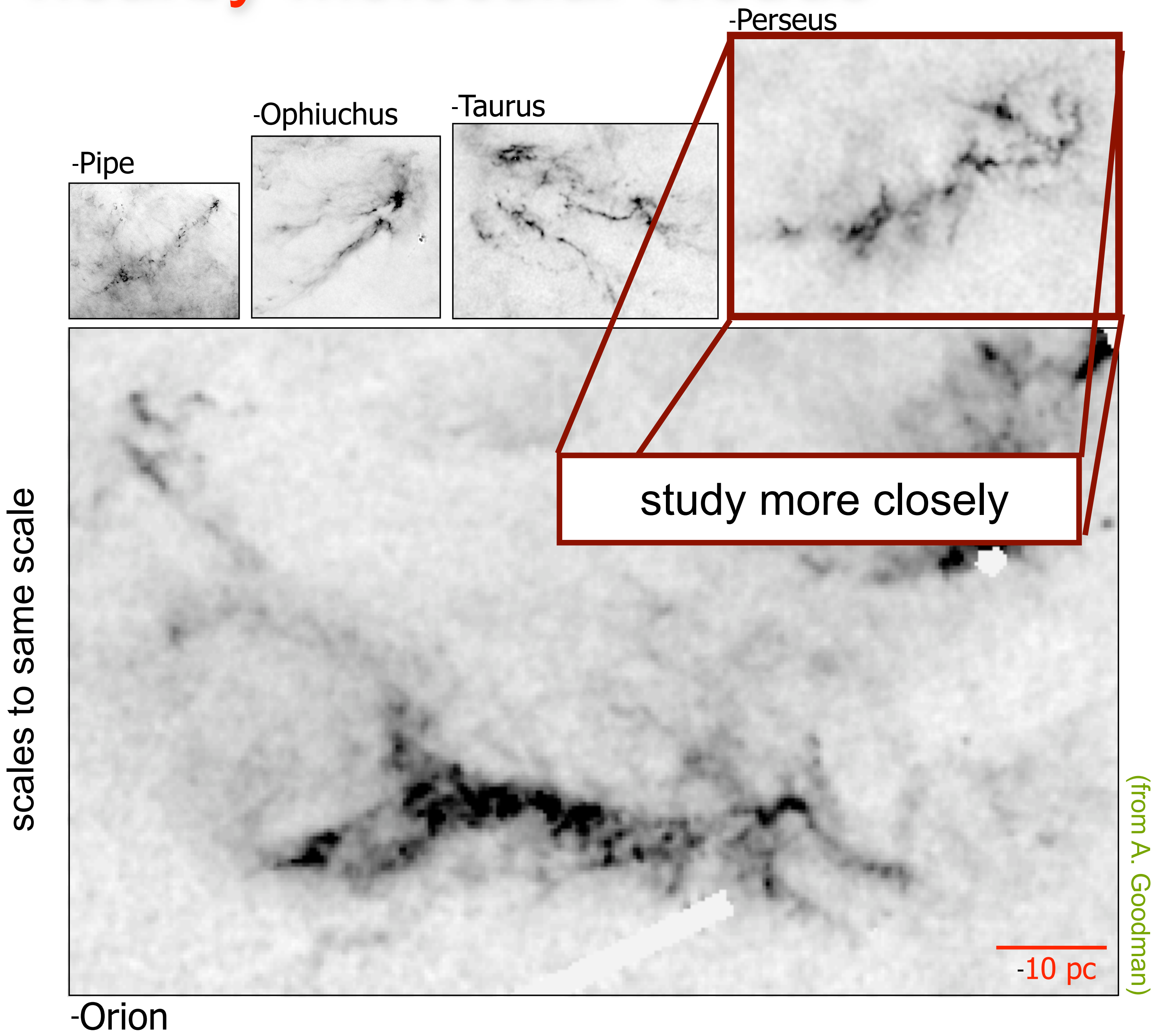
nearby molecular clouds



scales to same scale

(from A. Goodman)

nearby molecular clouds



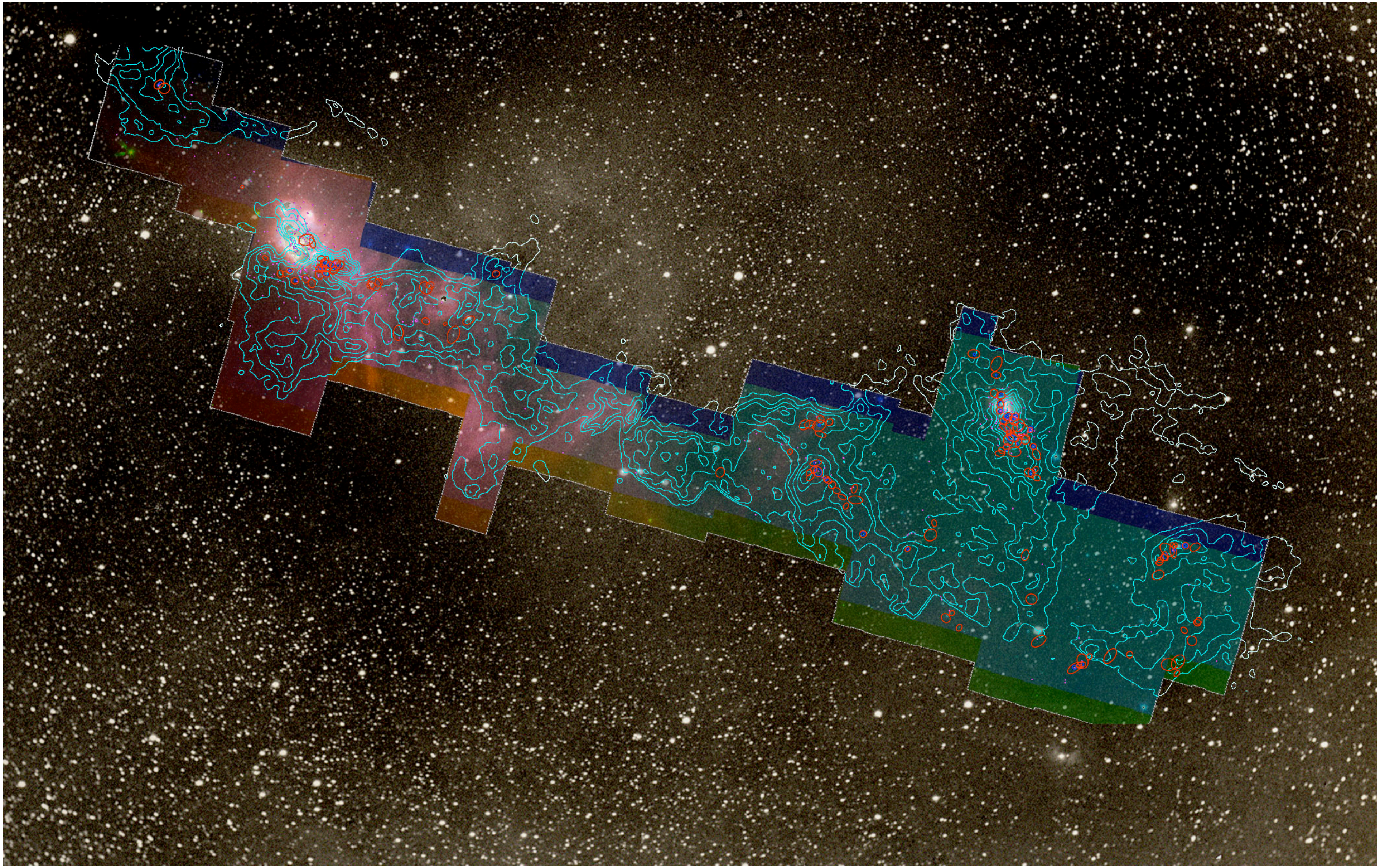
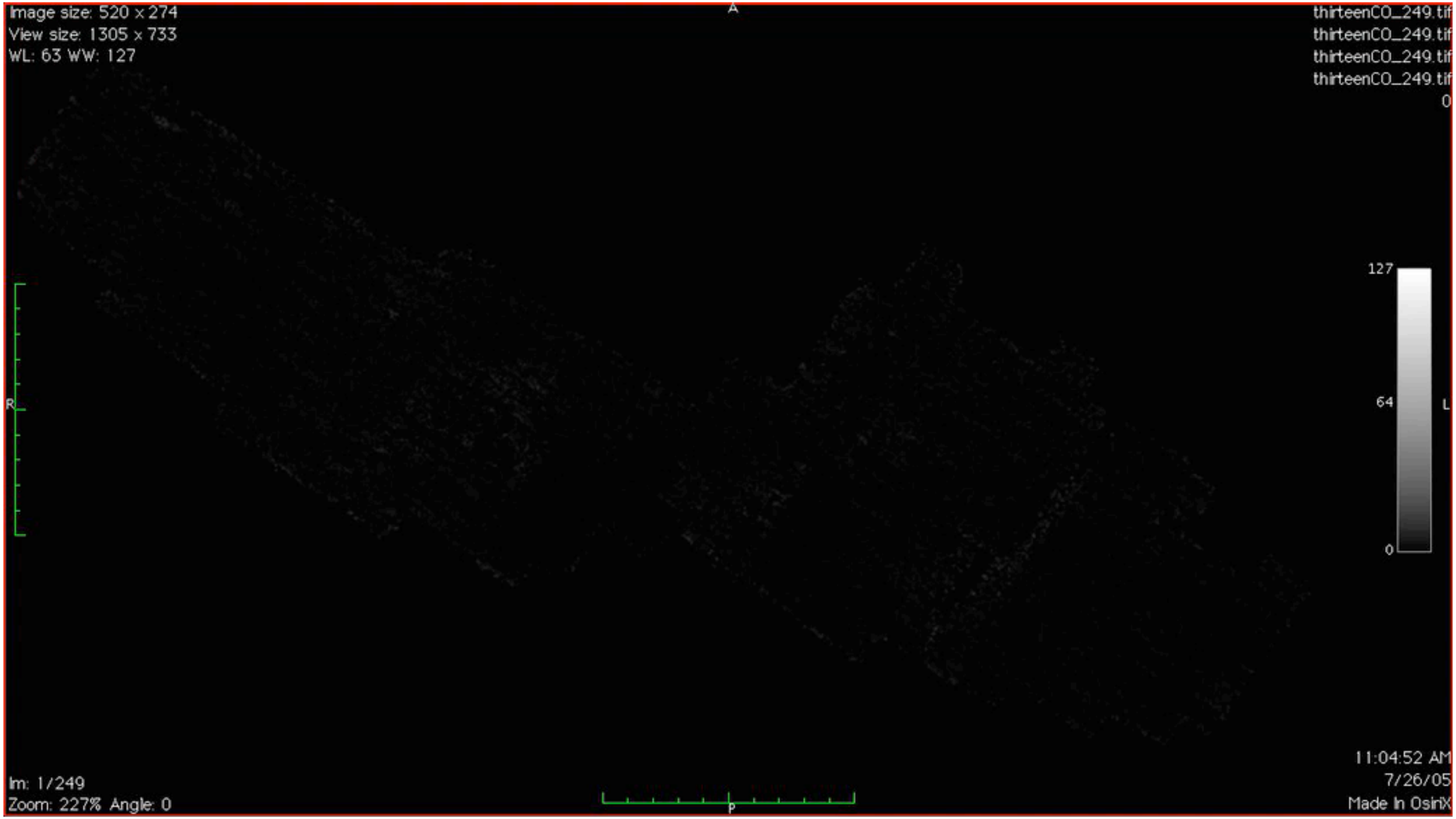


image from Alyssa Goodman: COMPLETE survey



LOS Geschwindigkeitsverteilung in Perseus

velocity cube from Alyssa Goodman: COMPLETE survey

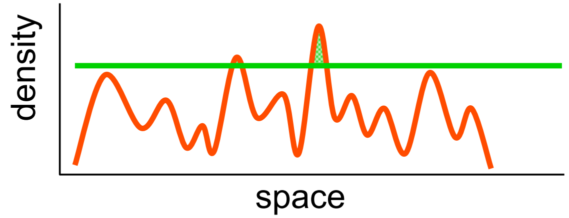
what we need to consider ...

- *correlation* between large and small scales in galaxy (stars “know” where to form and when)
- all stars form in *molecular cloud* complexes (star formation linked to molecular cloud formation)
- molecular clouds are *turbulent* (understand turbulence to understand star formation)
- stars form in *clusters* (importance of dynamical interactions during formation)
- star formation has *universal* characteristics (e.g. initial mass fuction)

basic idea

dynamical SF in a nutshell

- interstellar gas is highly *inhomogeneous*



  - *gravitational instability*
  - *thermal instability*
  - *turbulent compression* (in shocks $\delta\rho/\rho \propto \mathcal{M}^2$; in atomic gas: $\mathcal{M} \approx 1...3$)

- cold *molecular clouds* can form rapidly in high-density regions at *stagnation points* of *convergent large-scale flows*
  - chemical *phase transition*: atomic → molecular
  - process is *modulated* by large-scale *dynamics* in the galaxy

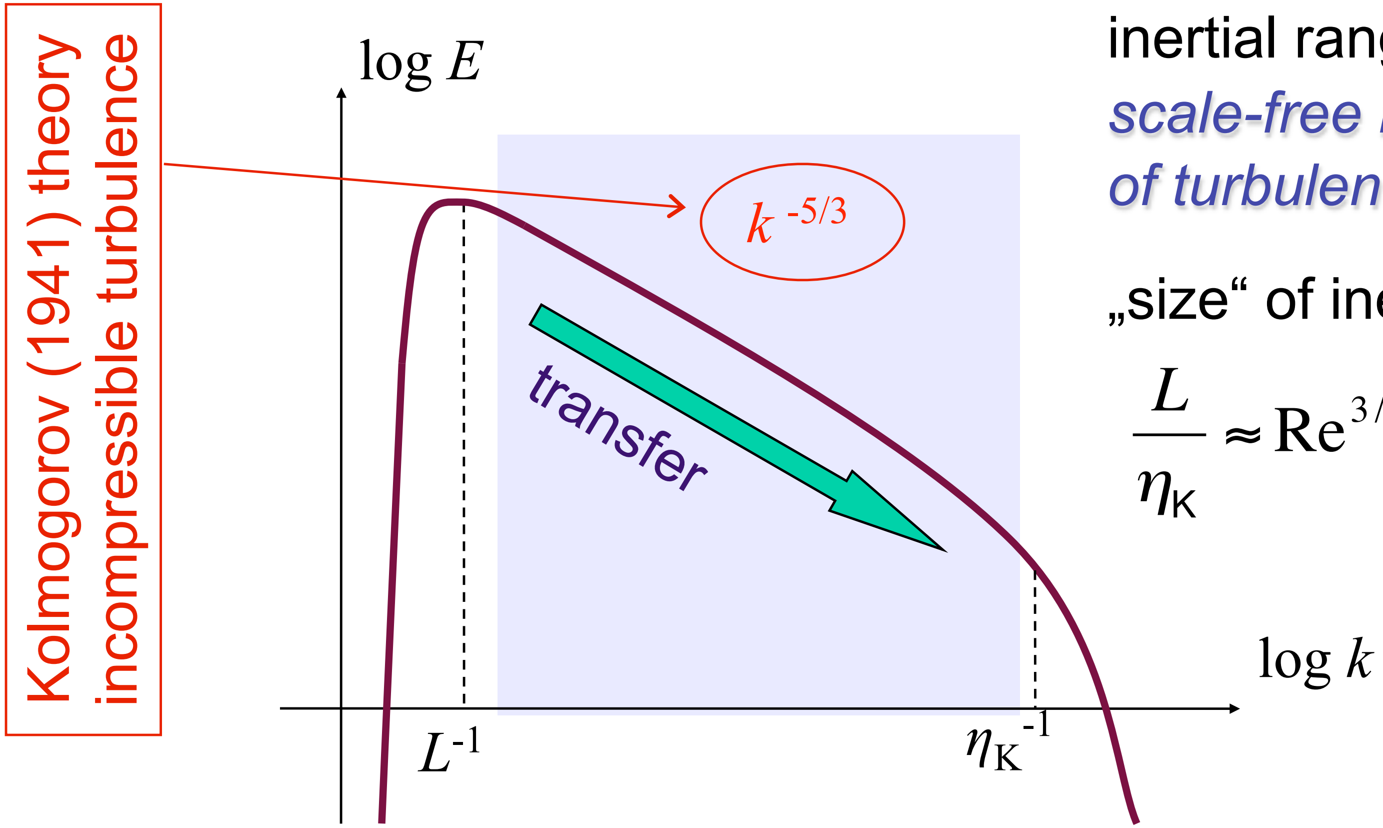
- inside *cold clouds:* turbulence is highly supersonic ($\mathcal{M} \approx 1...20$)
  → *turbulence* creates large density contrast,
  gravity selects for collapse

  → ***GRAVOTUBULENT FRAGMENTATION***

- *turbulent cascade:* local compression *within* a cloud provokes collapse
  → formation of individual *stars* and *star clusters*

(e.g. Mac Low & Klessen, 2004, Rev. Mod. Phys., 76, 125-194)

Turbulent cascade



energy
input
scale

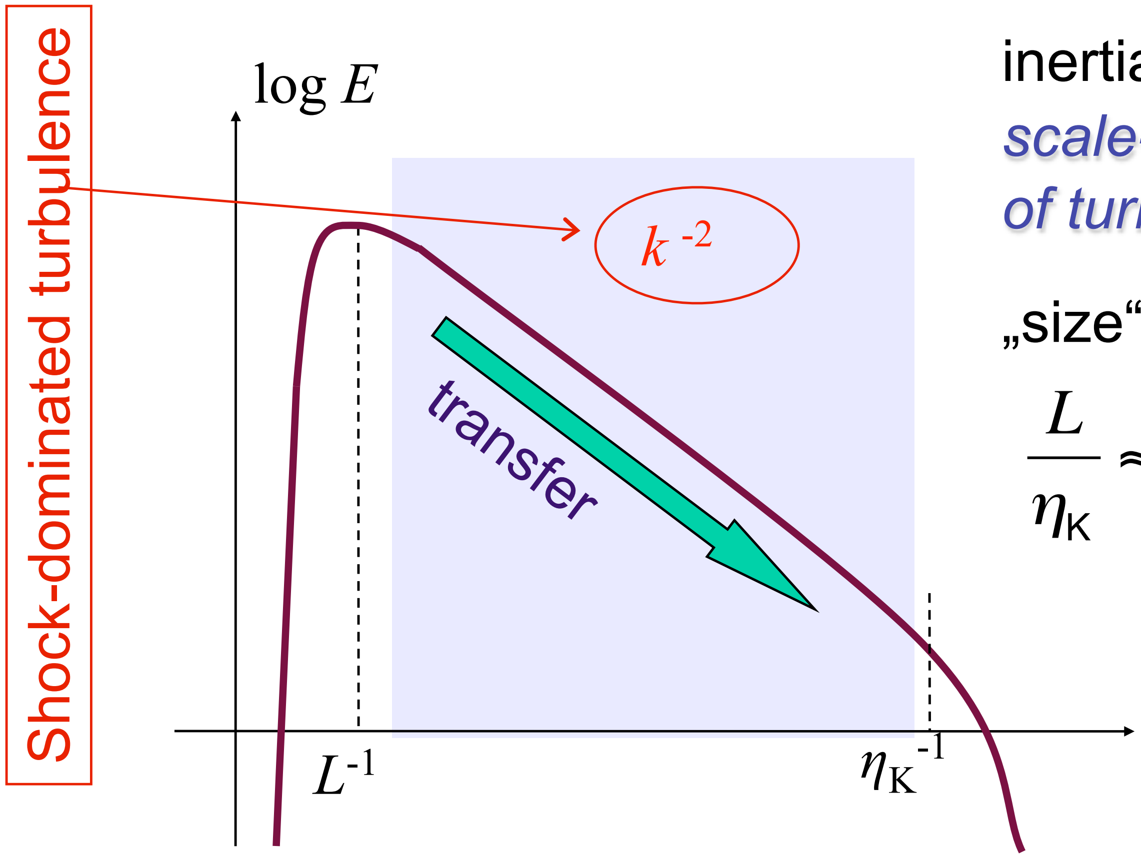
energy
dissipation
scale

inertial range:
scale-free behavior of turbulence

„size“ of inertial range:

$$\frac{L}{\eta_K} \approx \mathrm{Re}^{3/4}$$

Turbulent cascade



inertial range:

scale-free behavior of turbulence

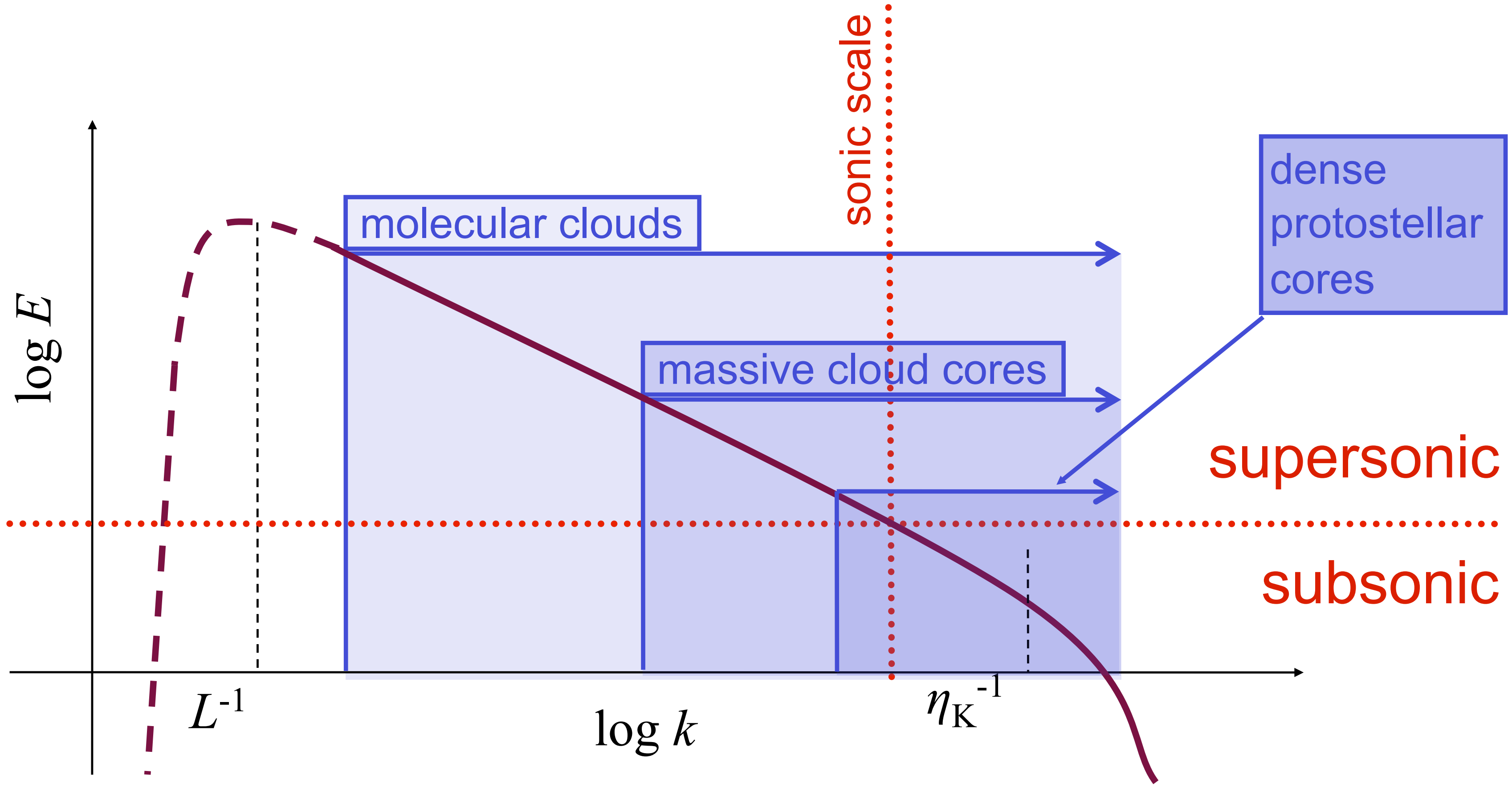
„size“ of inertial range:

$$\frac{L}{\eta_K} \approx \mathrm{Re}^{3/4}$$

energy input scale

energy dissipation scale

Turbulent cascade in ISM



energy source & scale *NOT known* (supernovae, winds, spiral density waves?)

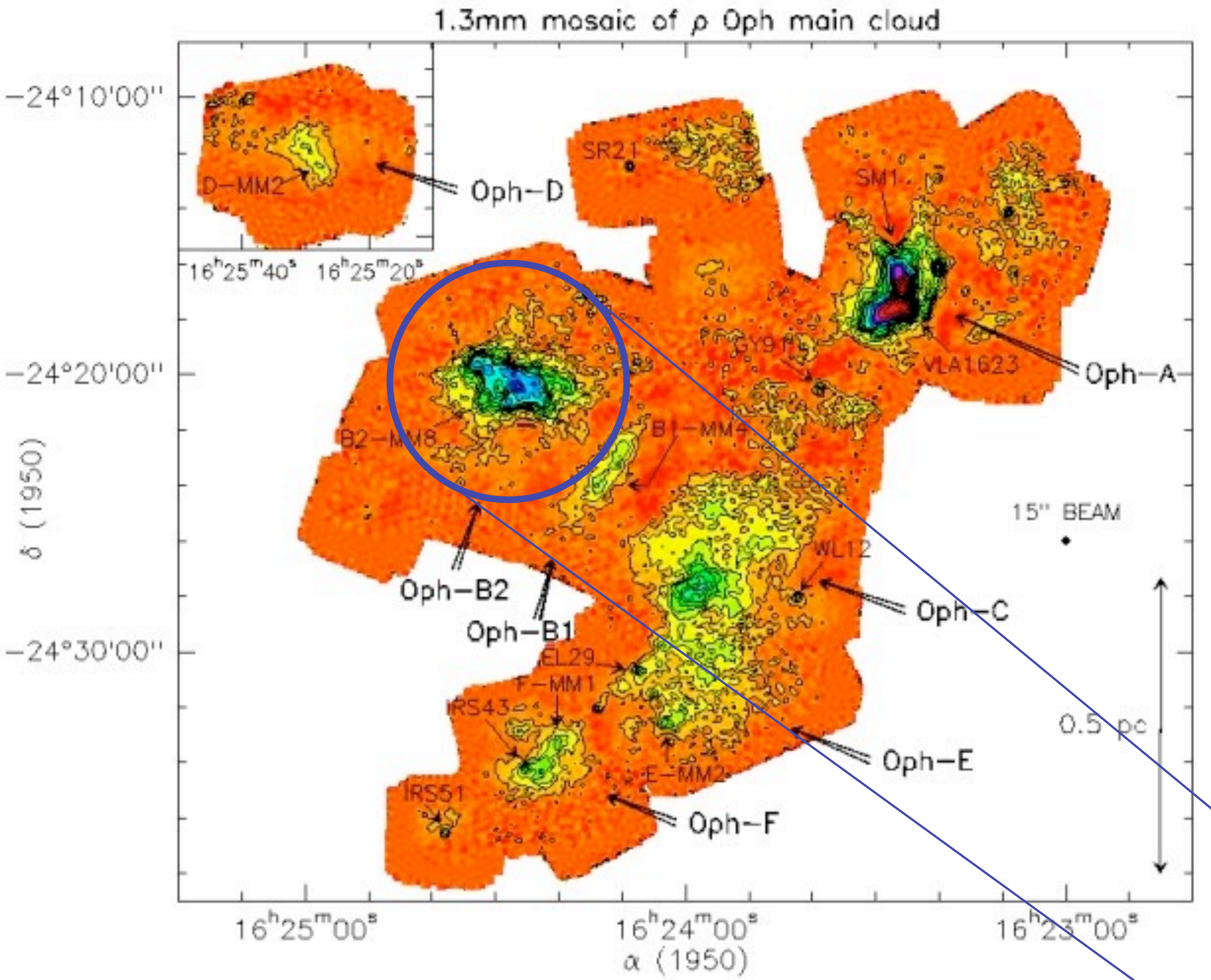
σ_{rms} << 1 km/s

$M_{rms} \leq 1$

L ≈ 0.1 pc

dissipation scale not known (ambipolar diffusion, molecular diffusion?)

Density structure of MC's



(Motte, André, & Neri 1998)

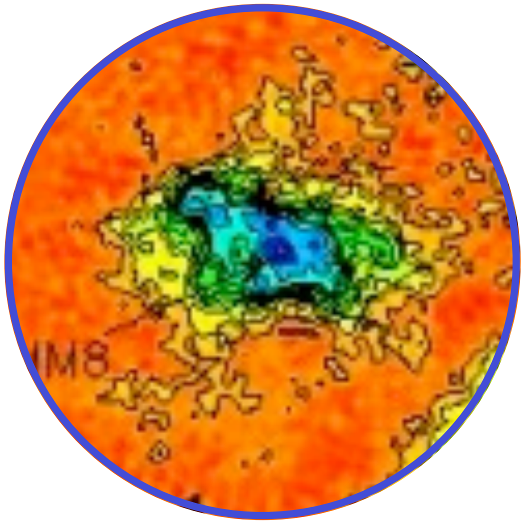
molecular clouds are highly inhomogeneous

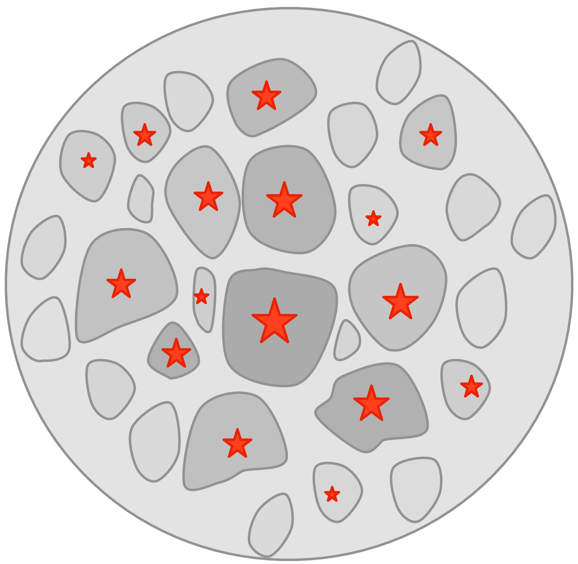
stars form in the densest and coldest parts of the cloud

ρ-Ophiuchus cloud seen in dust emission

let's focus on a cloud core like this one

Evolution of cloud cores

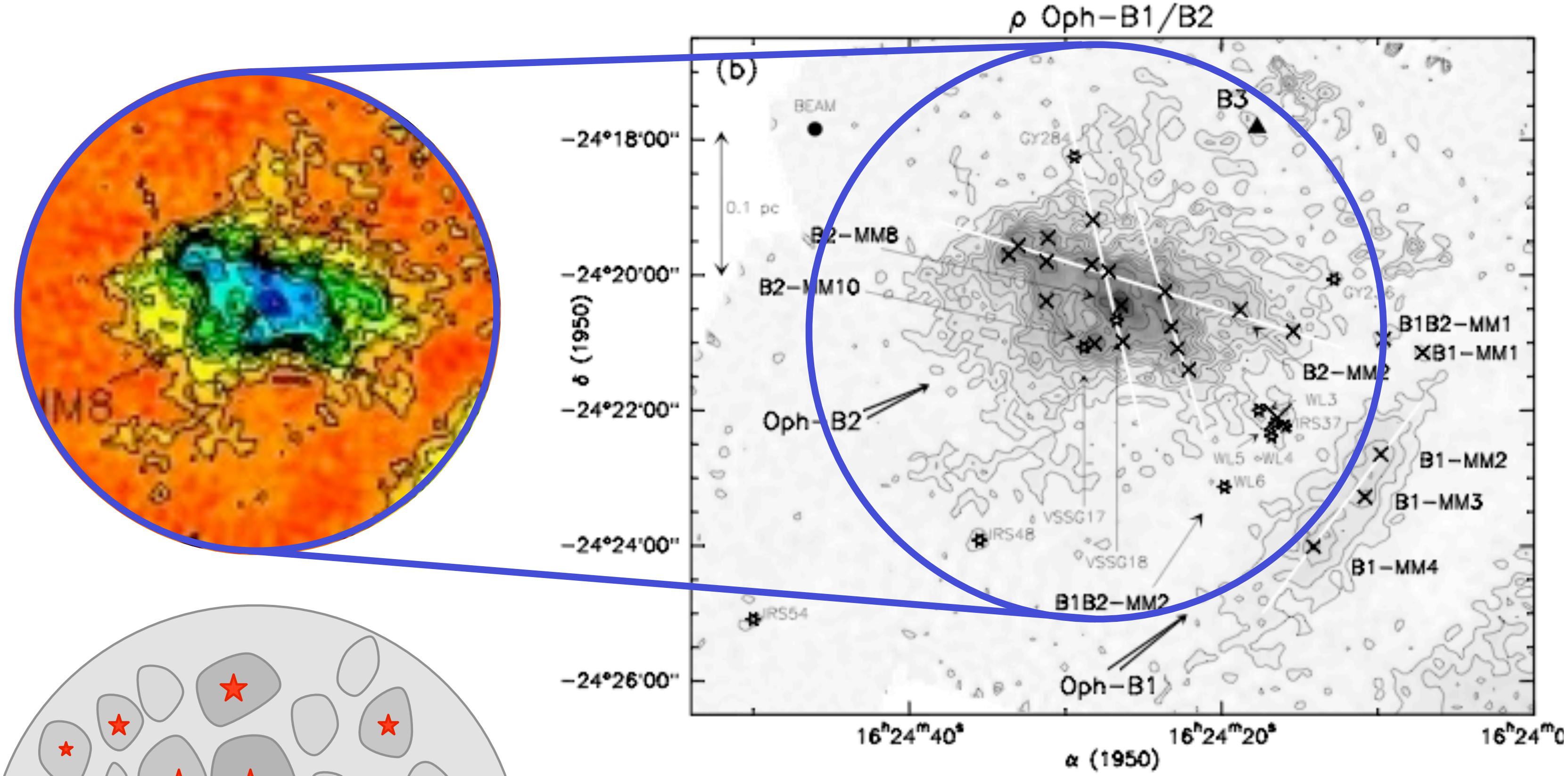


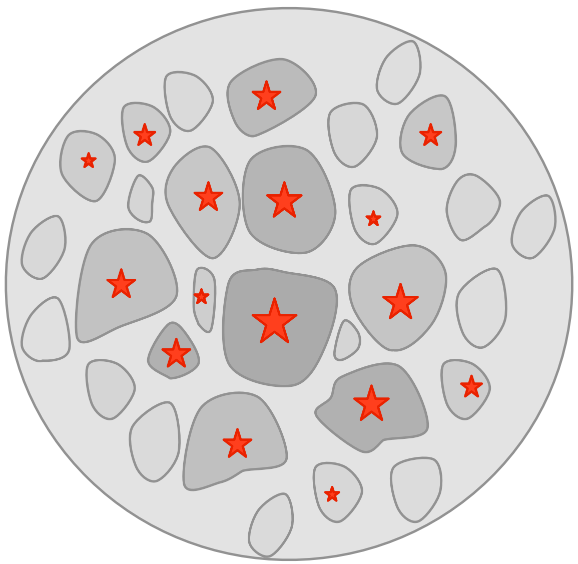


- How does this core evolve?
  Does it form one single massive star
  or cluster with mass distribution?

- Turbulent cascade „goes through“ cloud core
  --> NO *scale separation* possible
  --> NO *effective sound speed*
- Turbulence is supersonic!
  --> produces strong density contrasts:
  $\delta\rho/\rho \approx M^2$
  --> with typical $M \approx 10$ --> $\delta\rho/\rho \approx 100$!
- many of the shock-generated fluctuations are Jeans unstable and go into collapse
- --> expectation: *core breaks up and forms a cluster of stars*

Evolution of cloud cores



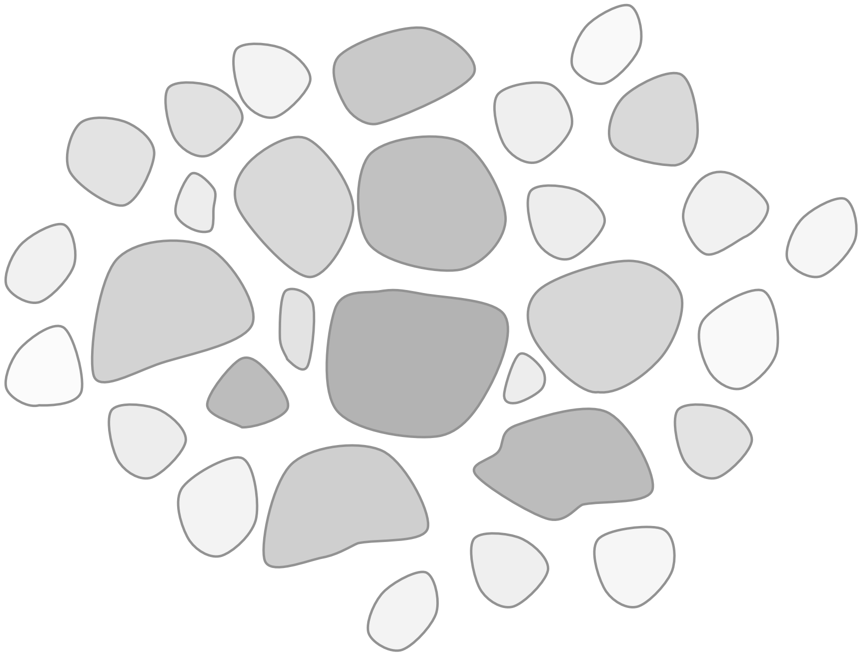


indeed ρ-Oph B1/2 contains several cores ("starless" cores are denoted by ×, cores with embedded protostars by ☆)

(Motte, André, & Neri 1998)

Formation and evolution of cores

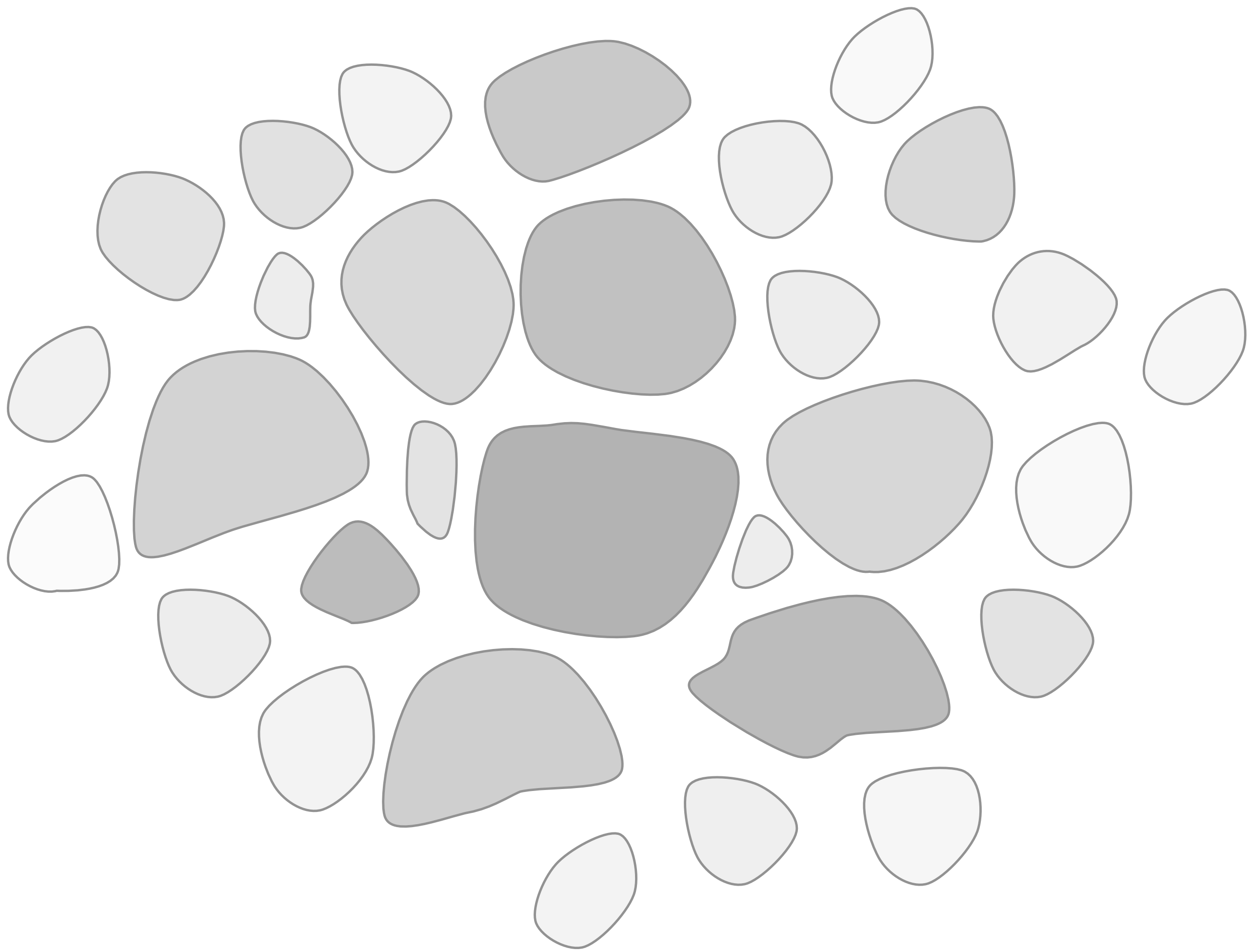
What happens to distribution of cloud cores?



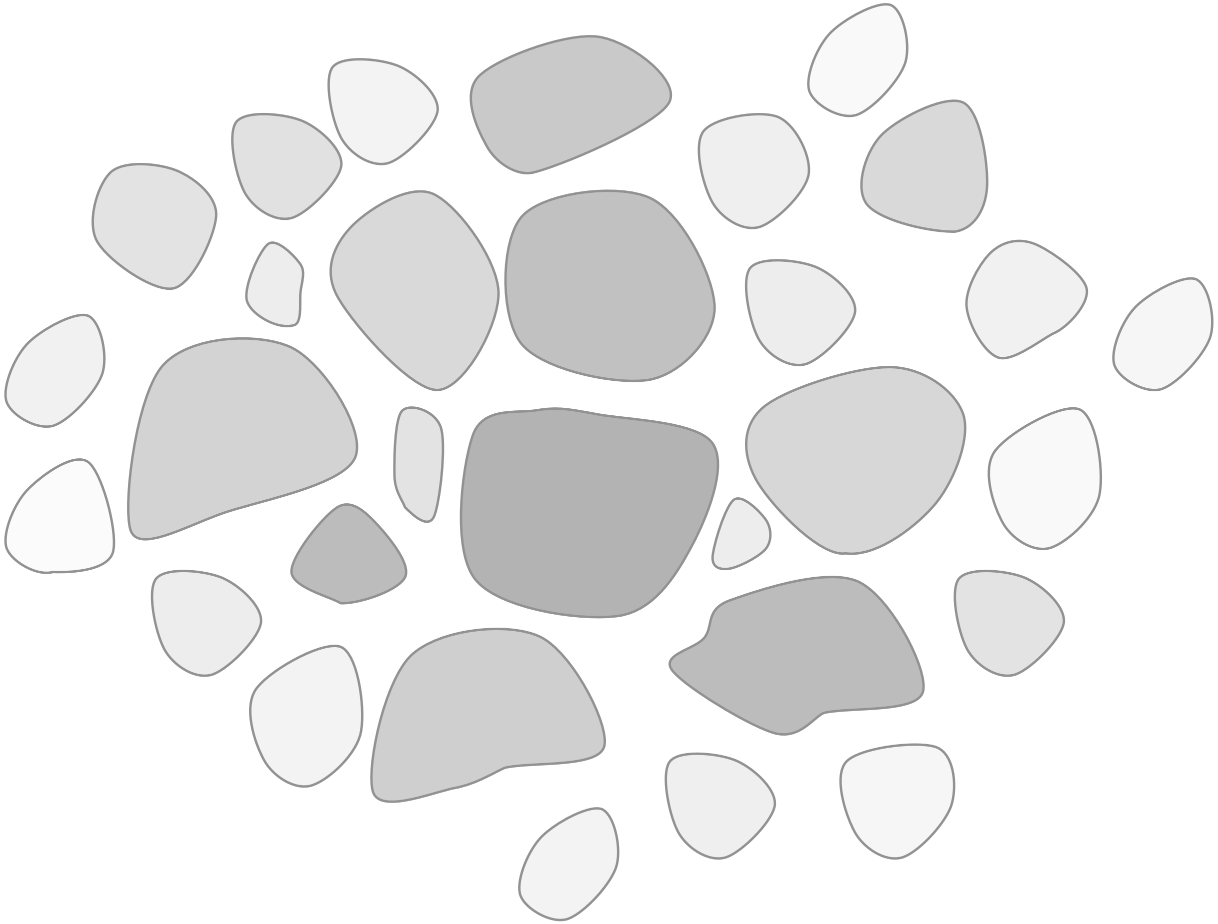
Two exteme cases:

(1) turbulence dominates energy budget:
$\alpha = E_{kin}/|E_{pot}| > 1$
--> individual cores do *not* interact
--> *collapse* of *individual* cores dominates *stellar mass growth*
--> *loose cluster of low-mass stars*

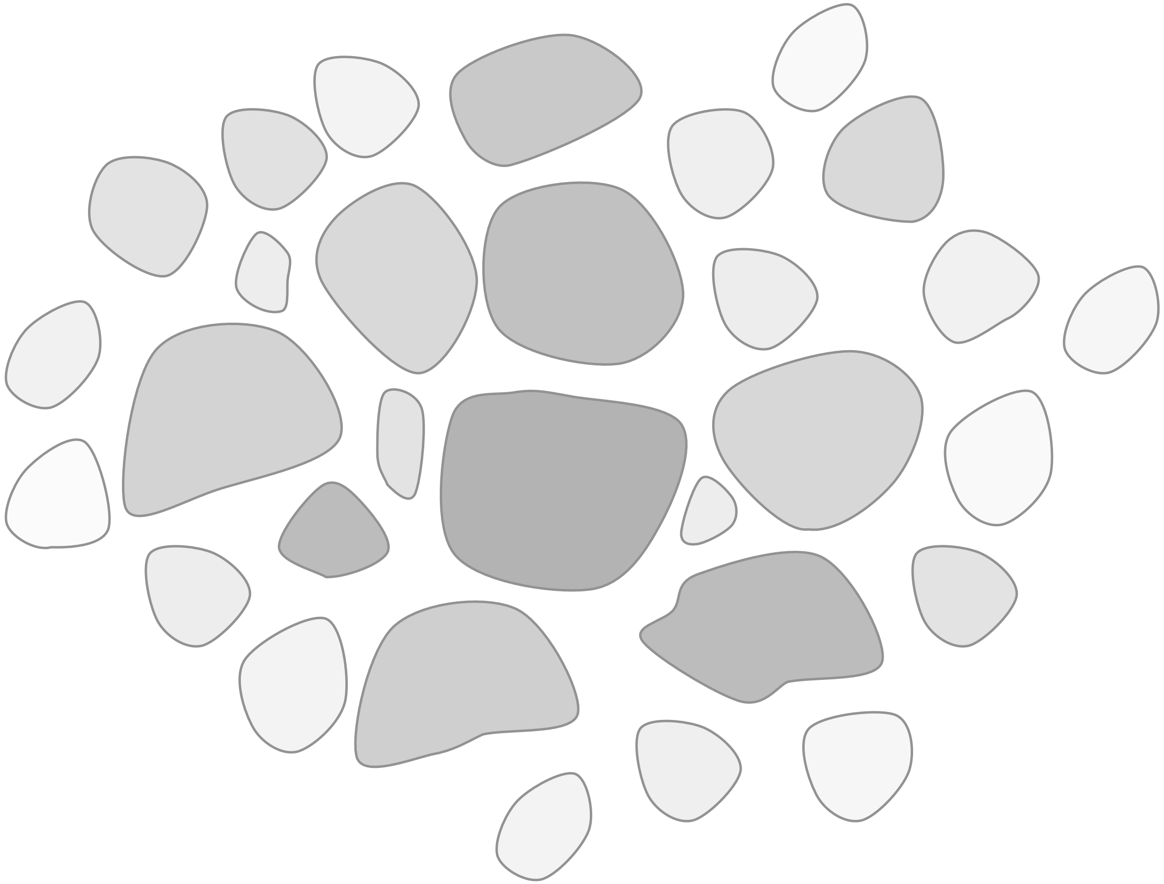
(2) turbulence decays, i.e. gravity dominates: $\alpha = E_{kin}/|E_{pot}| < 1$
--> *global contraction*
--> core do *interact* while collapsing
--> *competition* influences *mass growth*
--> *dense cluster with high-mass stars*



turbulence creates a hierarchy of clumps



as turbulence decays locally, contraction sets in



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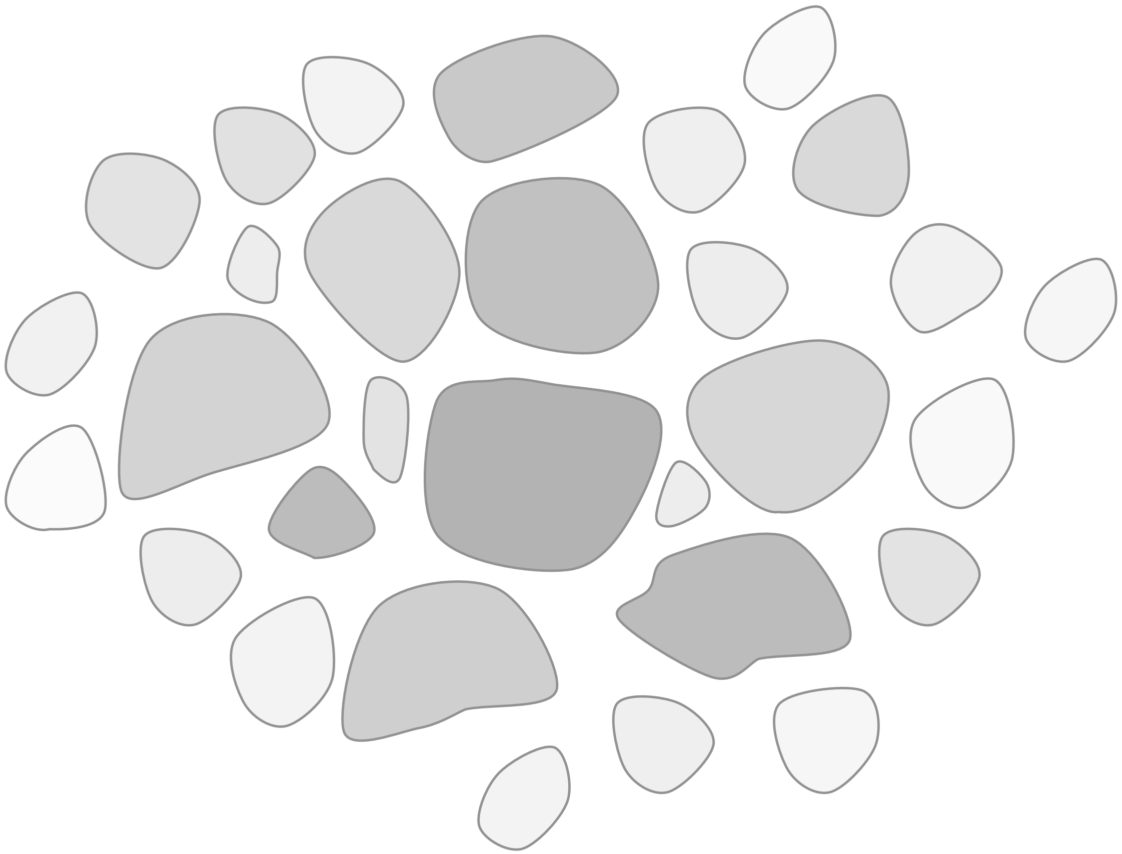
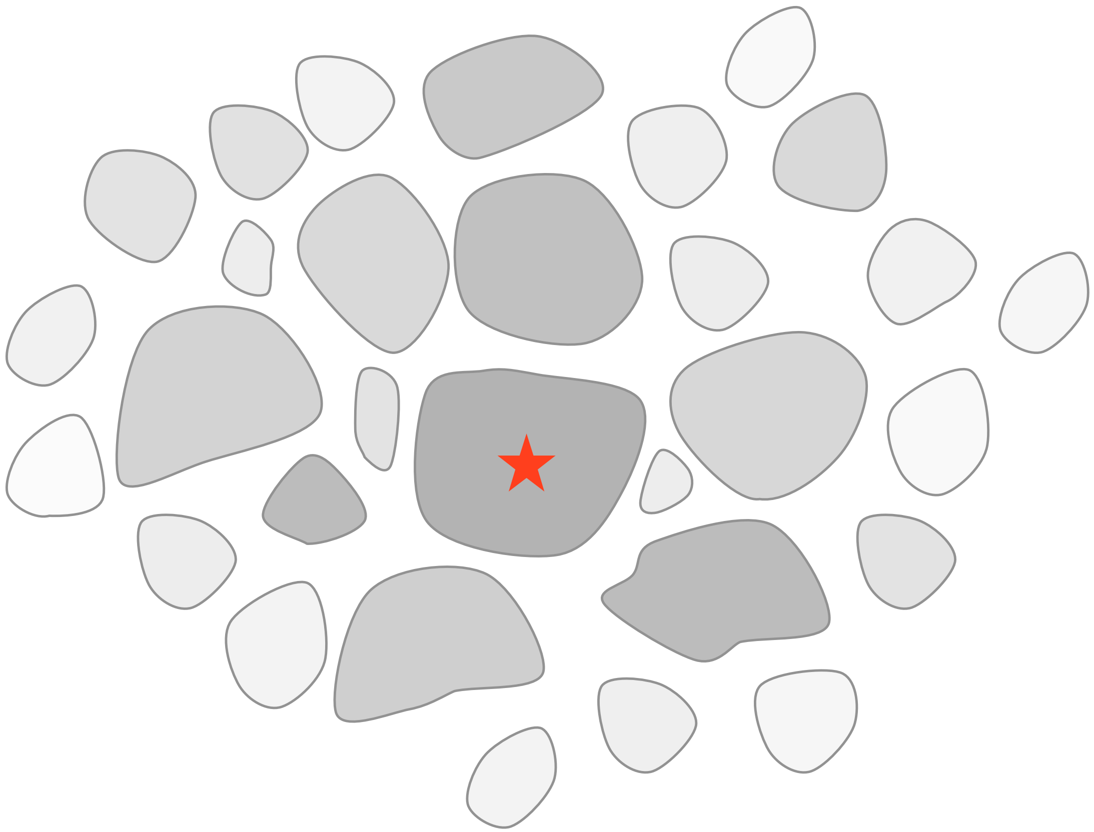


while region contracts, individual clumps collapse to form stars



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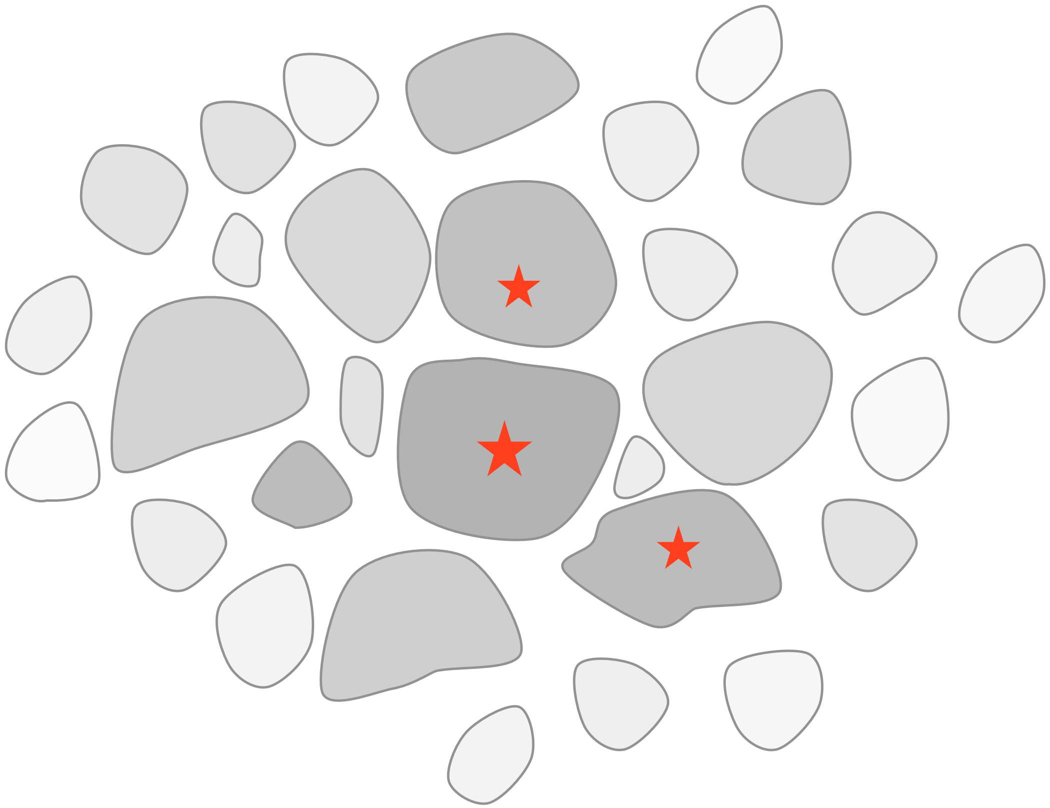
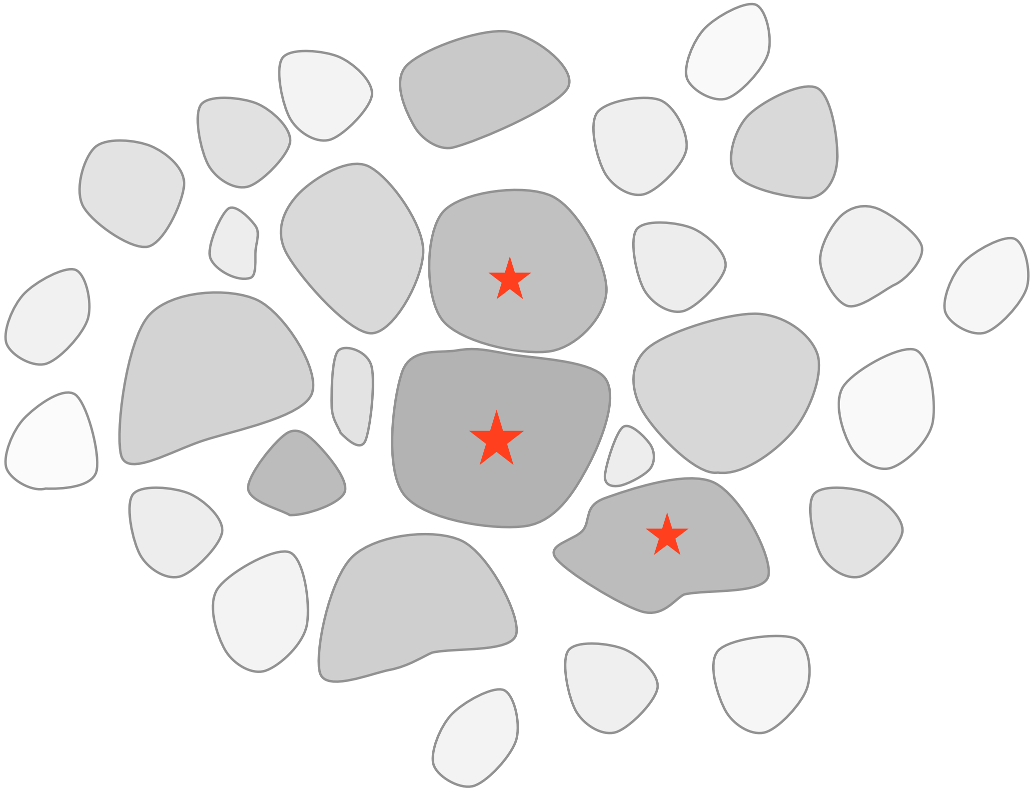


individual clumps collapse to form stars



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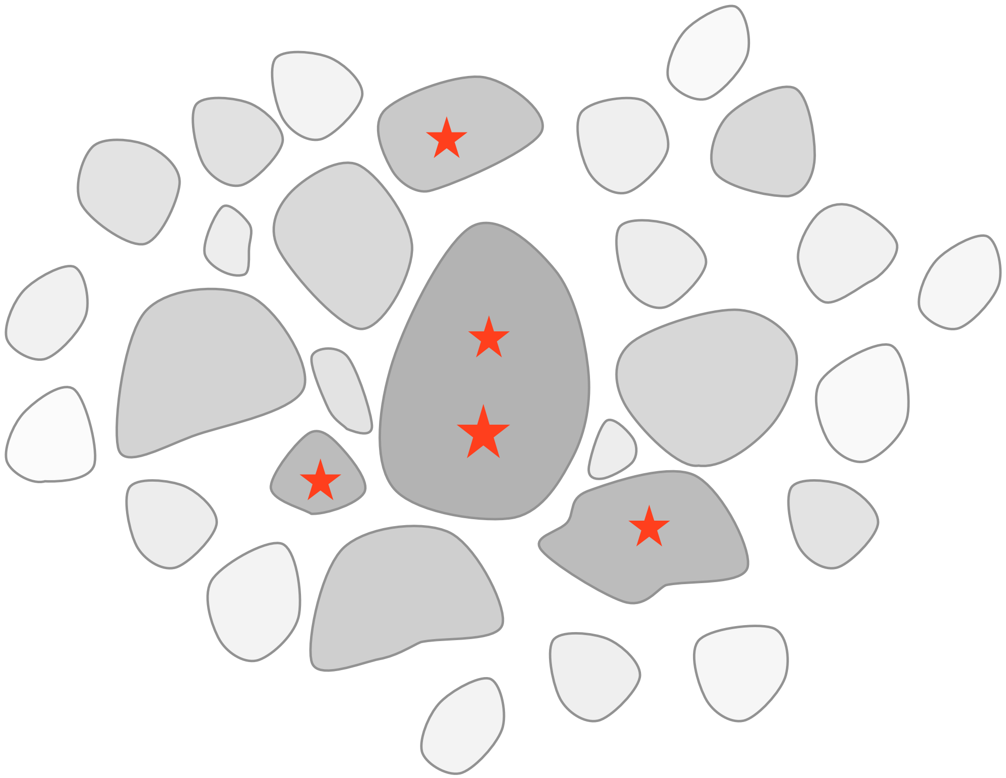
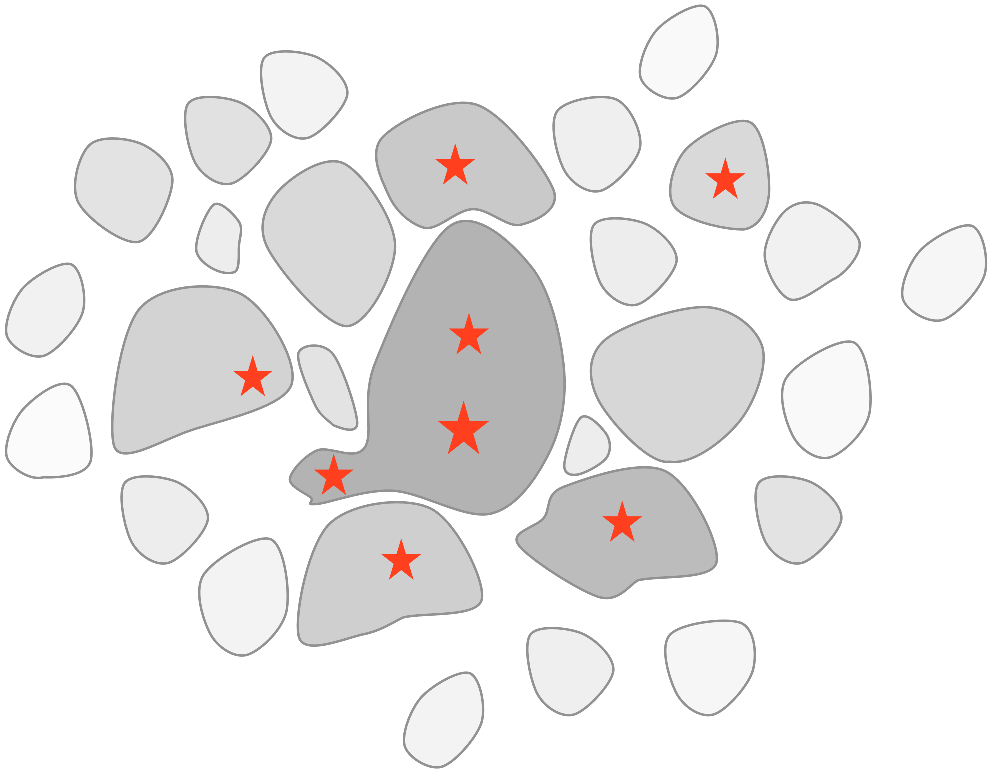


$\alpha = E_{kin}/|E_{pot}| < 1$

in *dense clusters*, clumps may merge while collapsing
--> then contain multiple protostars



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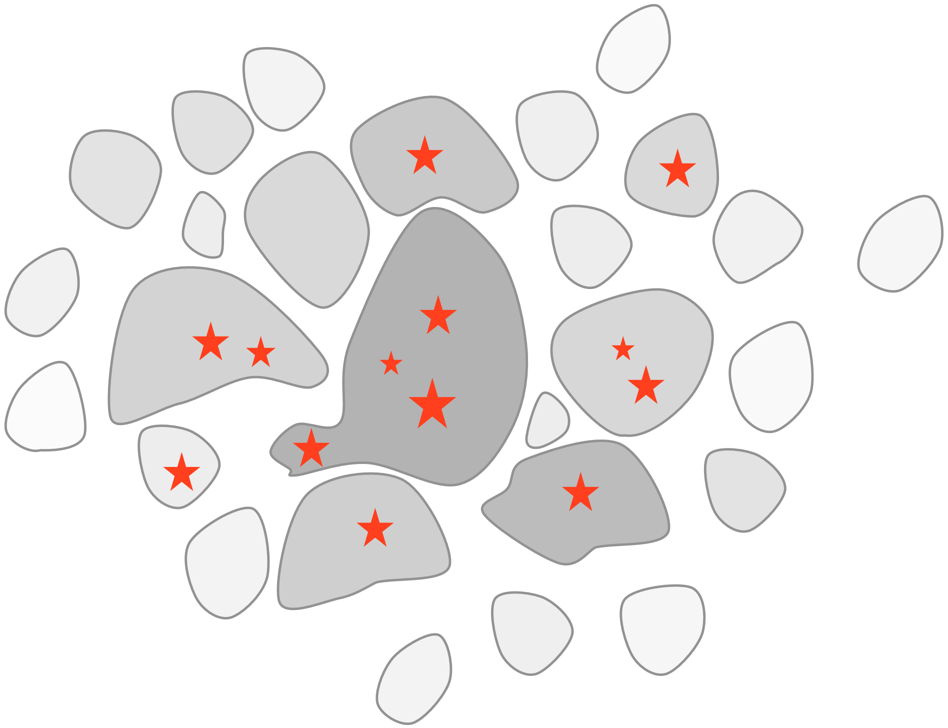

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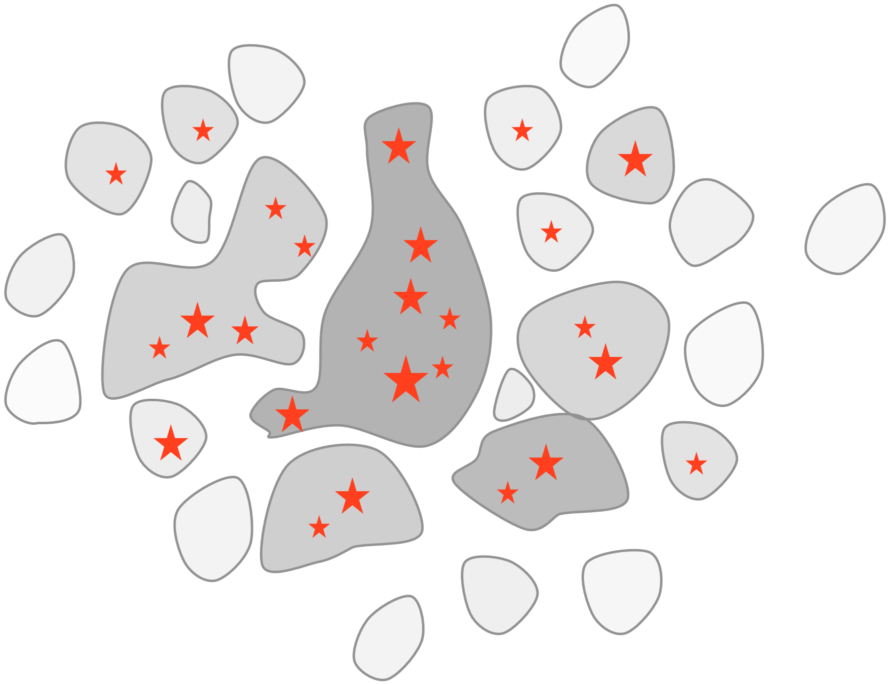
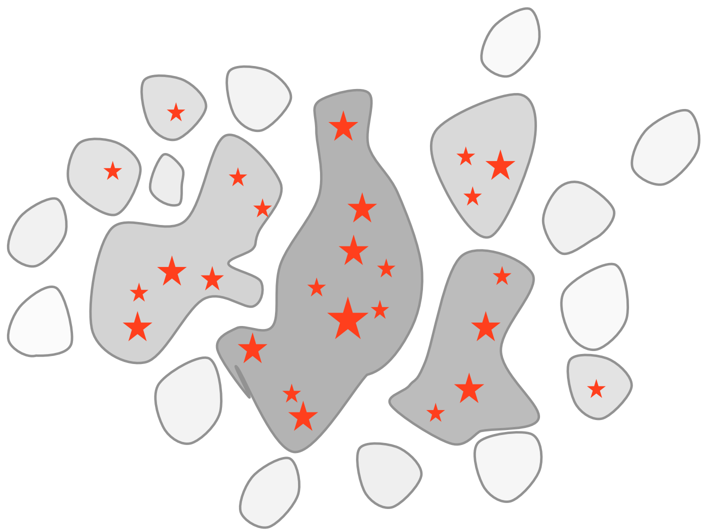


in *dense clusters*, competitive mass growth becomes important



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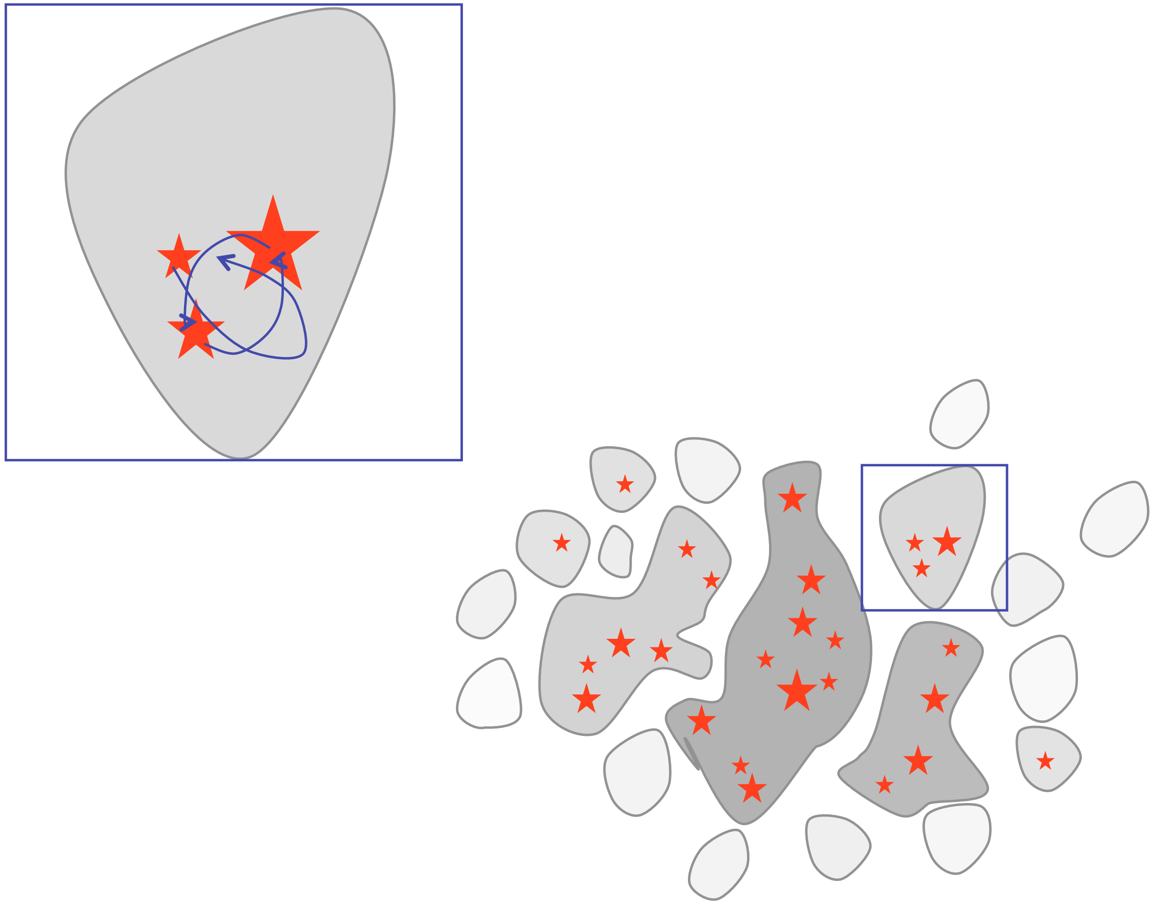
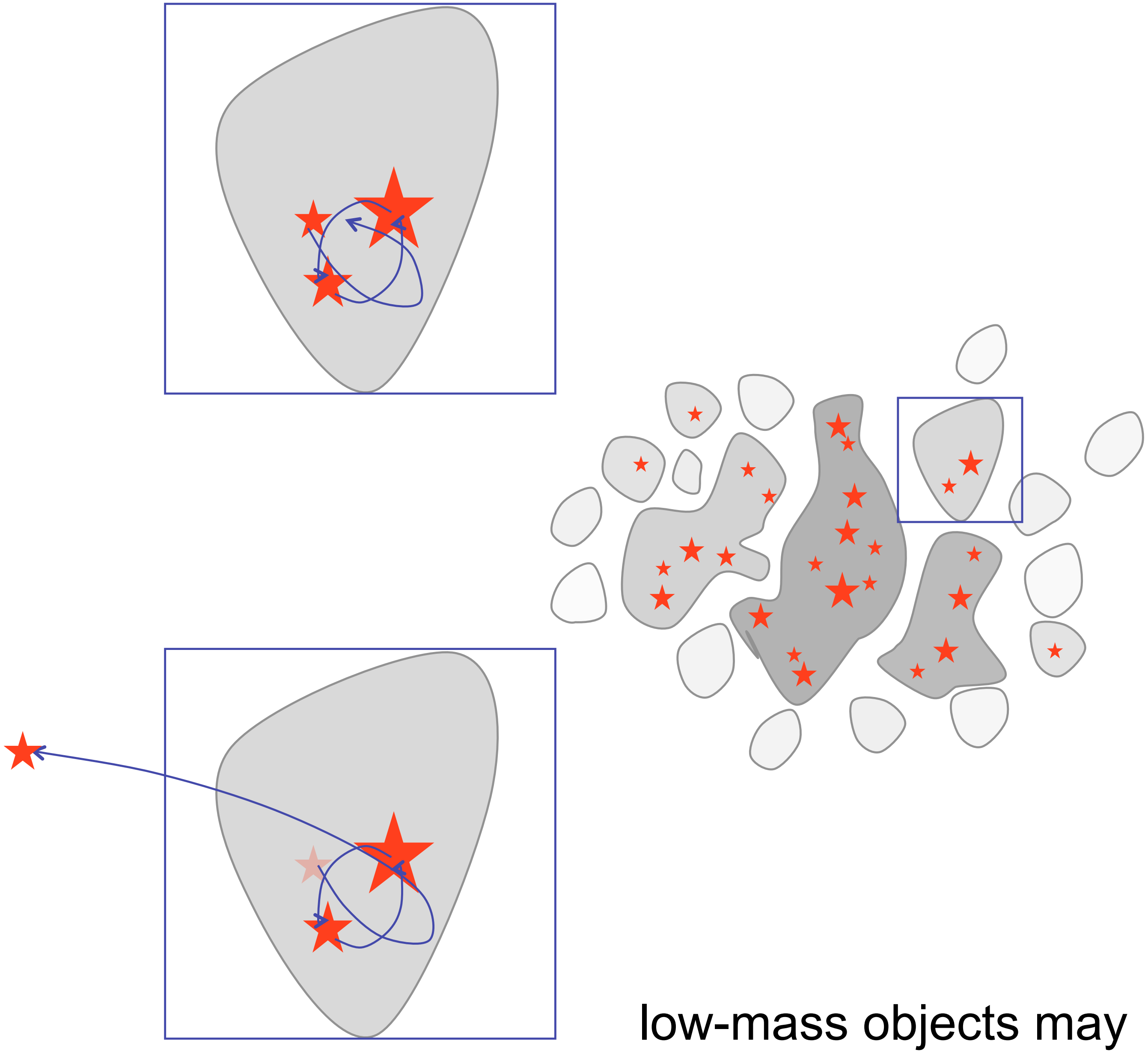
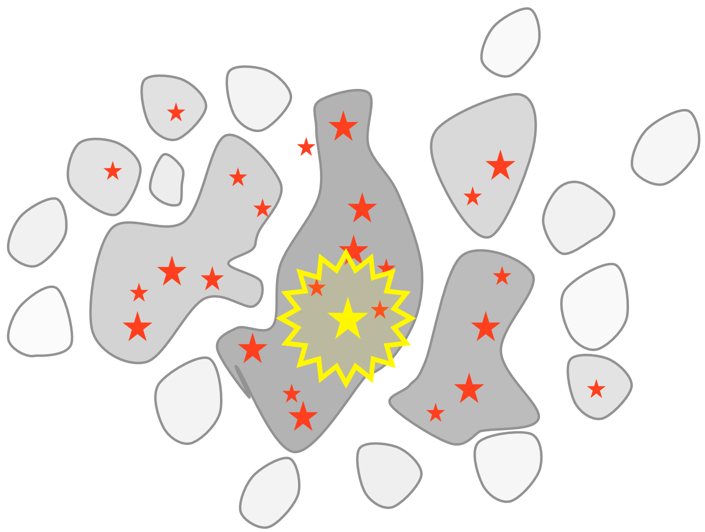


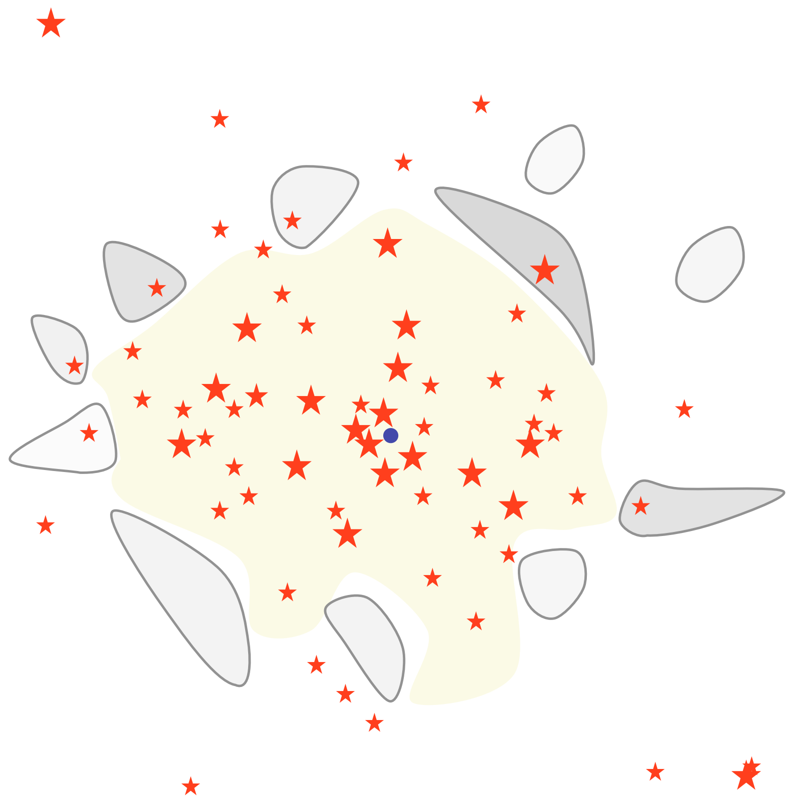
in *dense clusters*, *N*-body effects influence mass growth



low-mass objects may
become ejected --> accretion stops



feedback terminates star formation



result: *star cluster*, possibly with HII region

NGC 602 in the LMC: Hubble Heritage Image

result: *star cluster* with HII region

applications

two examples

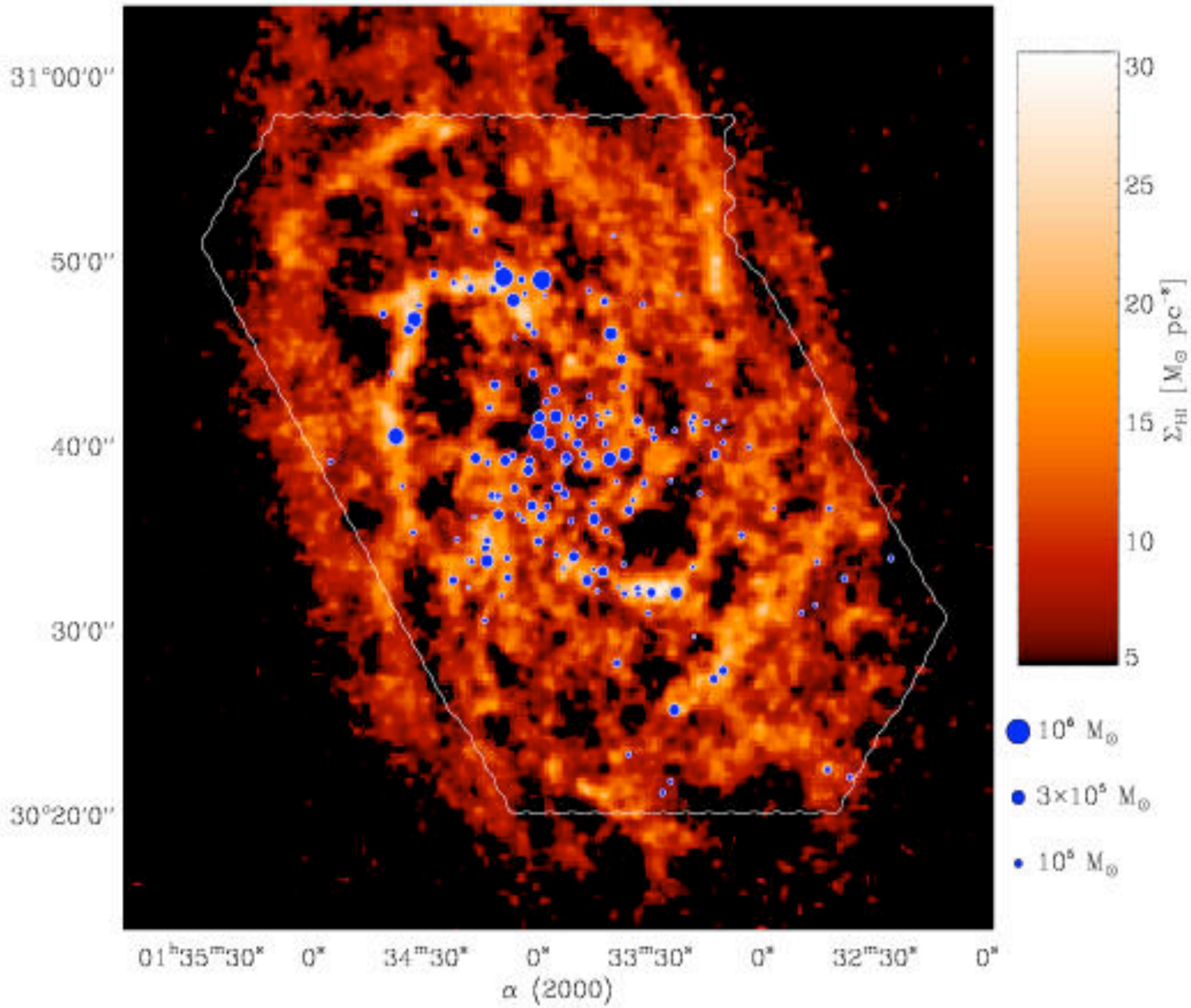
- formation of molecular clouds in the disk of the Milky Way
  - timescales
  - dynamic properties
  - x-factor
- formation of star clusters inside these clouds
  - IMF

molecular cloud formation

molecular cloud formation

- star formation on galactic scales
  → missing link so far:
  formation of molecular clouds
- questions
  - *where* and *when* do molecular clouds form?
  - *what* are their properties?
  - *how* does that correlation to star formation?
  - global correlations? → *Schmidt law*

molecular cloud formation



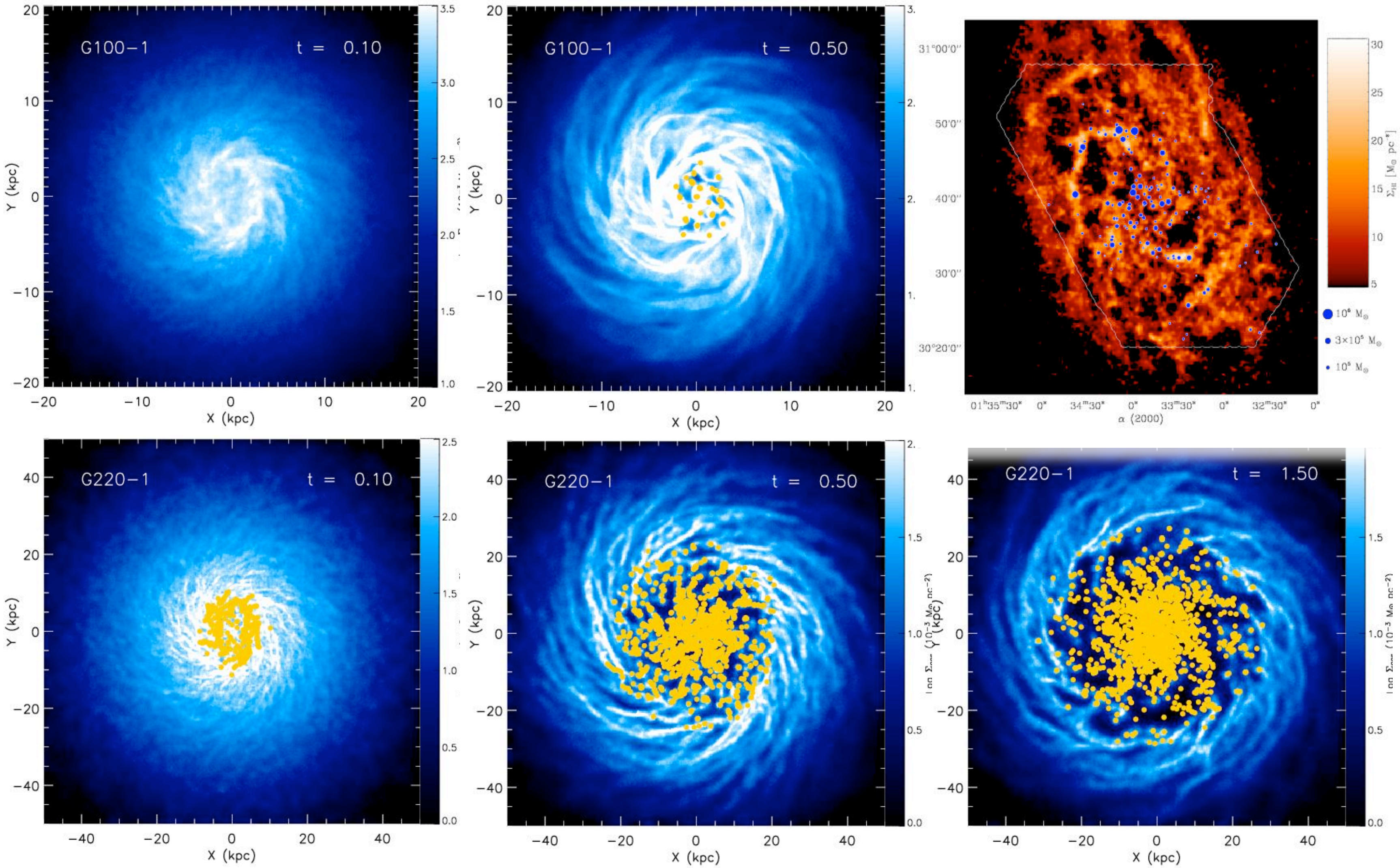
Thesis:

Molecular clouds form at *stagnation points* of large-scale convergent flows, mostly triggered by global (or external) perturbations.

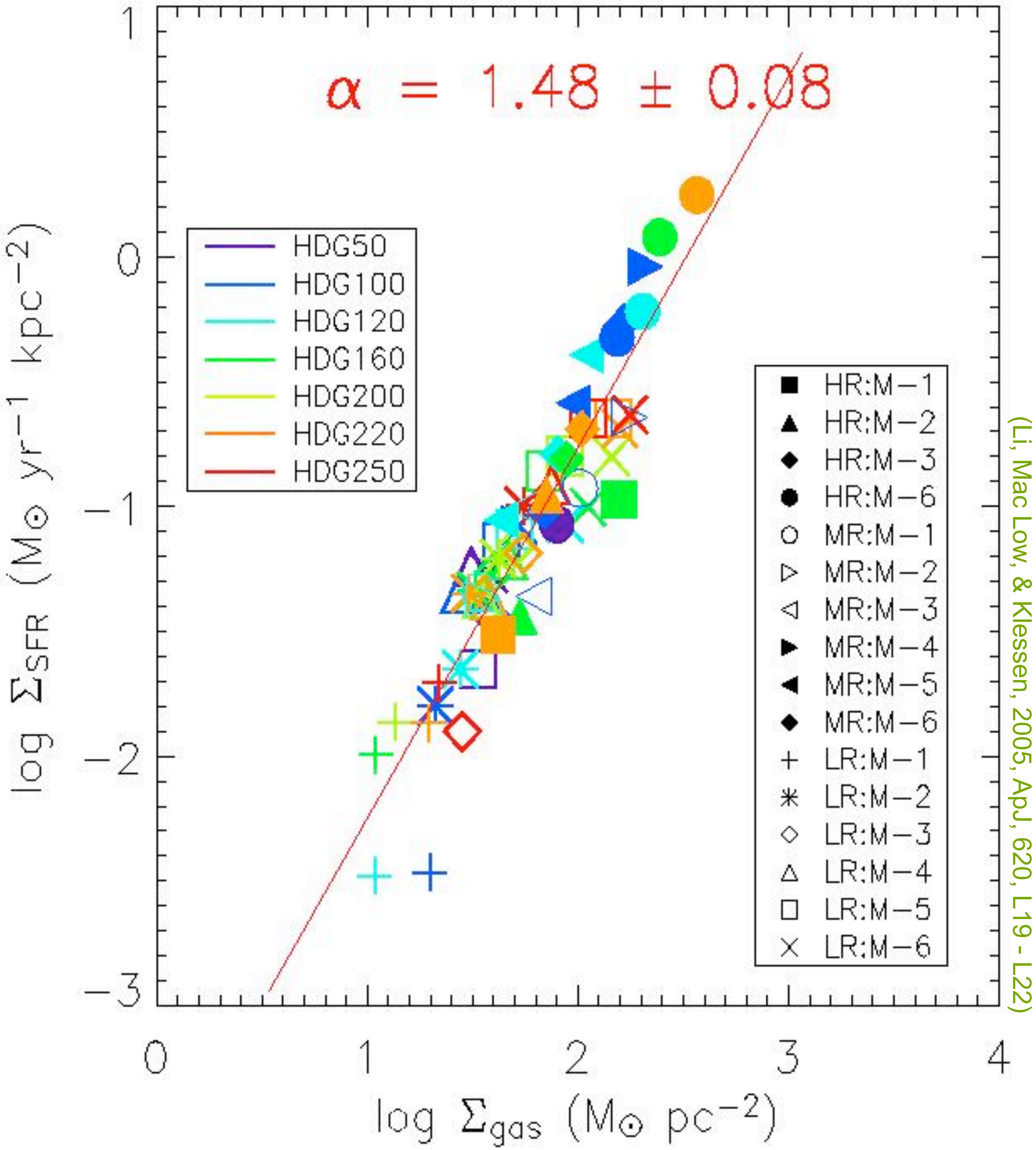
(Deul & van der Hulst 1987, Blitz et al. 2004)

modeling galactic SF

SPH calculations of self-gravitating disks of stars and (isothermal) gas in dark-matter potential, sink particles measure local collapse --> star formation



(Li, Mac Low, & Klessen, 2005, ApJ, 620,L19 - L22)

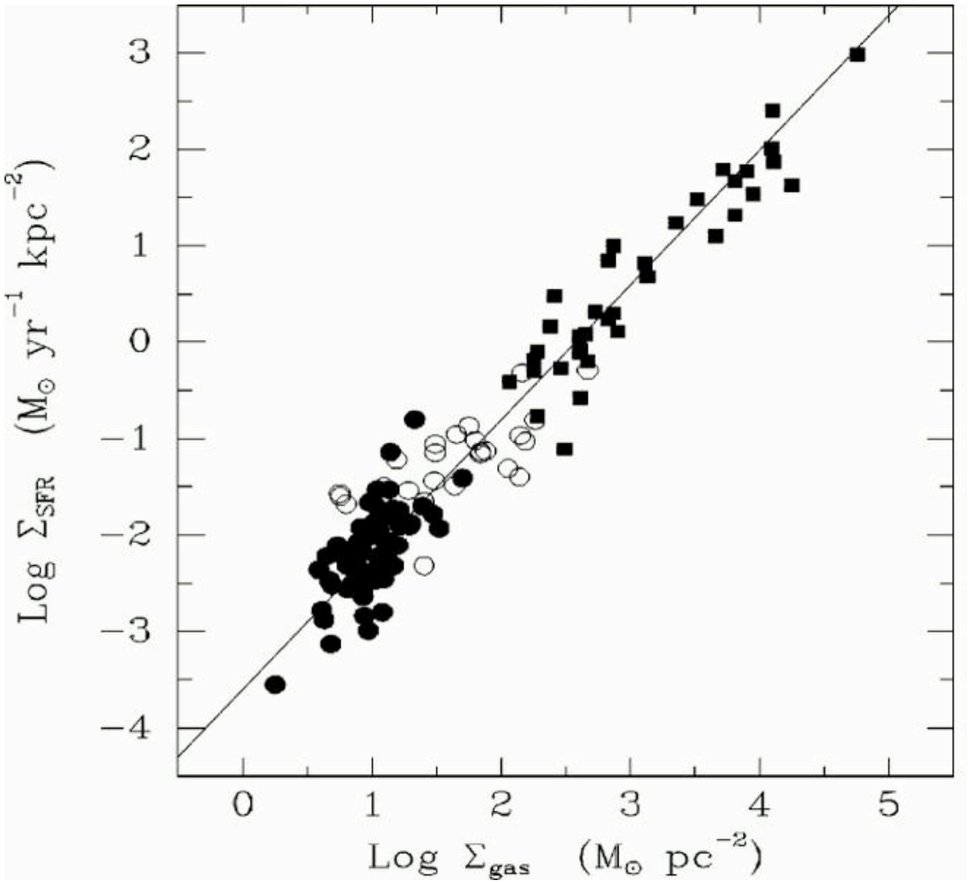


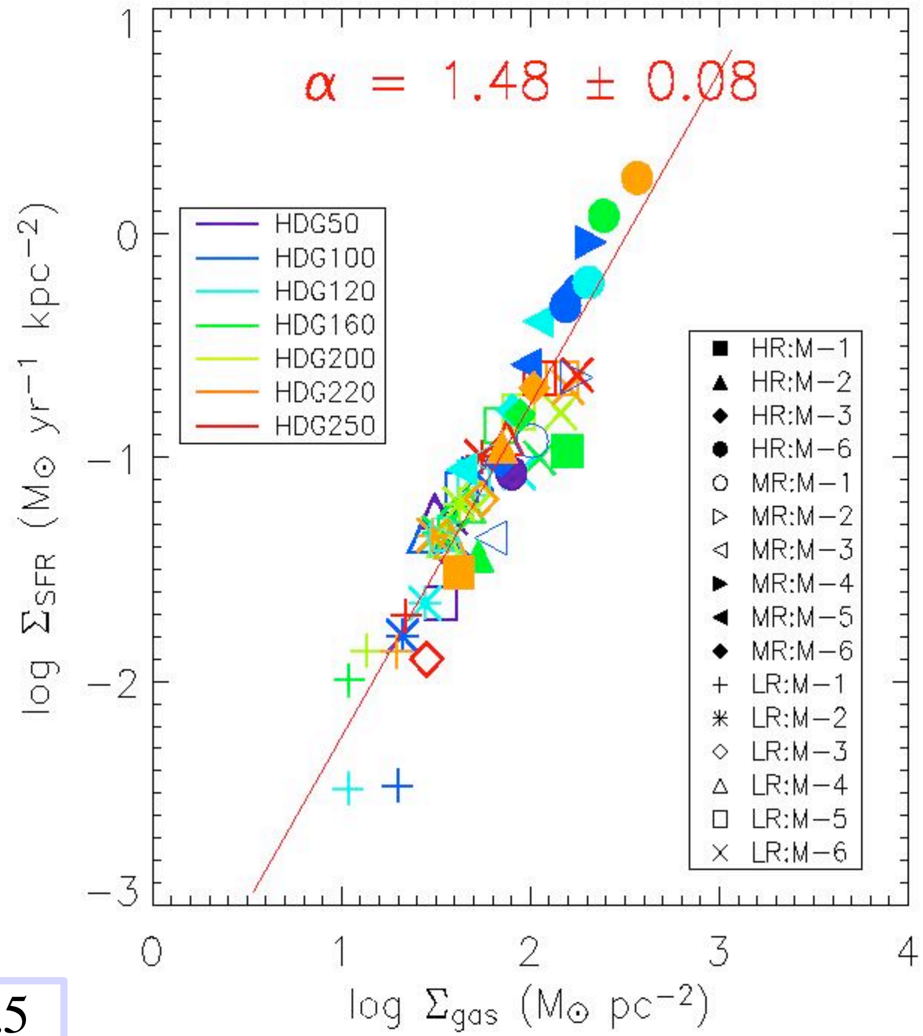
We find correlation between *star formation rate* and *gas surface density:*

$$\Sigma_{SFR} \propto \Sigma_{gas}^{1.5}$$

global Schmidt law

observed Schmidt law

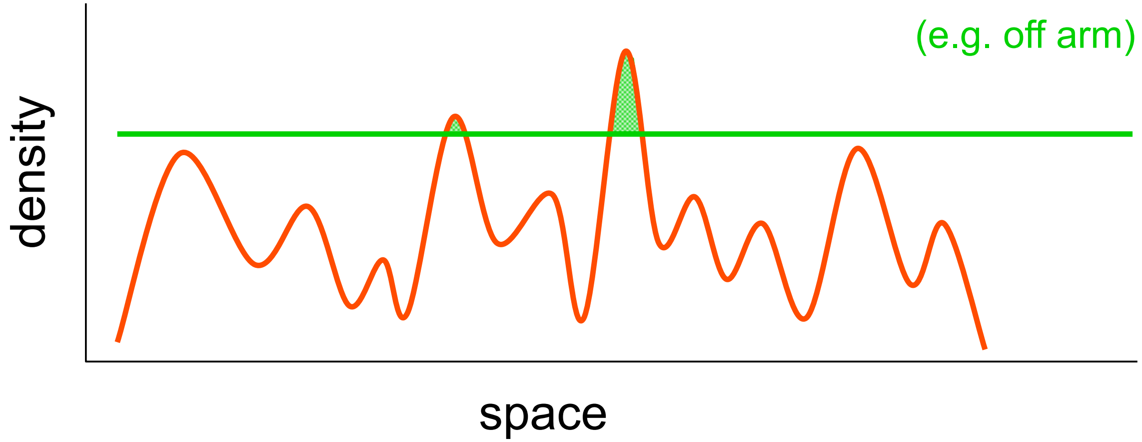


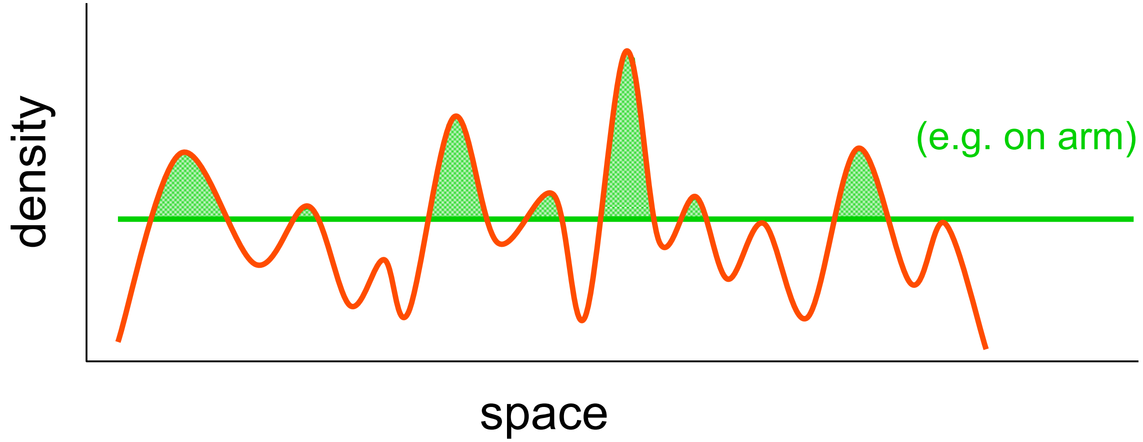


in both cases: $\Sigma_{SFR} \propto \Sigma_{gas}^{1.5}$

(from Kennicutt 1998)

correlation with large-scale perturbations





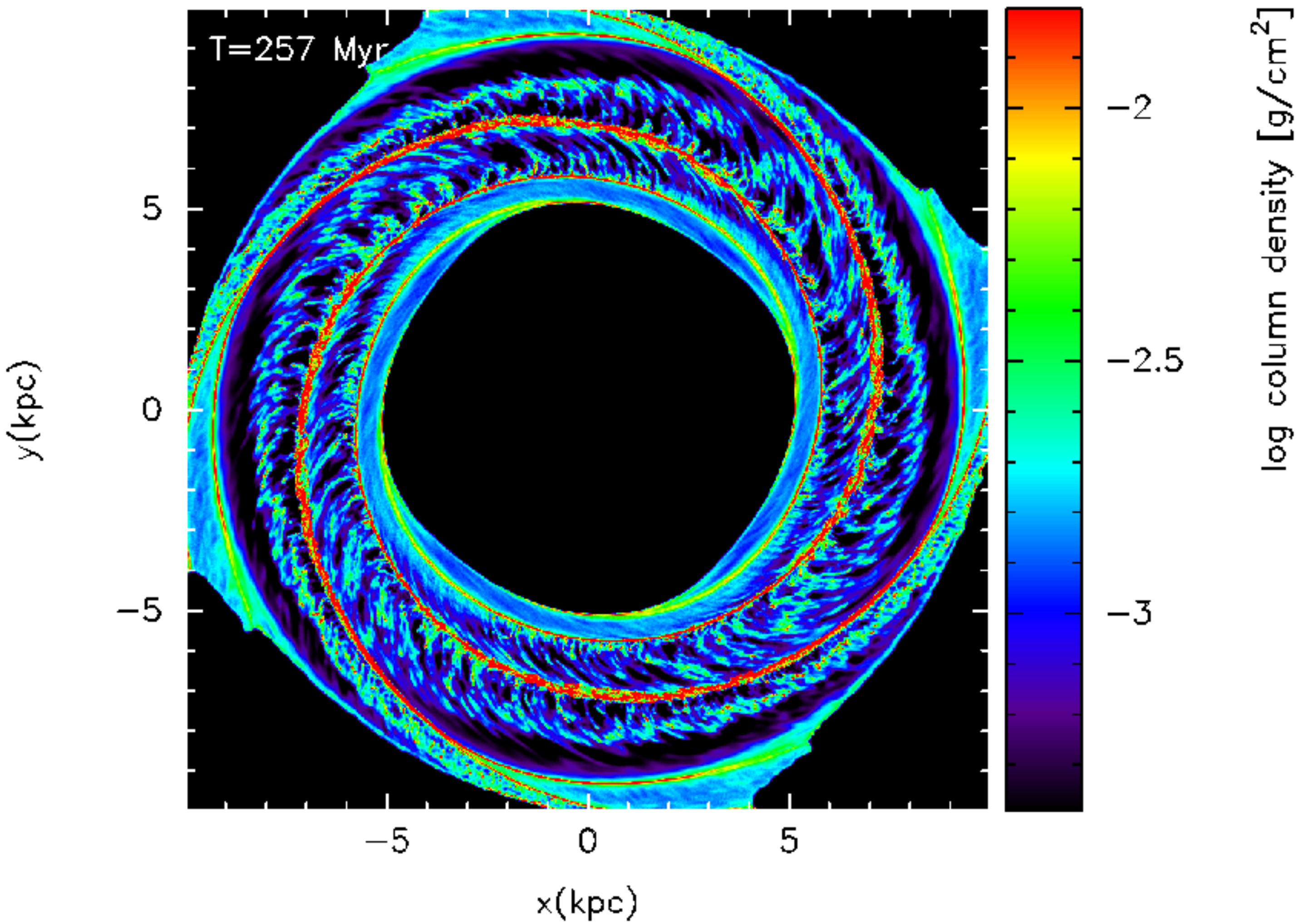
density/temperature fluctuations in warm atomar ISM are caused by *thermal/gravitational instability* and/or *supersonic turbulence*

some fluctuations are *dense* enough to *form* H_2 within "*reasonable time*"
→ *molecular cloud*

(Glover & Mac Low 2007a,b)

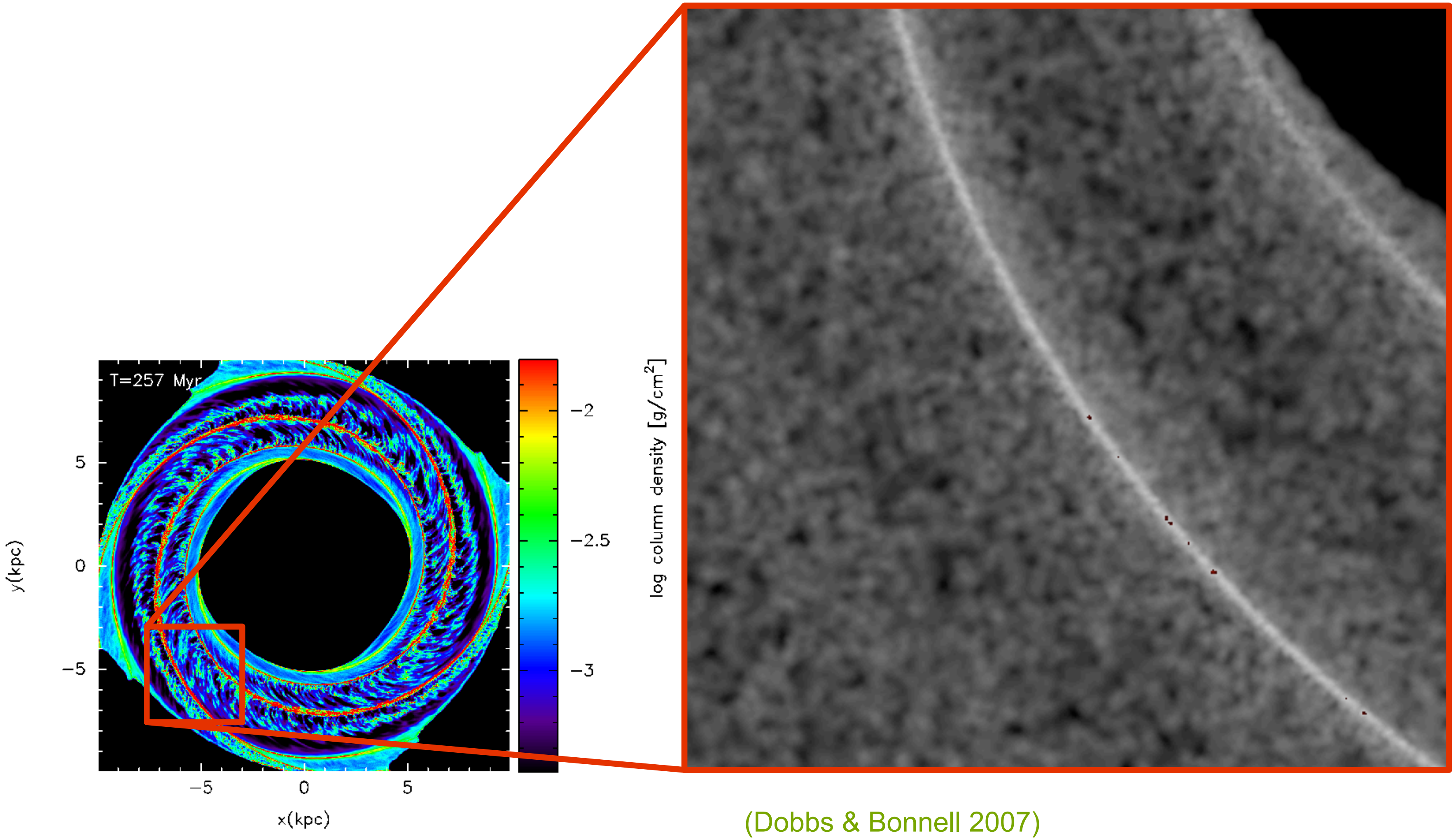
external perturbuations (i.e. potential changes) *increase* likelihood

molecular cloud formation



(from Dobbs, Glover, Clark, Klessen 2008)

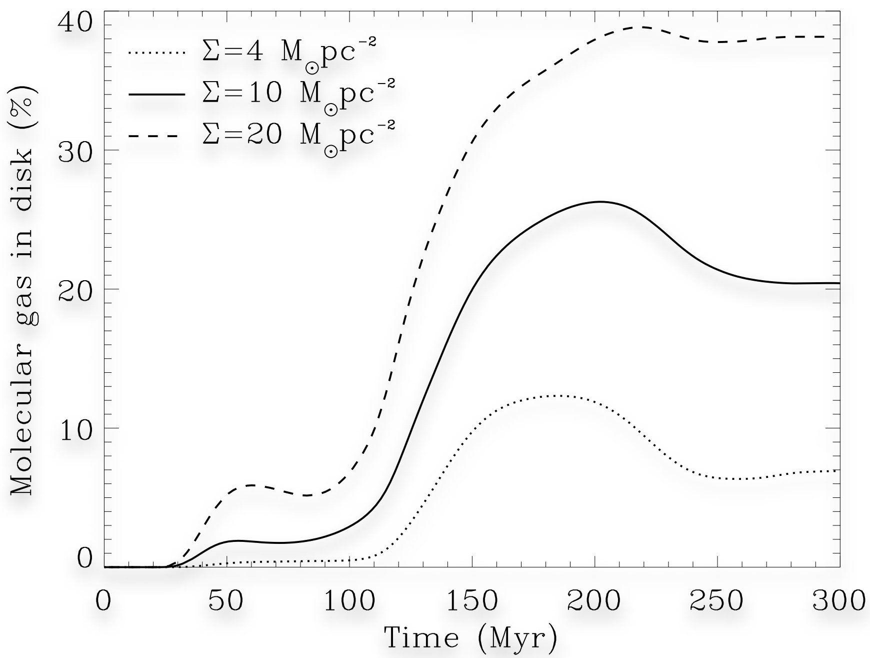
molecular cloud formation



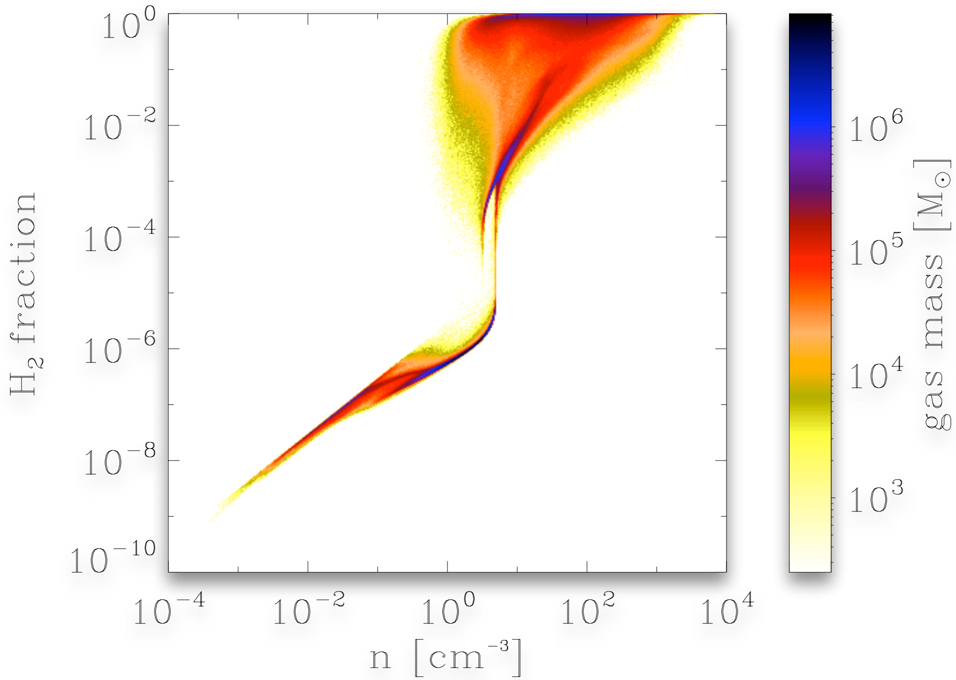
(Dobbs & Bonnell 2007)

molecular cloud formation

molecular gas fraction as function of time



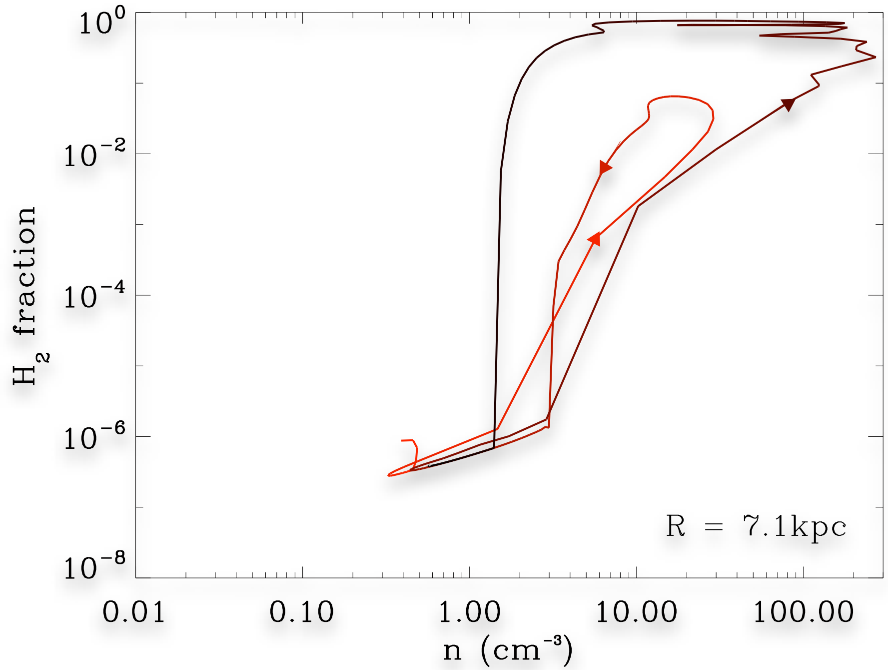
molecular gas fraction as function of density



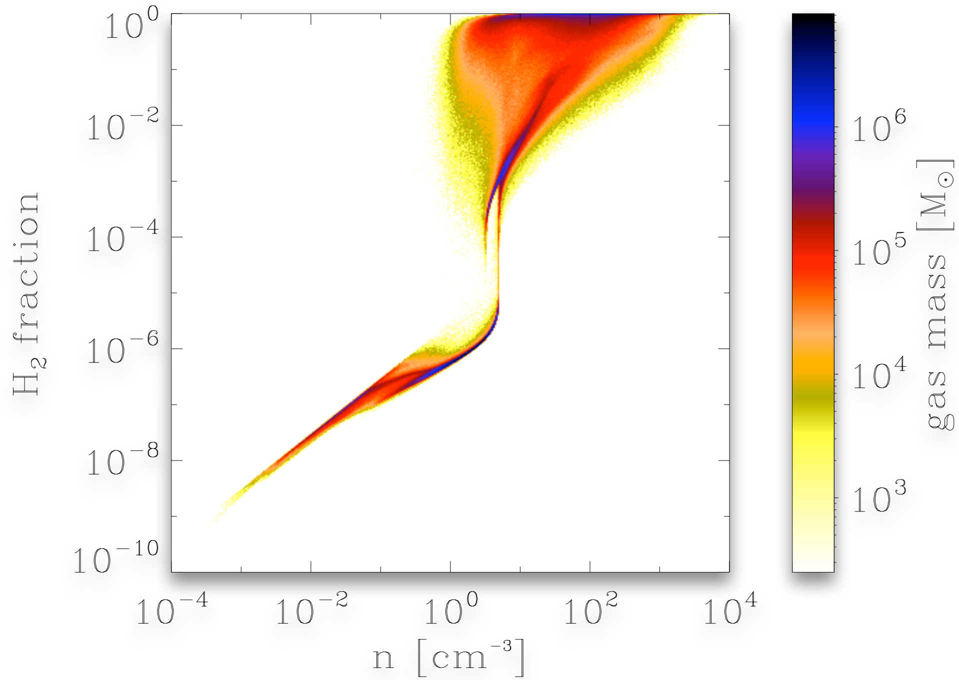
(Dobbs et al. 2008)

molecular cloud formation

molecular gas fraction of fluid element as function of time

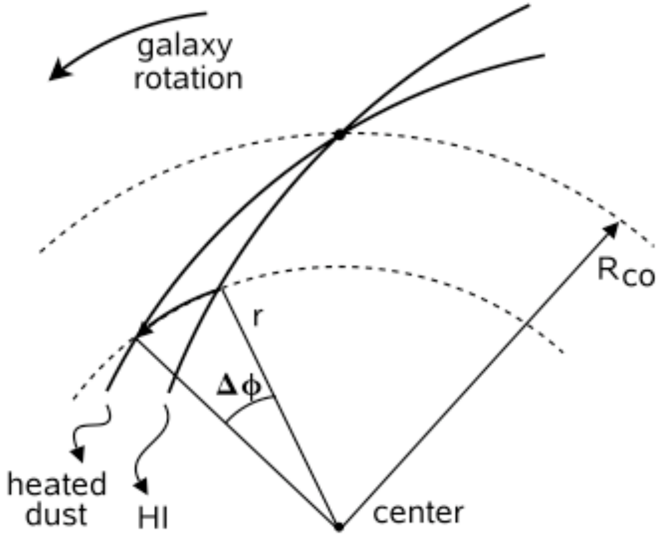


molecular gas fraction as function of density



(Dobbs et al. 2008)

observed timescales



Tamburro et al. (2008)

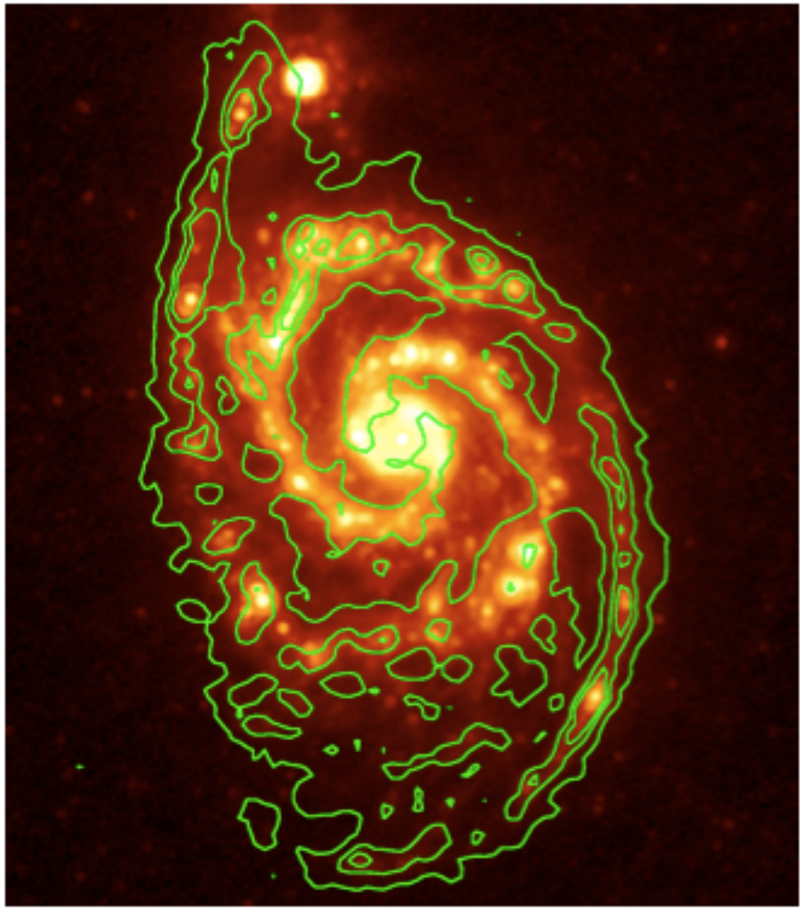
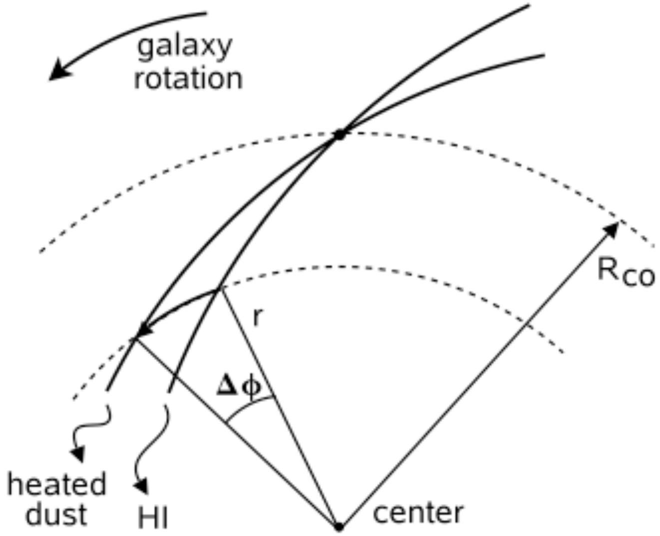


Fig. 1.— NGC 5194: the 24 μm band image is plotted in color scale; the H I emission map is overlayed with green contours.

observed timescales



Tamburro et al. (2008)

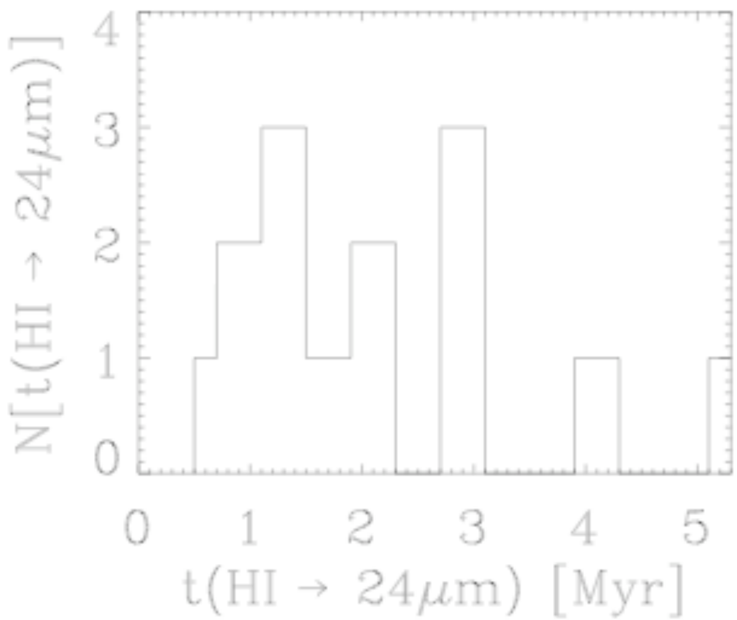
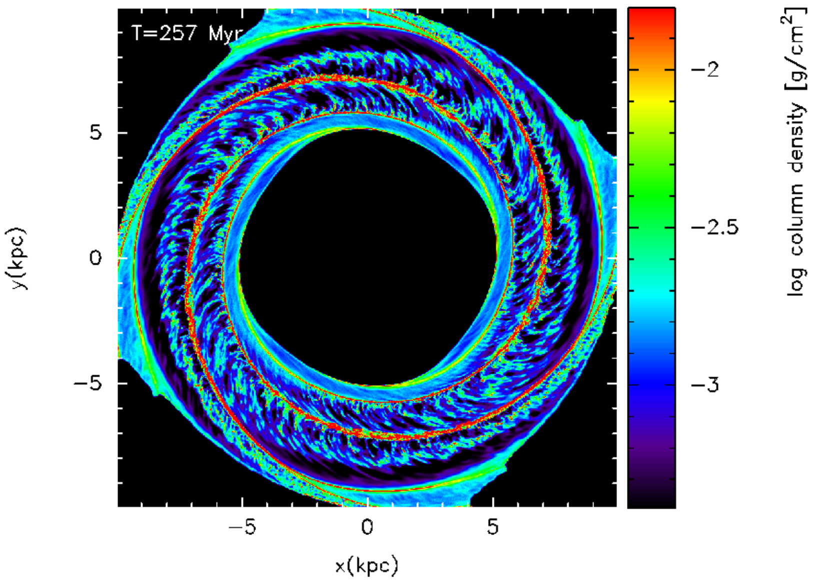


Fig. 5.— Histogram of the time scales $t_{HI \mapsto 24\,\mu m}$ derived from the fits in Figure 4 and listed in Table. 2 for the 14 sample galaxies listed in Table. 1. The timescales range between 1 and 4 Myr for almost all galaxies.

calculated timescales



Dobbs et al. (2008)

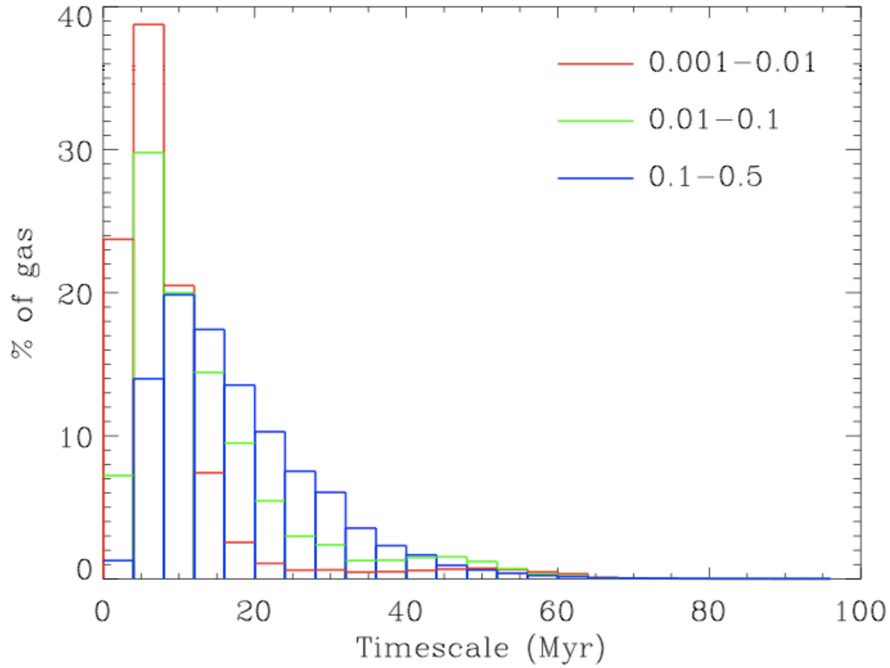


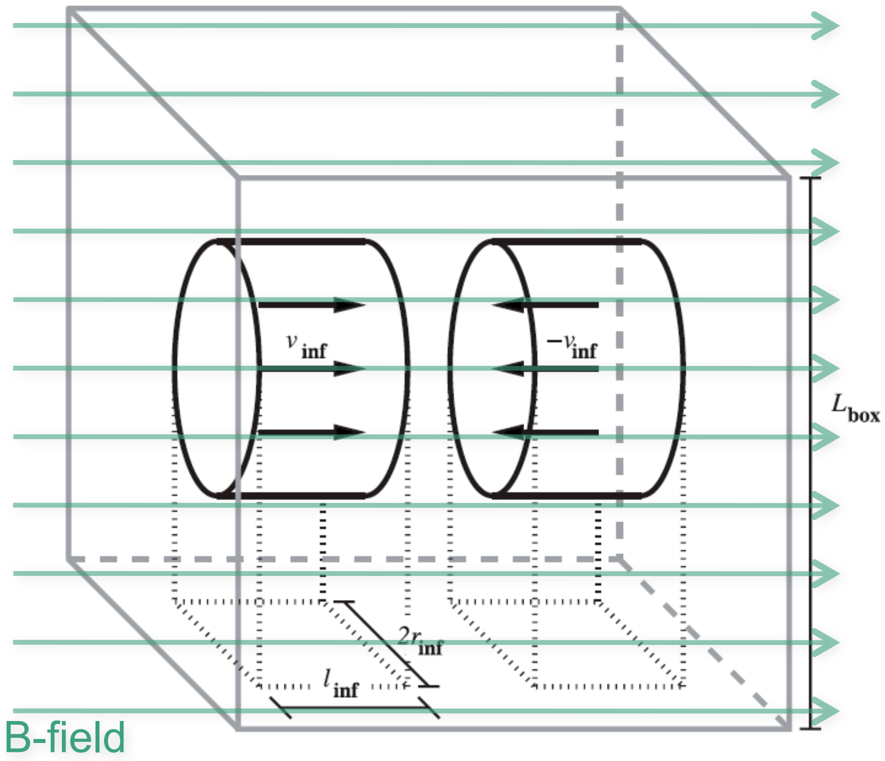
Figure 16. This histogram gives the distribution of timescales over which the gas reaches certain molecular gas fractions. The timescales denote the time for the H_2 fraction of a particle to increase from 0.001 to 0.01, 0.01 to 0.1 and 0.1 to 0.5, as indicated.

from atomic gas to molecular clouds

- let's look at the details:
  - how does molecular cloud material form in convergent flows, e.g., as triggered by spiral density waves...
  - do sequence of idealized numerical experiments
- questions
  - are molecular clouds truly "multi-phase" media?
  - turbulence? dynamical & morphological properties?
  - what is relation to initial & environmental conditions?
  - magnetic field structure?

convergent flows: set-up

numerical set-up



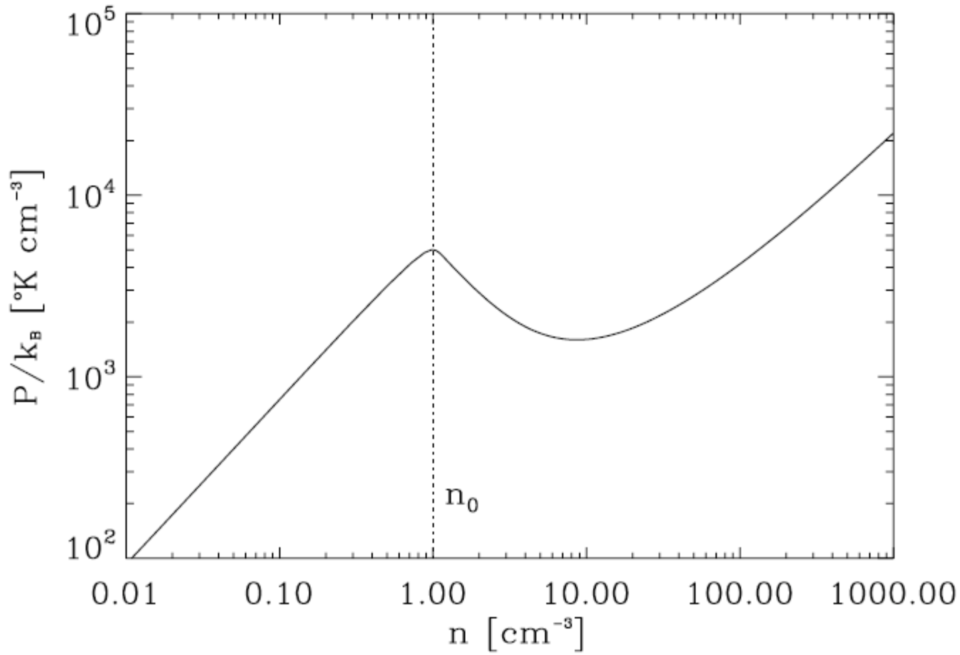
from Vazquez-Semadeni et al. (2007)

- convergent flow studies
  - atomic flows collide
  - cooling curve (soon chemistry)
  - gravity
  - magnetic fields
  - numerics: AMR, BGK, SPH

see studies by Banerjee et al., Heitsch et al., Hennebelle et al., Vazquez-Semadeni et al.

convergent flows: set-up

adopted cooling curve



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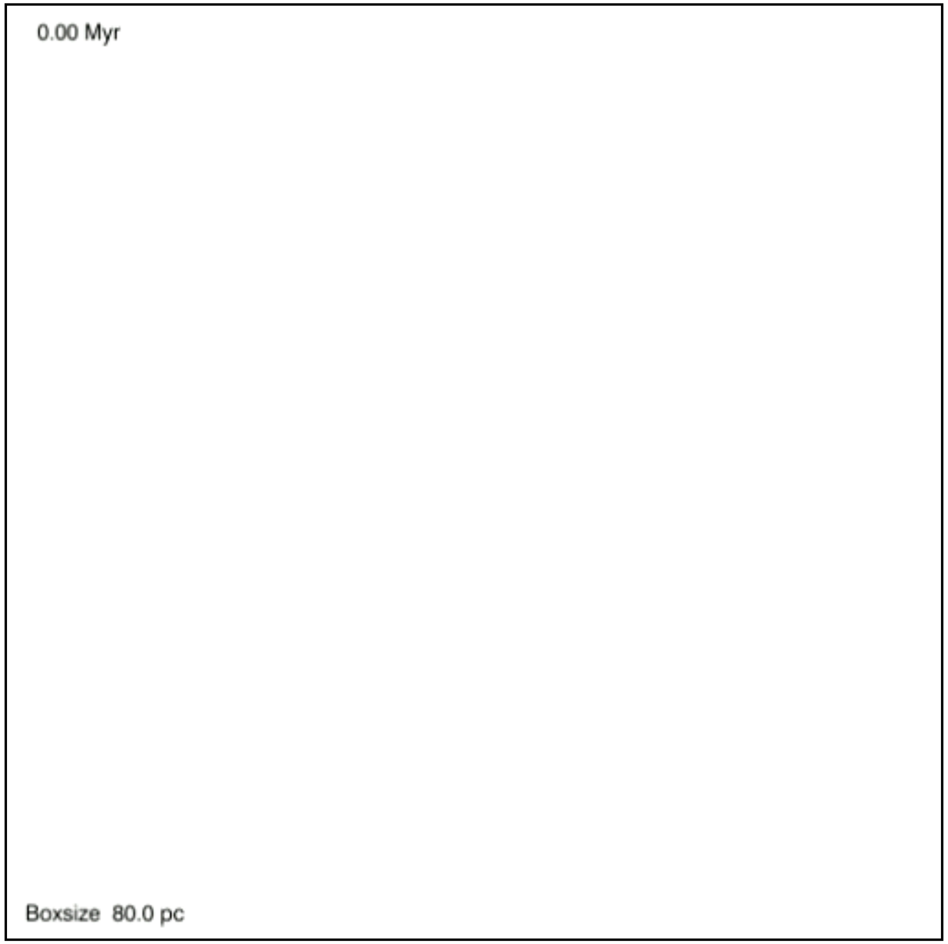
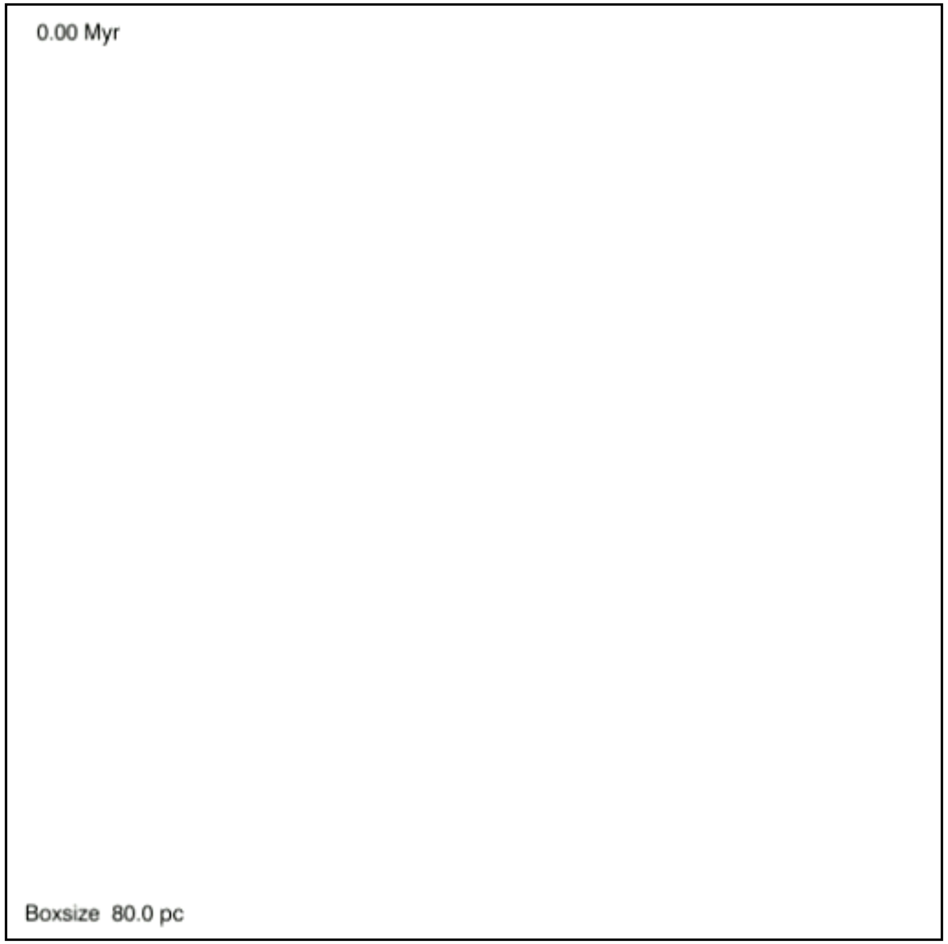
MC formation in convergent flows

the non-magnetic case





-edge-on view

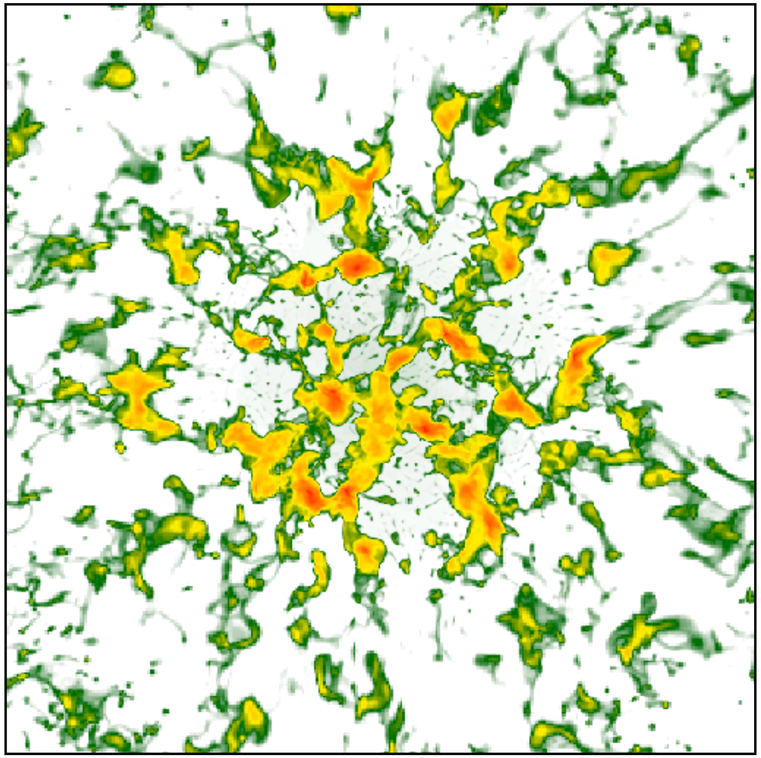
-face-on view

thermal instability + gravity creates complex molecular cloud structure:

MC formation in convergent flows

this simple set-up reproduces (and explains!) some of the main properties of MCs:

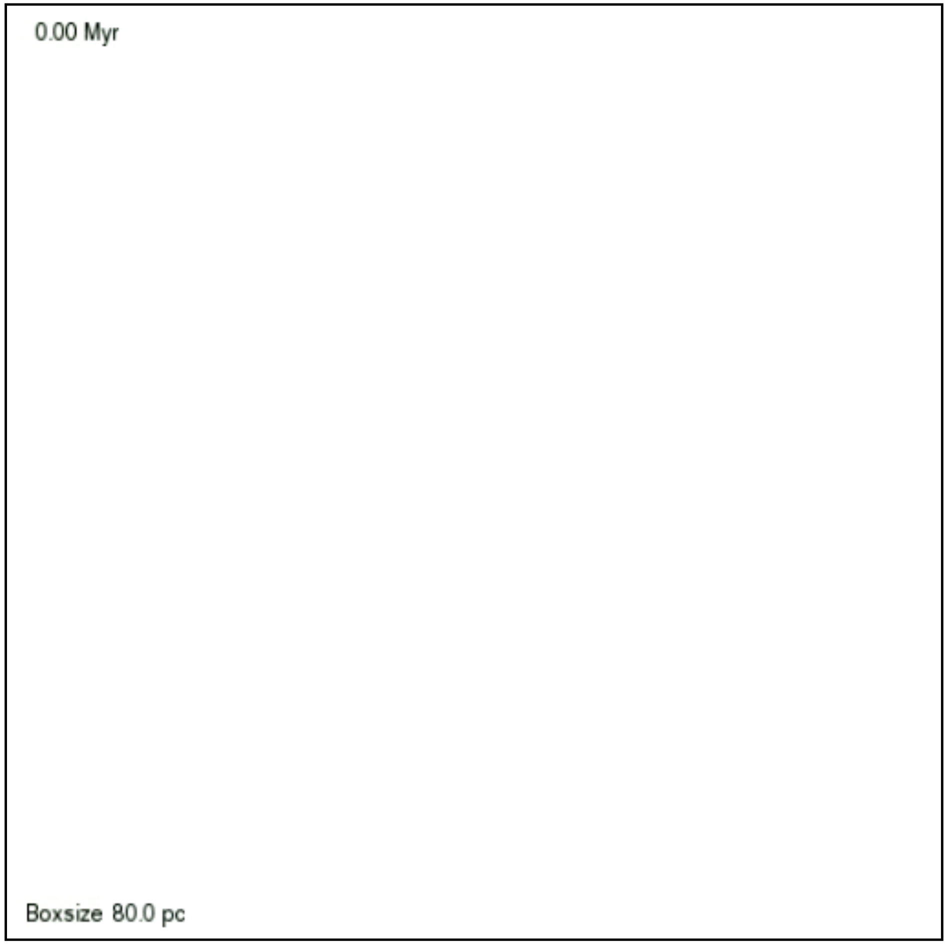
- highly patchy and clumpy
- high fraction of substructure
- cold dense molecular clumps coexist with warm atomic gas
- not a well bounded entity
- dynamical evolution (different star formation modes: from low mass to high mass SF?)



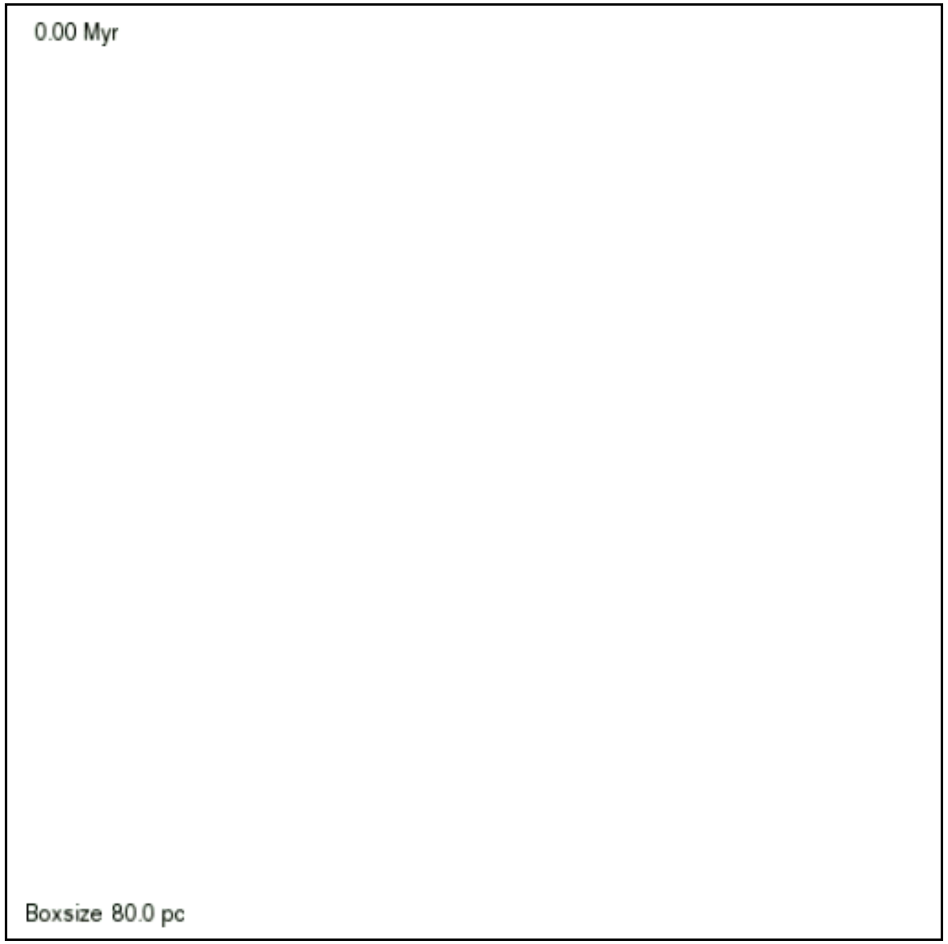
from Banerjee et al. (2008)
(see also studies by Hennebelle et al. and Vazquez-Semadeni et al. and Heitsch et al.)

MC formation in convergent flows

the weakly magnetized ($B_x = 1\mu G$) case



edge-on view



face-on view

from Banerjee et al. (2008)
(see also studies by Hennebelle et al. and Vazquez-Semadeni et al. and Heitsch et al.)

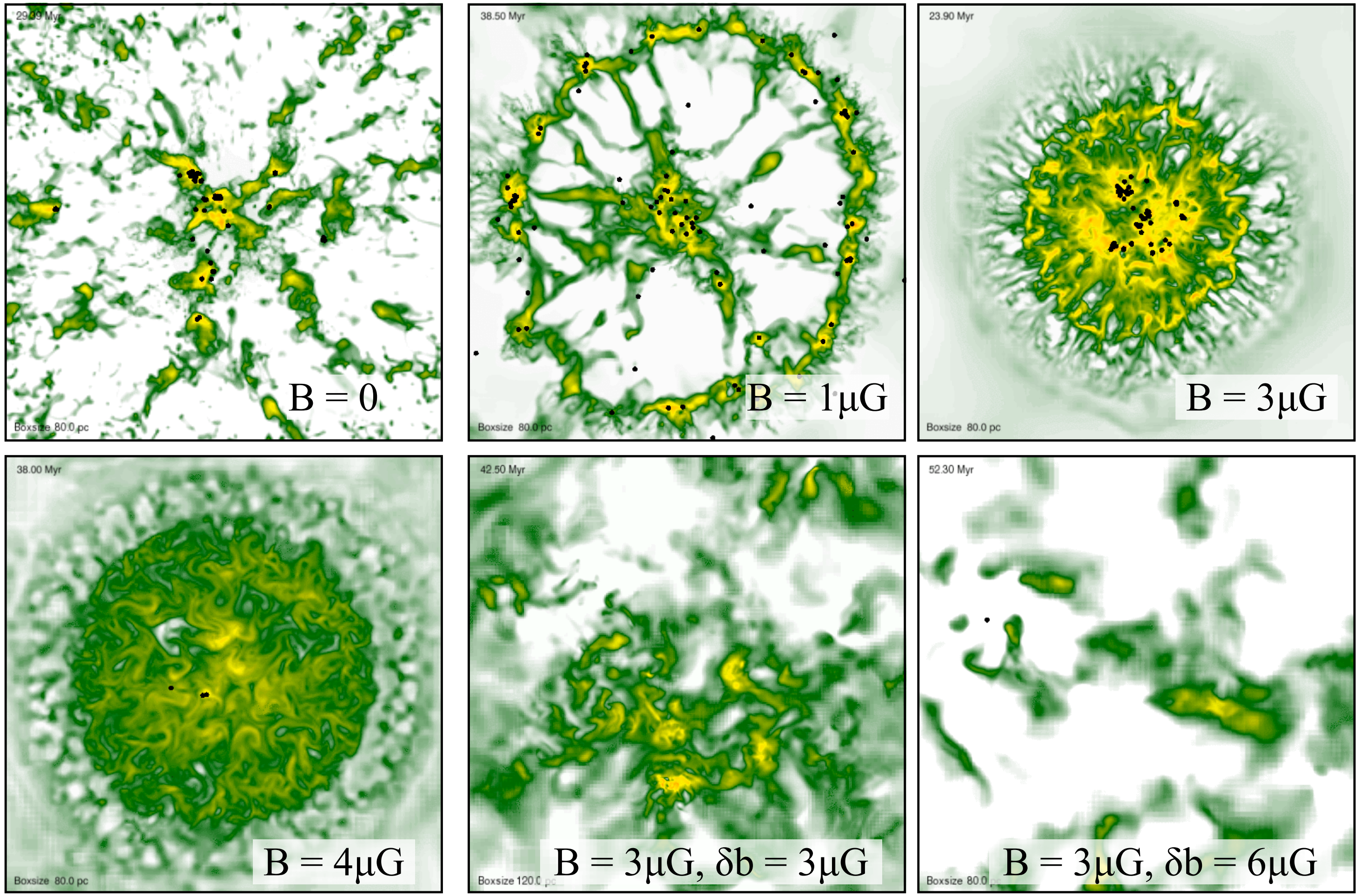
MC formation in convergent flows

with random component: $B_x = 3\mu G + \delta b = 3\mu G$

Banerjee et al. in prep.

face-on view

MC formation in convergent flows

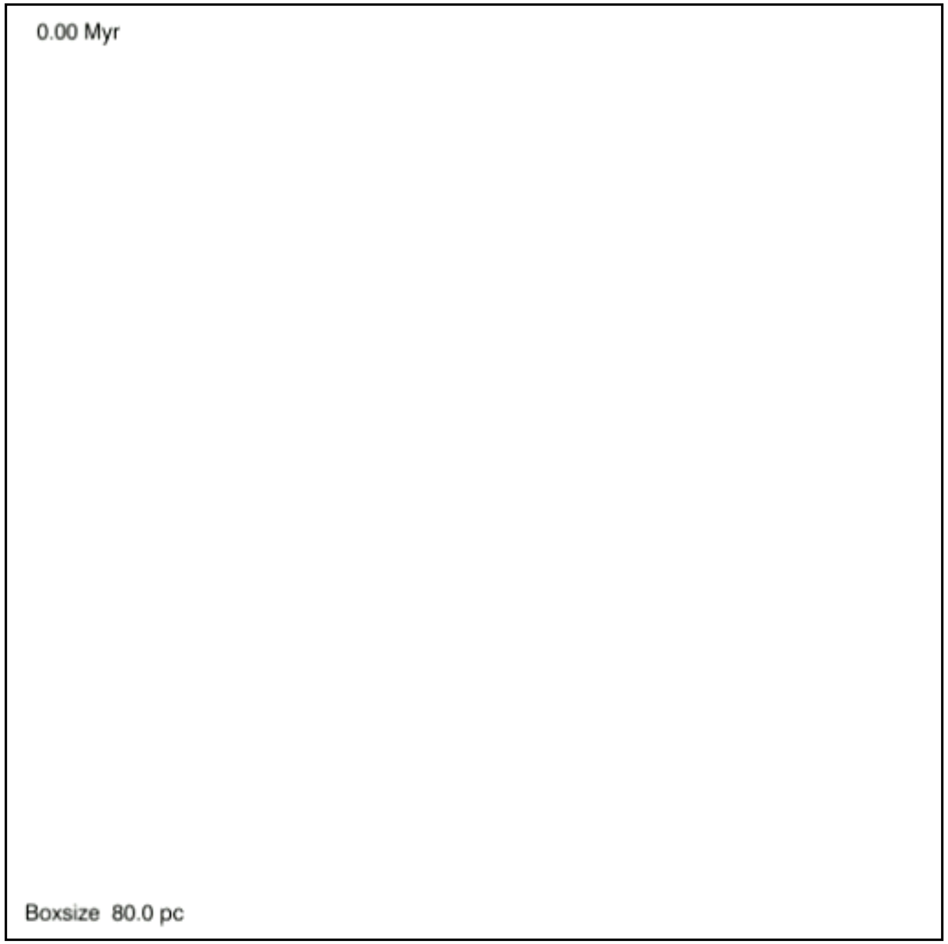


Morphology of the molecular cloud and star formation efficiency depends on the strength of the magnetic field

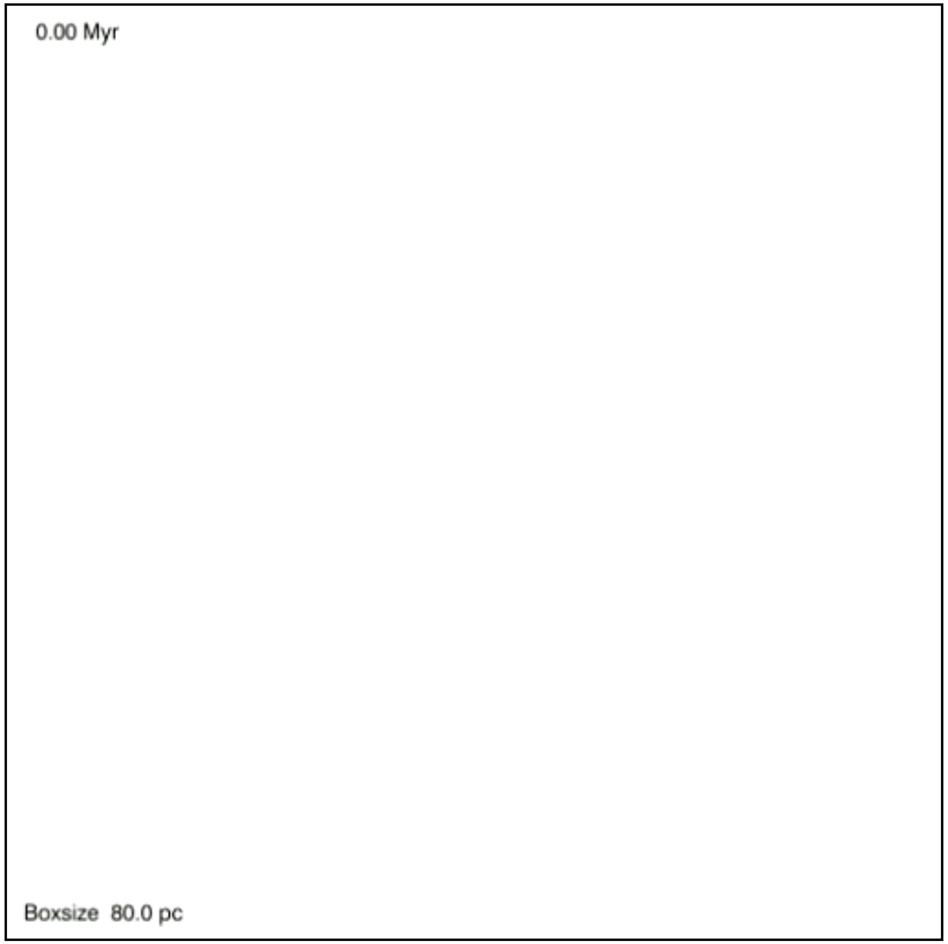
Banerjee et al. in prep.

MC formation in convergent flows

Influence of Ambipolar Diffusion: $B_x = 3\mu G$ (super-critical)



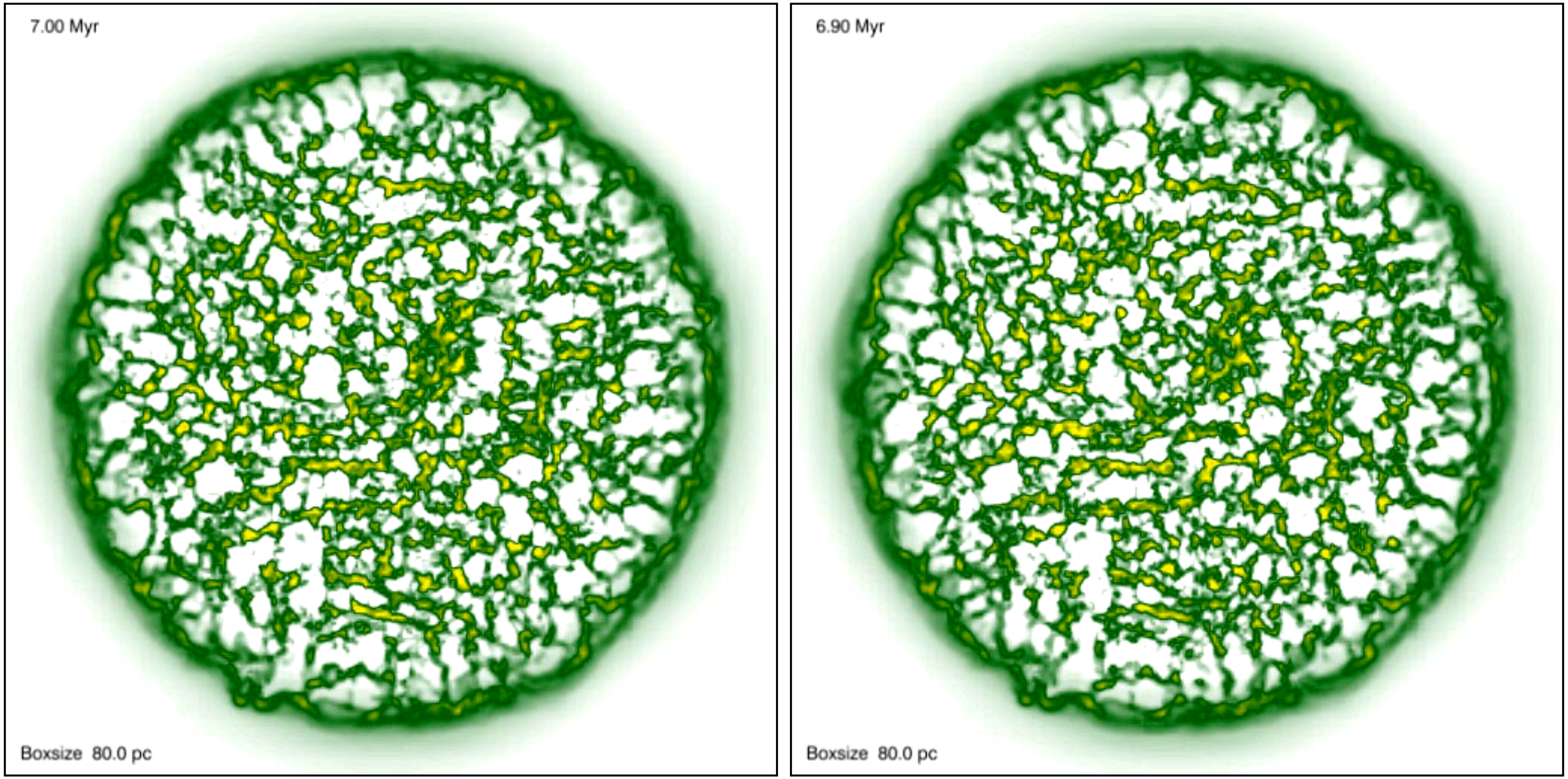
Ideal MHD



with AD

MC formation in convergent flows

Influence of Ambipolar Diffusion: $B_x = 4\mu G$ (critical)



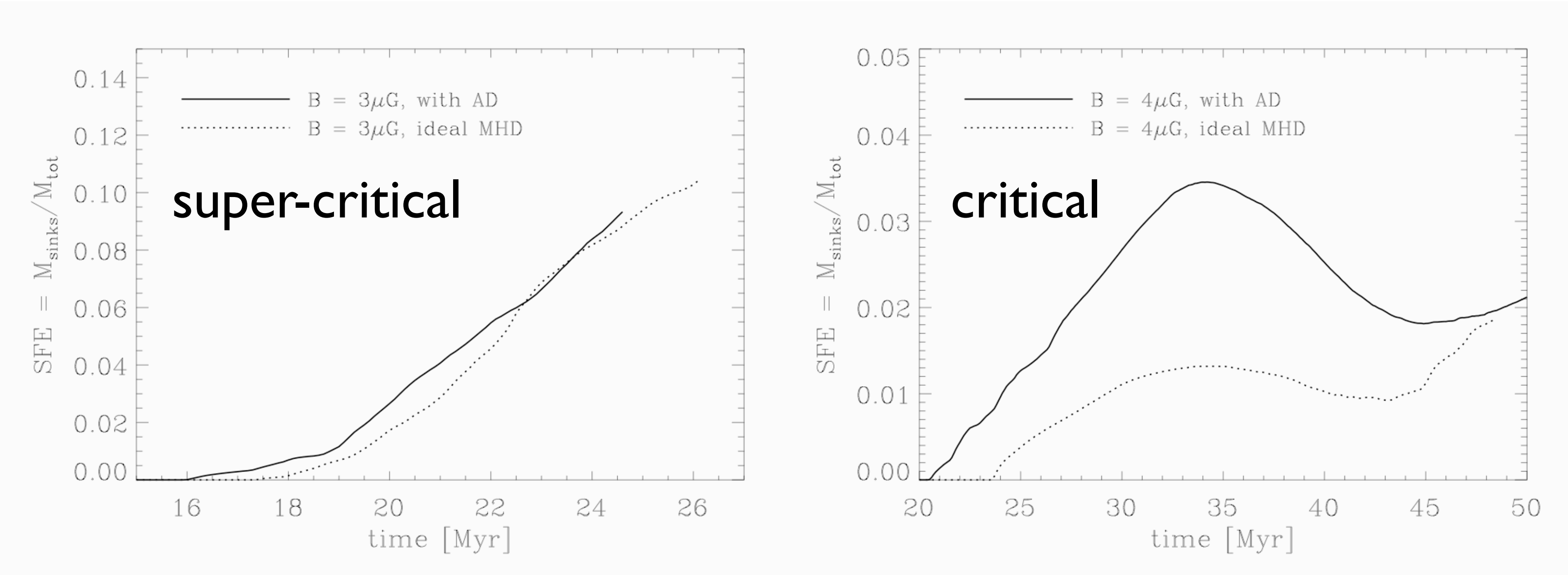
Ideal MHD

with AD

Banerjee et al. in prep.

MC formation in convergent flows

Influence of Ambipolar Diffusion

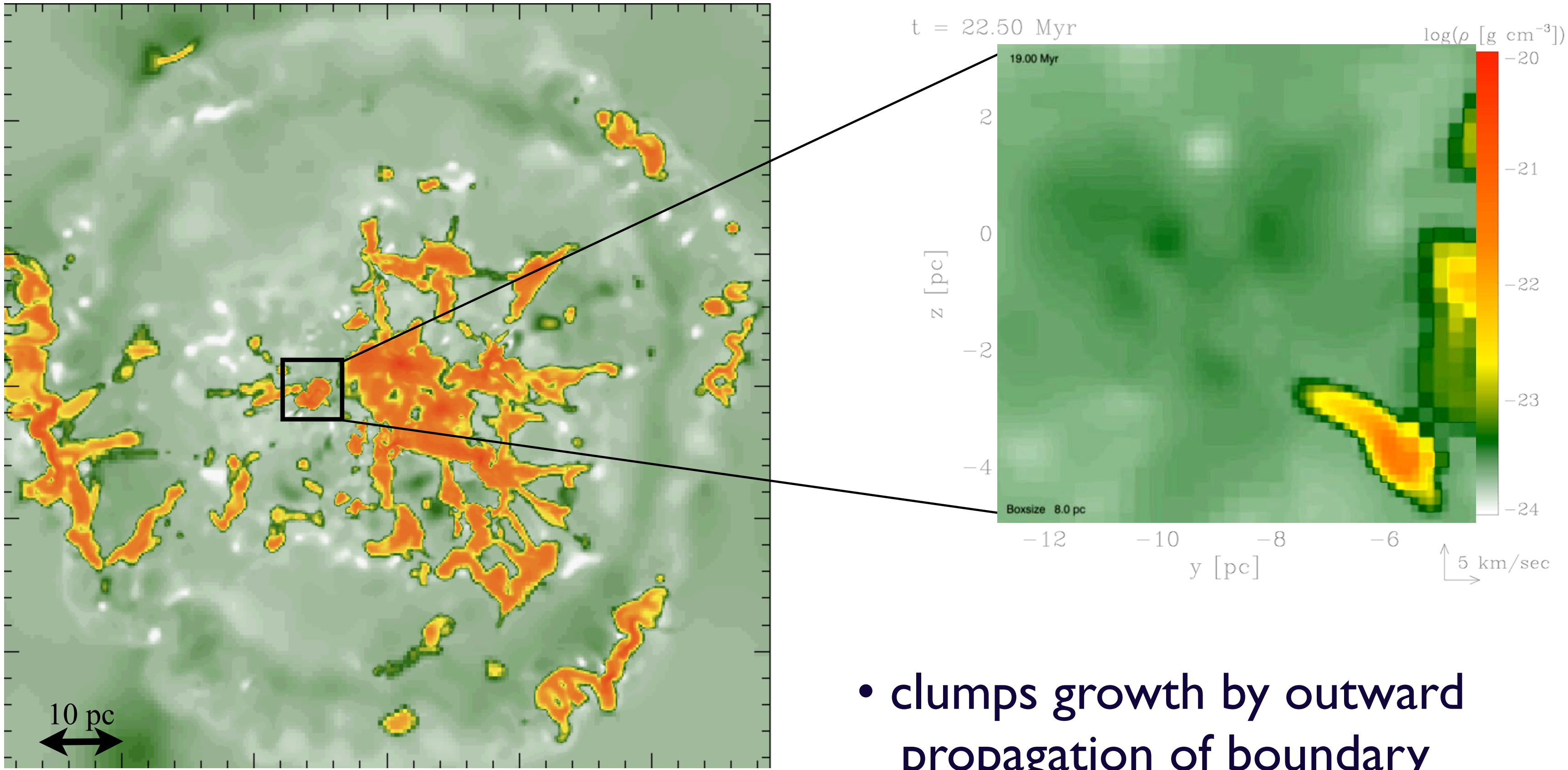


- Ambipolar diffusion is **not** a major player for star formation on molecular cloud scales
- this is different during protostellar collapse (Hennebelle et al.)

Banerjee et al. in prep.

MC formation in convergent flows

morphology and clump evolution



- MCs are inhomogeneous
- cold clumps embedded in warm atomic gas

- clumps growth by outward propagation of boundary layers and
- coalescence at later times

see studies by Banerjee et al., Heitsch et al., Hennebelle et al., Vazquez-Semadeni et al.

some results: growth of cores

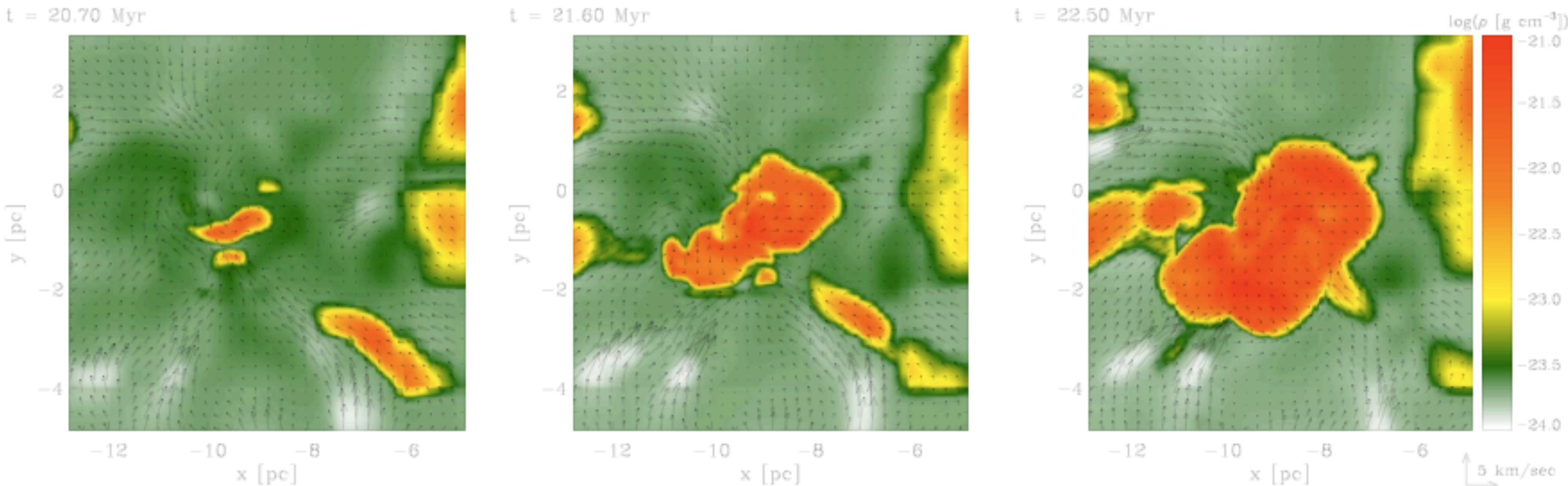
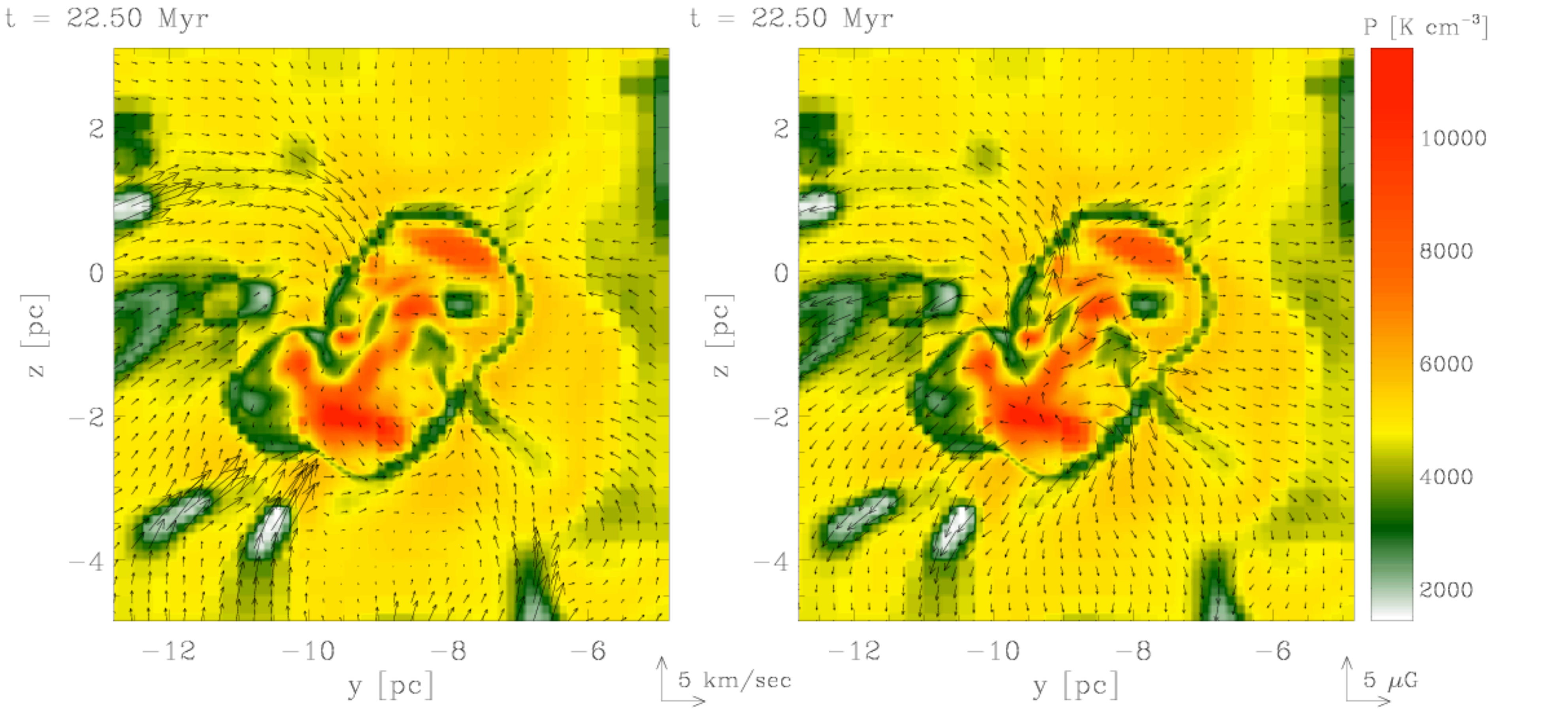


Figure 2. Shows the time evolution of a typical clump which initially develops out of the thermally unstable WNM in shock layers of turbulent flows. A small cold condensate grows by outward propagation of its boundary layer. Coalescence and merging with nearby clumps further increases the size and mass of these clumps. The global gravitational potential of the proto-cloud enhances the merging probability with time. The images show 2D slices of the density (logarithmic colour scale) and the gas velocity (indicated as arrows) in the plane perpendicular to the large scale flows.

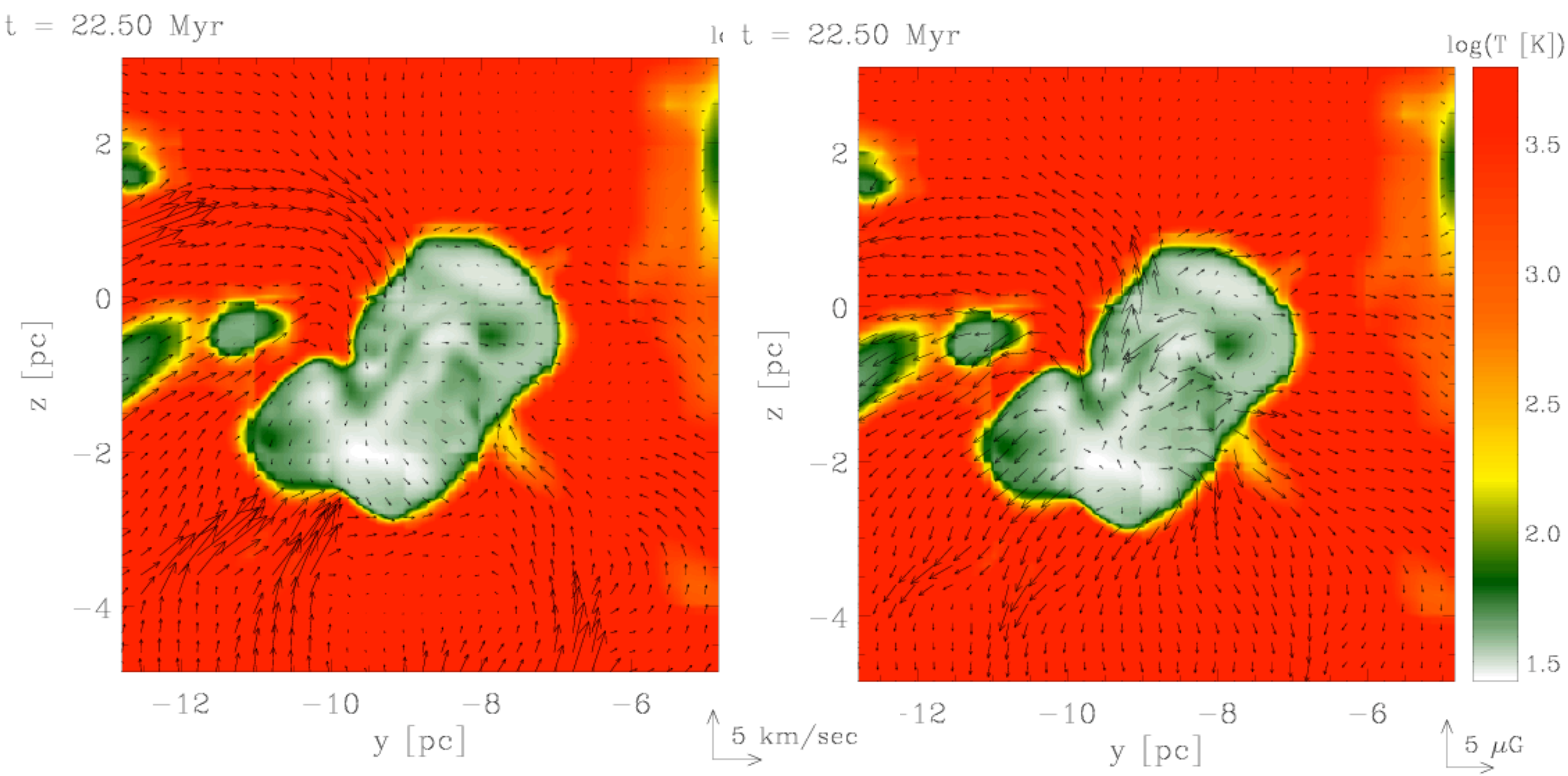
two phases of core growth:
(1) by *outward propagation* of *boundary layer* → Jeans sub-critical phase
(2) *core mergers* → super-Jeans → gravitational collapse & star formation
example: *Pipe nebula* ???

from Banerjee et al. (2008)



- cores roughly in pressure balance with surroudings
- relation between flow and magnetic field: mass flow mostly along field lines

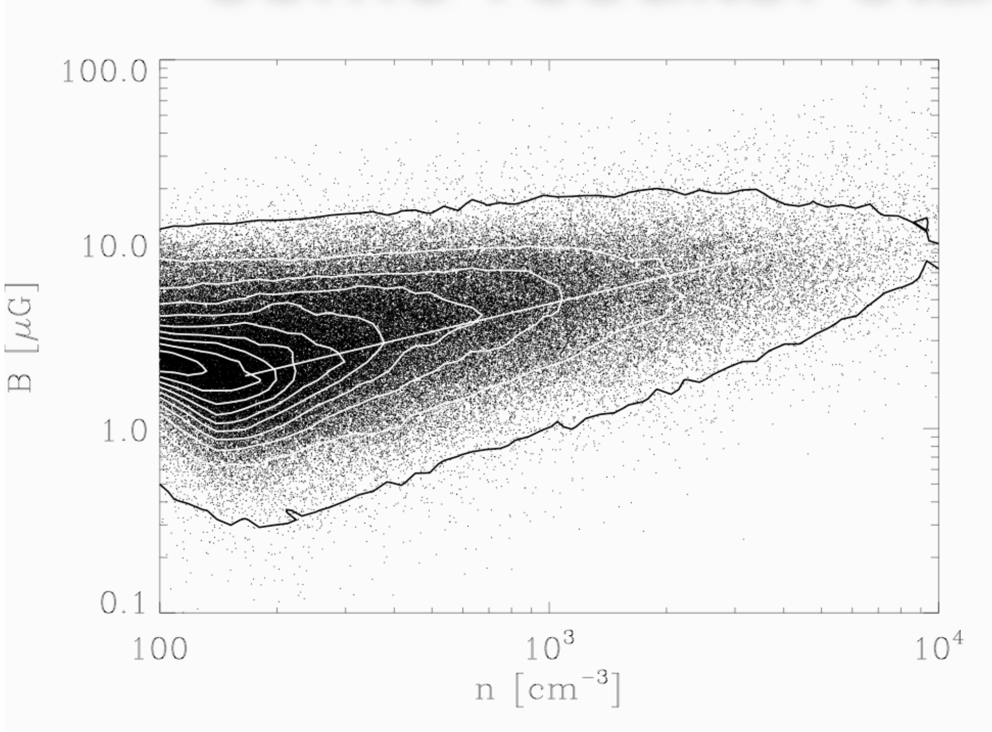
from Banerjee et al. (2008)



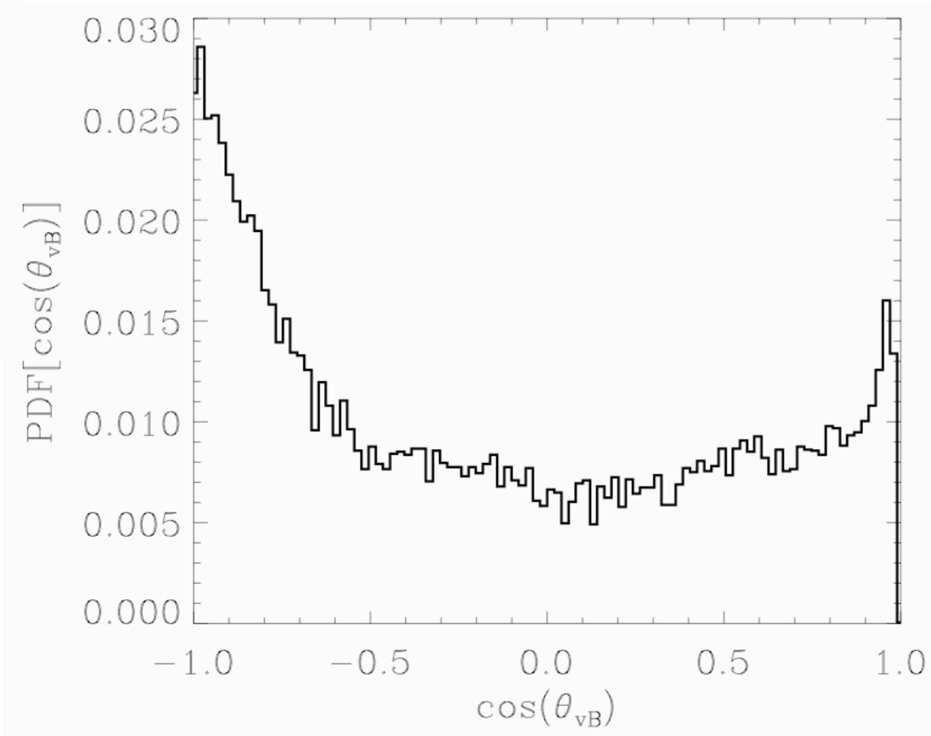
- typical core densities $n \sim 2 - 5 \times 10^3\ cm^{-3}$
- typical core temperatures $T \sim 30 - 50\ K$

from Banerjee et al. (2008)

some results: statistical correlations



- strong correlation of gas streams and magnetic field lines

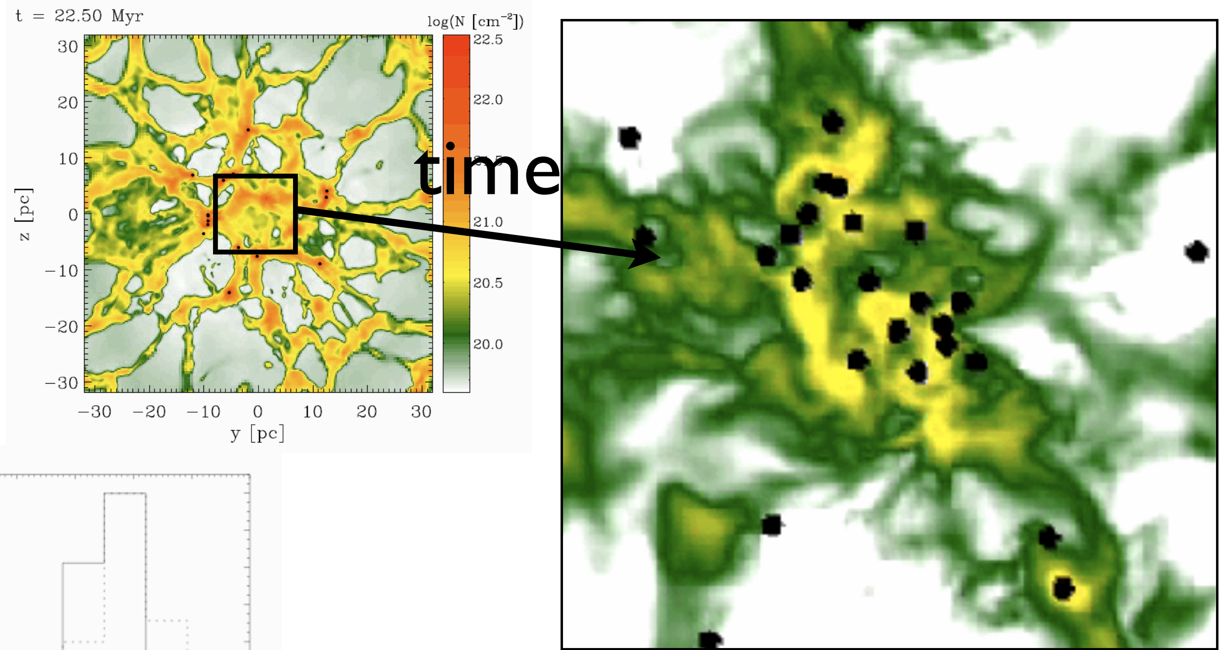


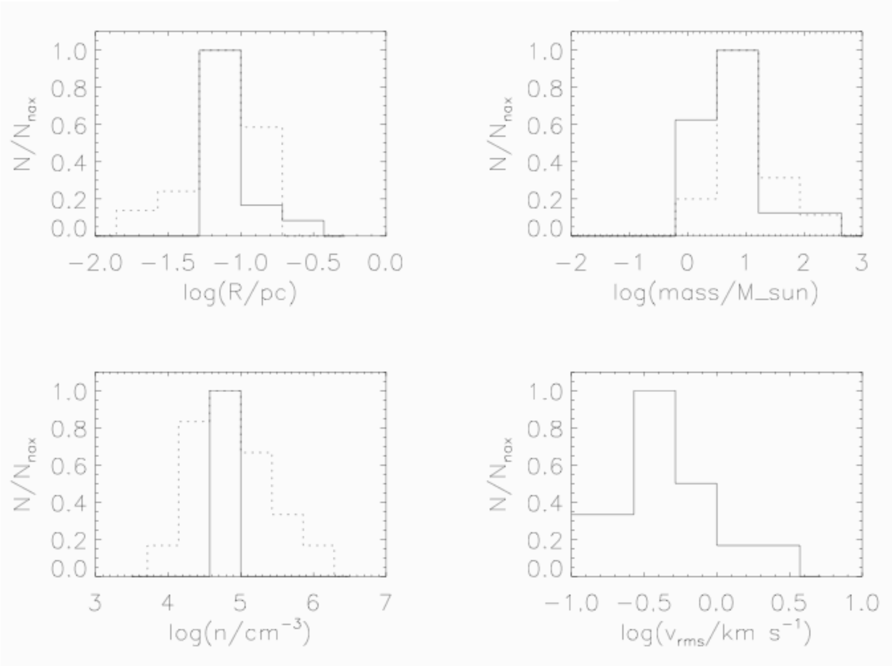
- large scatter of magnetic field strengths: sub- and super-critical cores exist
- median slope: $B \propto n^{0.5}$ (e.g. *Crutcher 1999*)

some results: loci of high-mass stars

global contraction phase

center of the cloud
→ birthplace for
massive stars?
(eg. Zinnecker & Yorke 2007)





comparison of core properties with observation of Cygnus X by *Motte et al 2007*

Vazquez-Semadeni et al. 2008

initial mass function

initial mass function

- what is the relation between molecular cloud fragmentation and the distribution of stars?
- important quantity: *IMF*
- equally important CAVEAT:
  "everyone" gets the right IMF
  → better look for secondary indicators
  - *stellar multiplicity*
  - protostellar *spin* (including disk)
  - *spatial distribution* + *kinematics* in young clusters
  - *magnetic field strength* and *orientation*

IMF

- distribution of stellar masses depends on
  - turbulent initial conditions
    --> mass spectrum of prestellar cloud cores
  - collapse and interaction of prestellar cores
    --> competitive accretion and *N*-body effects
  - thermodynamic properties of gas
    --> balance between heating and cooling
    --> EOS (determines which cores go into collapse)
  - (proto) stellar feedback terminates star formation
    ionizing radiation, bipolar outflows, winds, SN

(e.g. Larson 2003, Prog. Rep. Phys.; Mac Low & Klessen, 2004, Rev. Mod. Phys, 76, 125 - 194)

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example: model of Orion cloud

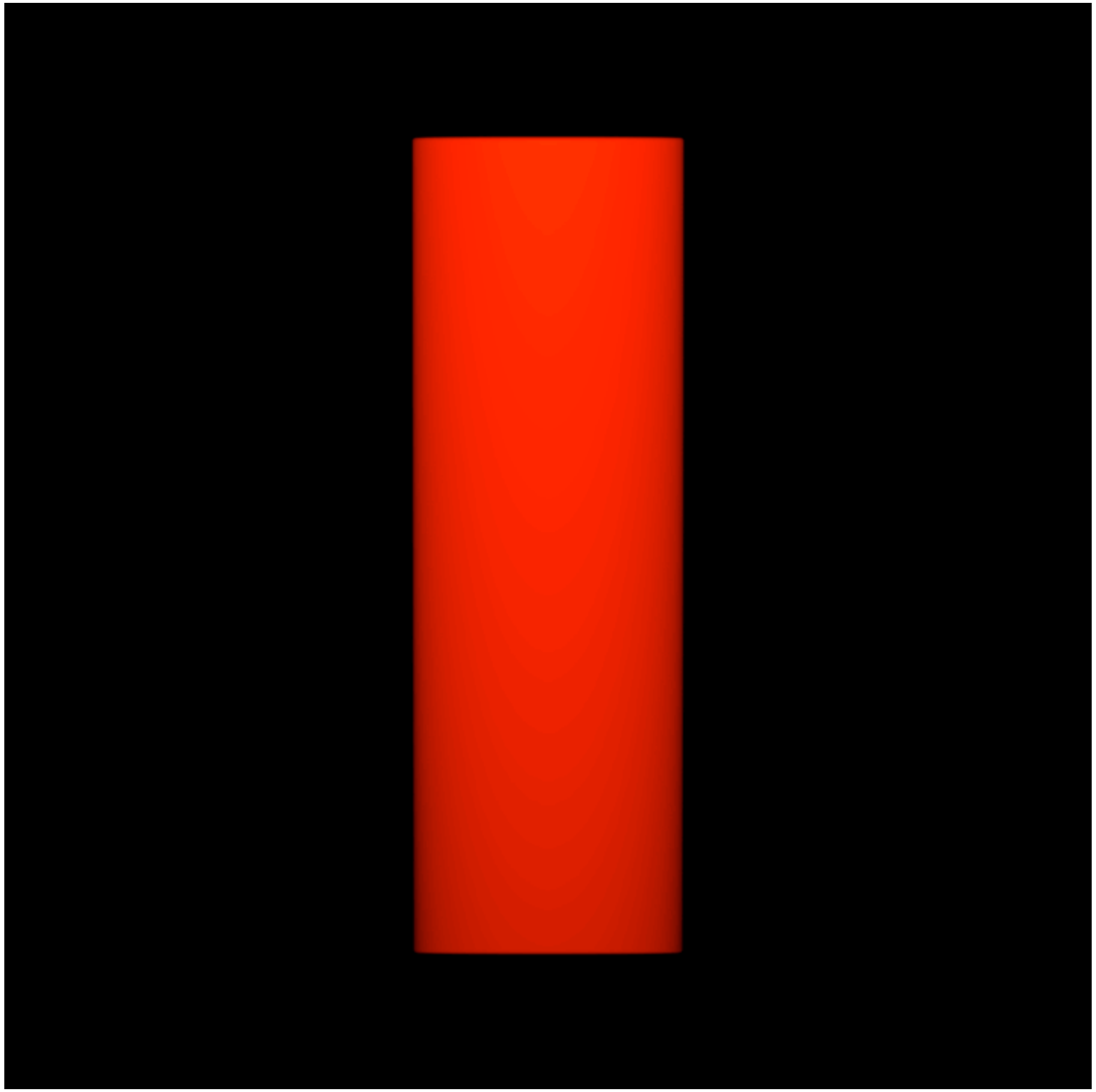
„model" of Orion cloud: 15.000.000 SPH particles, 10^4 M_{sun} in 10 pc, mass resolution 0,02 M_{sun}, forms ~2.500 „stars" (sink particles)

isothermal EOS, top bound, bottom unbound

has clustered as well as distributed „star" formation

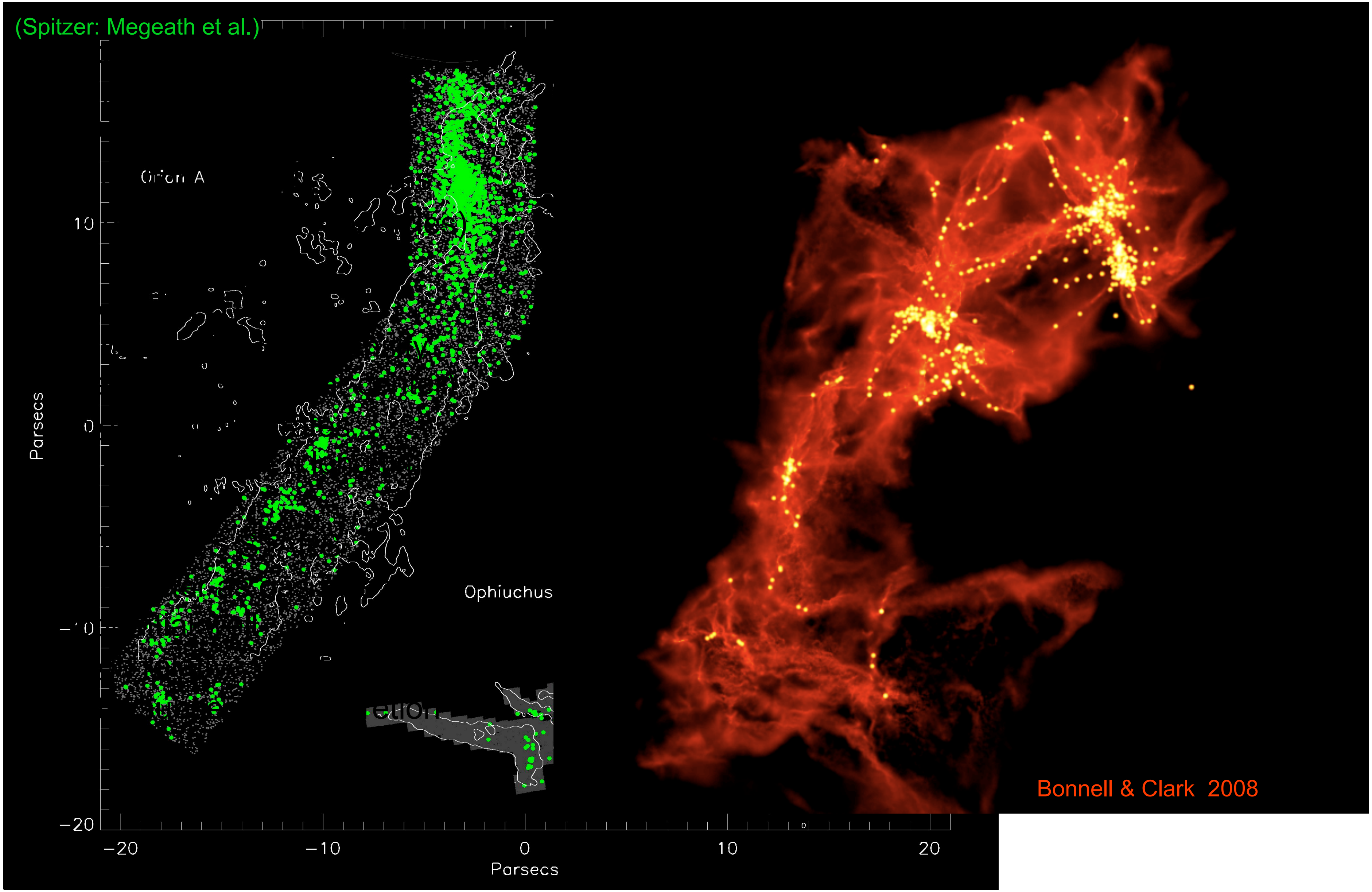
efficiency varies from 1% to 20%

develops full IMF
(distribution of sink particle masses)



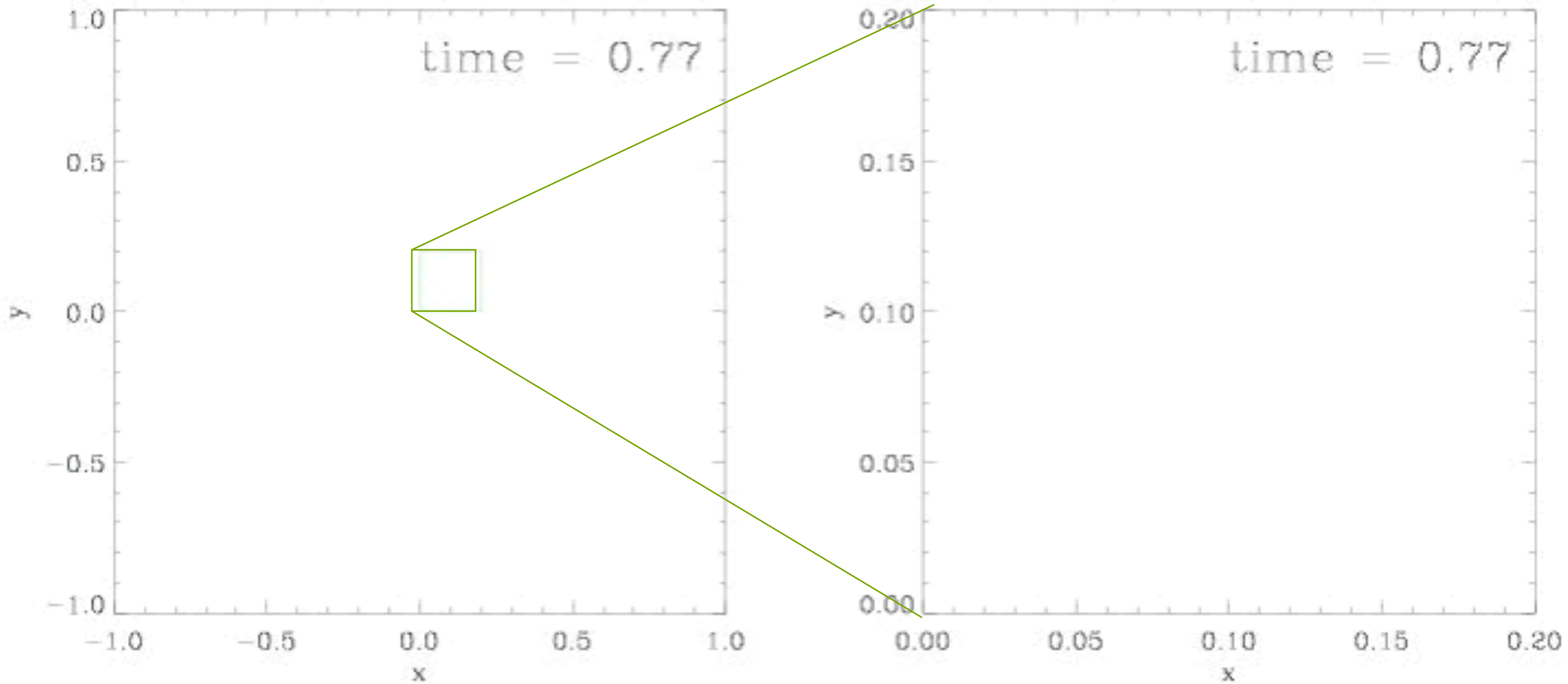
(Bonnell & Clark 2008)

example: model of Orion cloud



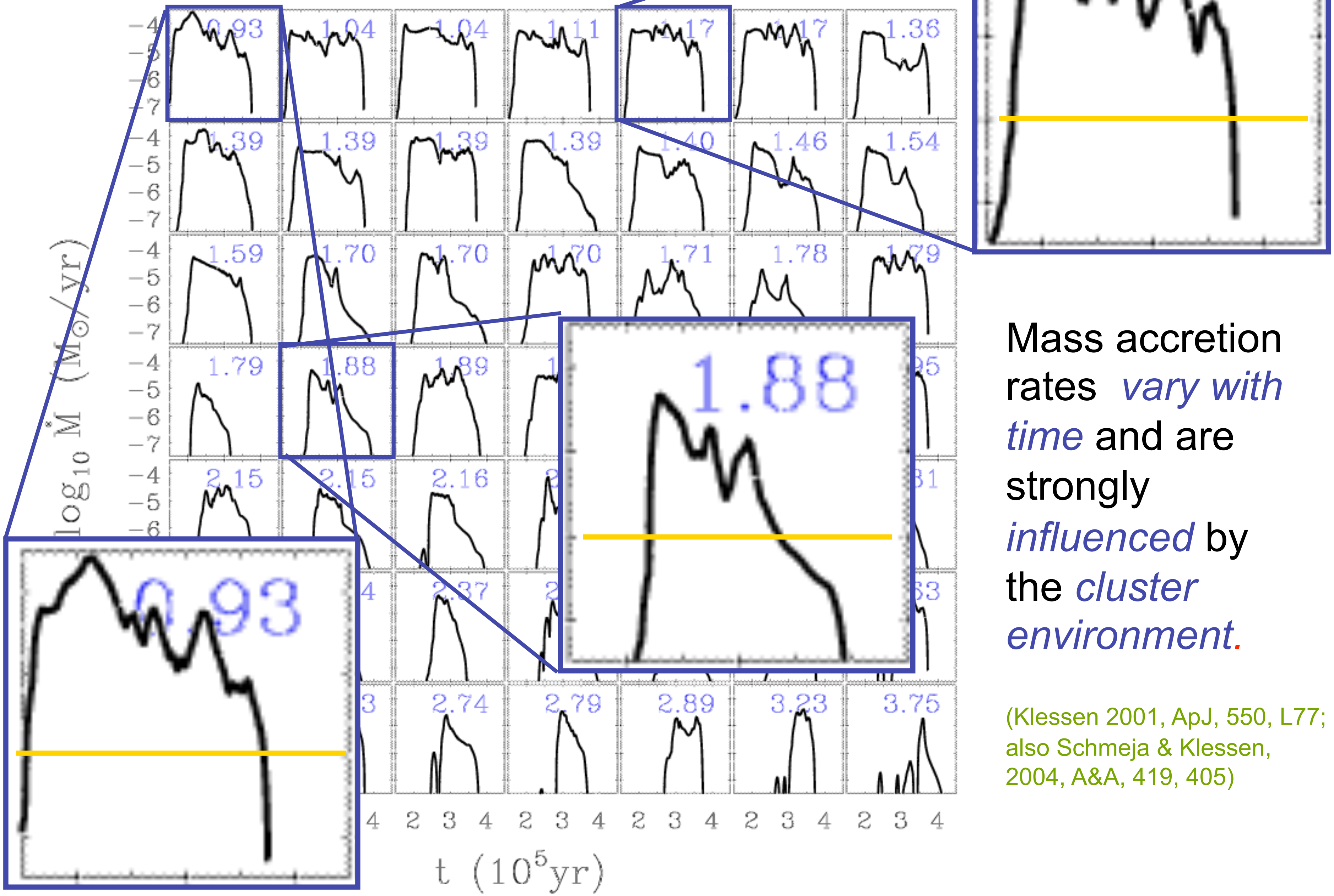
dynamics of nascent star cluster

in dense clusters protostellar interaction may be come important!



Trajectories of protostars in a nascent dense cluster created by gravoturbulent fragmentation
(from Klessen & Burkert 2000, ApJS, 128, 287)

accretion rates in clust



Mass accretion rates *vary with time* and are strongly *influenced* by the *cluster environment*.

(Klessen 2001, ApJ, 550, L77; also Schmeja & Klessen, 2004, A&A, 419, 405)

IMF

- distribution of stellar masses depends on
  - turbulent initial conditions
    --> mass spectrum of prestellar cloud cores
  - collapse and interaction of prestellar cores
    --> competitive accretion and *N*-body effects
  - *thermodynamic properties of gas*
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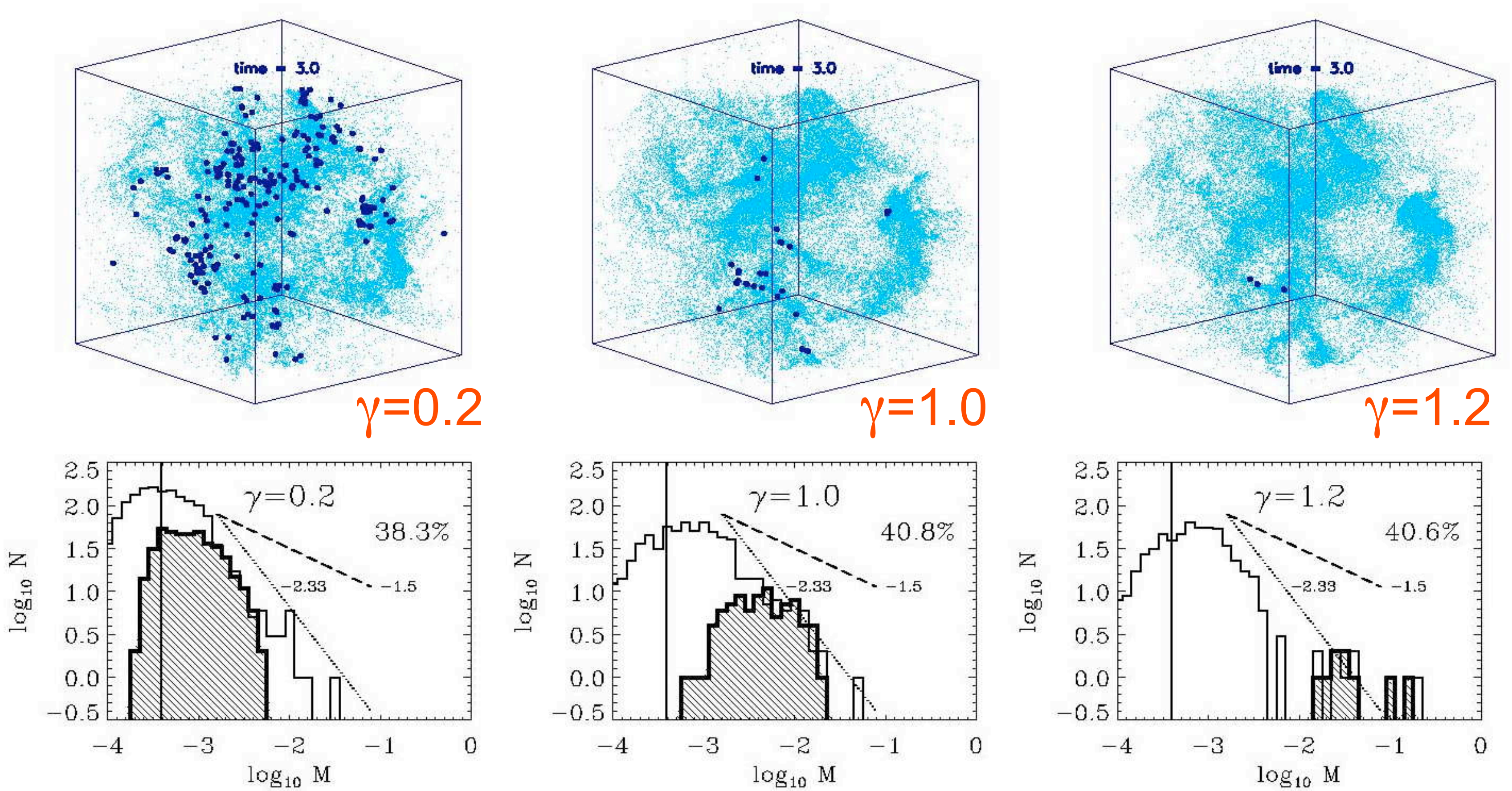
(e.g. Larson 2003, Prog. Rep. Phys.; Mac Low & Klessen, 2004, Rev. Mod. Phys, 76, 125 - 194)

dependency on EOS

- degree of fragmentation depends on *EOS!*

- polytropic EOS: **p** $\propto \rho^{\gamma}$
- $\gamma<1$: dense cluster of low-mass stars
- $\gamma>1$: isolated high-mass stars
- (see Li, Klessen, & Mac Low 2003, ApJ, 592, 975; also Kawachi & Hanawa 1998, Larson 2003)

dependency on EOS



for $\gamma<1$ fragmentation is enhanced → *cluster of low-mass stars*

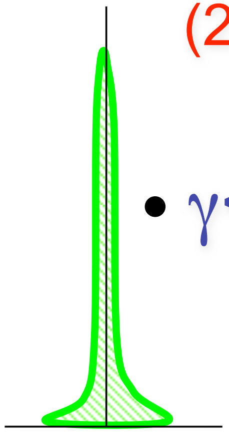
for $\gamma>1$ it is suppressed → formation of *isolated massive stars*

(from Li, Klessen, & Mac Low 2003, ApJ, 592, 975)

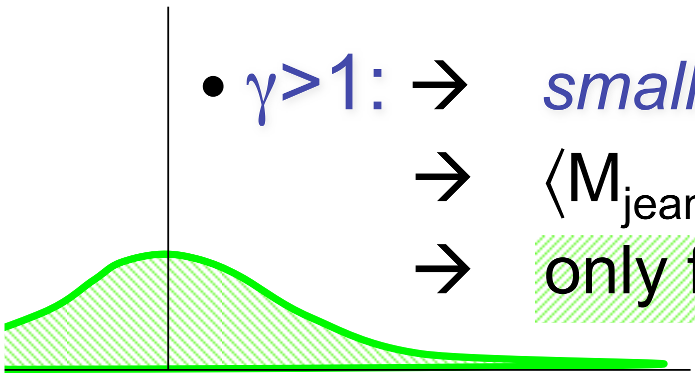
how does that work?

(1) $\mathbf{p} \propto \rho^{\gamma} \rightarrow \rho \propto \mathbf{p}^{1/\gamma}$

(2) $\mathbf{M_{jeans}} \propto \gamma^{3/2} \rho^{(3\gamma-4)/2}$



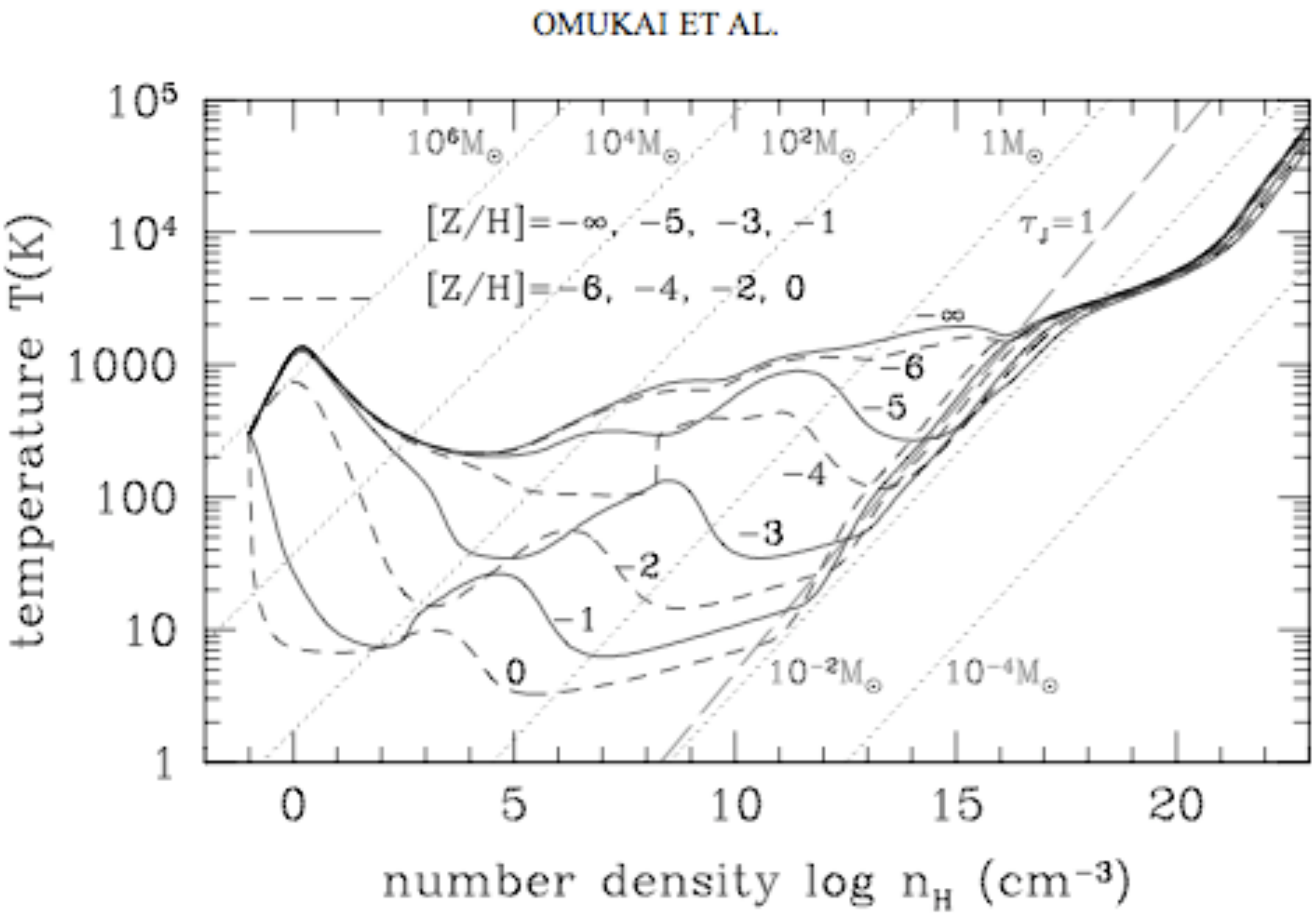
- $\gamma<1$: → *large* density excursion for given pressure
  - → $\langle M_{jeans} \rangle$ becomes small
  - → number of fluctuations with $M > M_{jeans}$ is large



- $\gamma>1$: → *small* density excursion for given pressure
  - → $\langle M_{jeans} \rangle$ is large
  - → only few and massive clumps exceed M_{jeans}

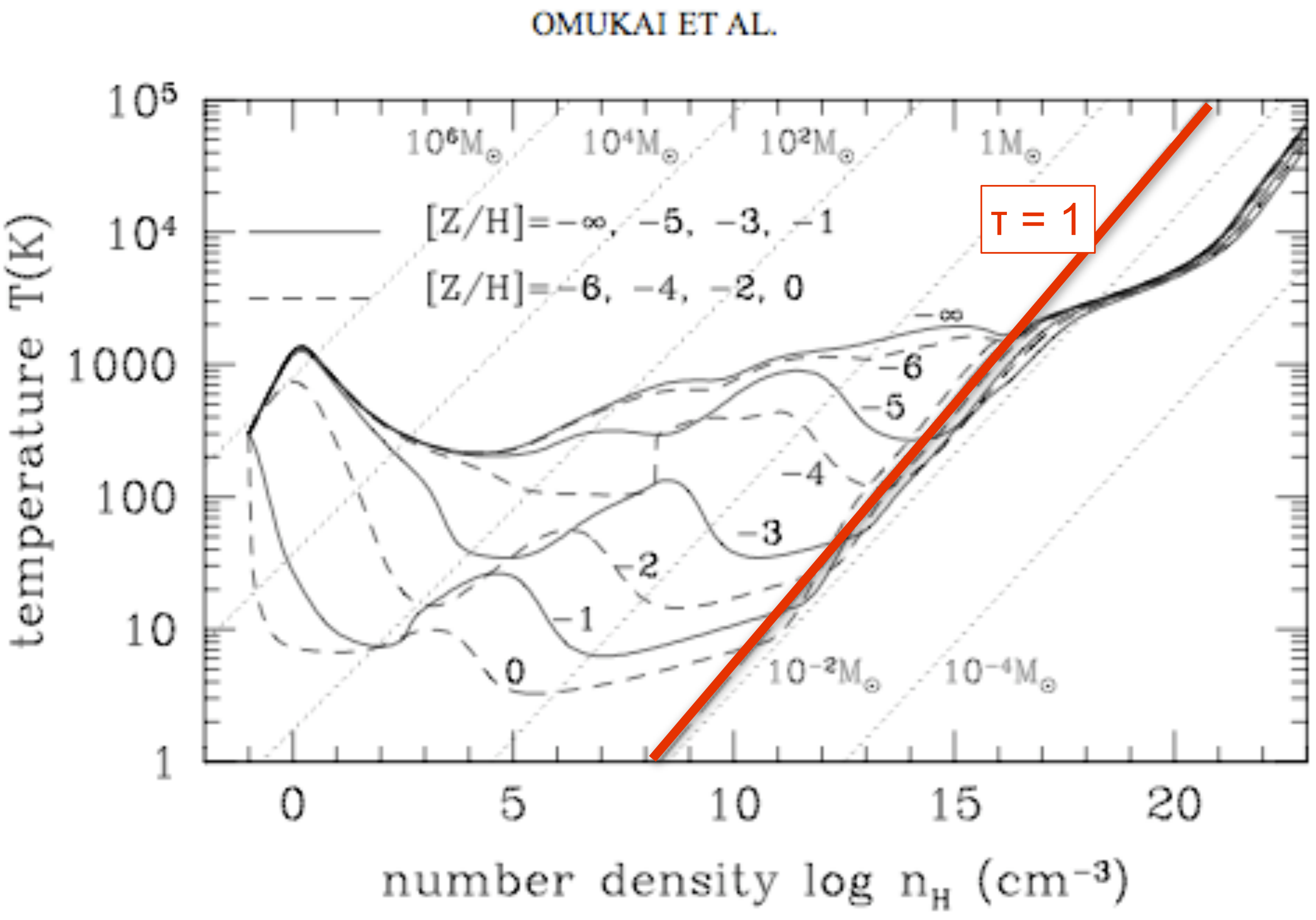
EOS in different environments

EOS as function of metallicity



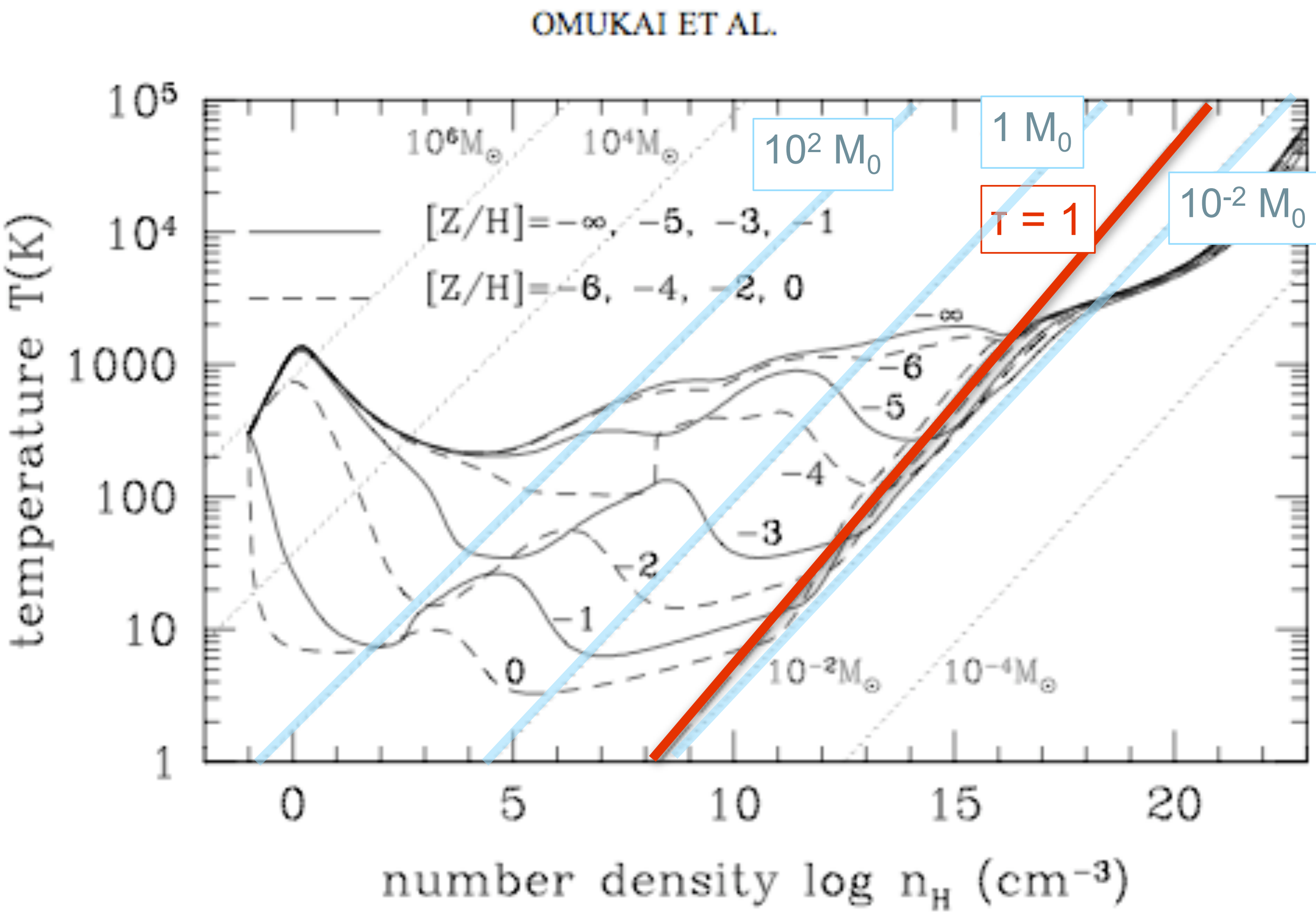
(Omukai et al. 2005)

EOS as function of metallicity



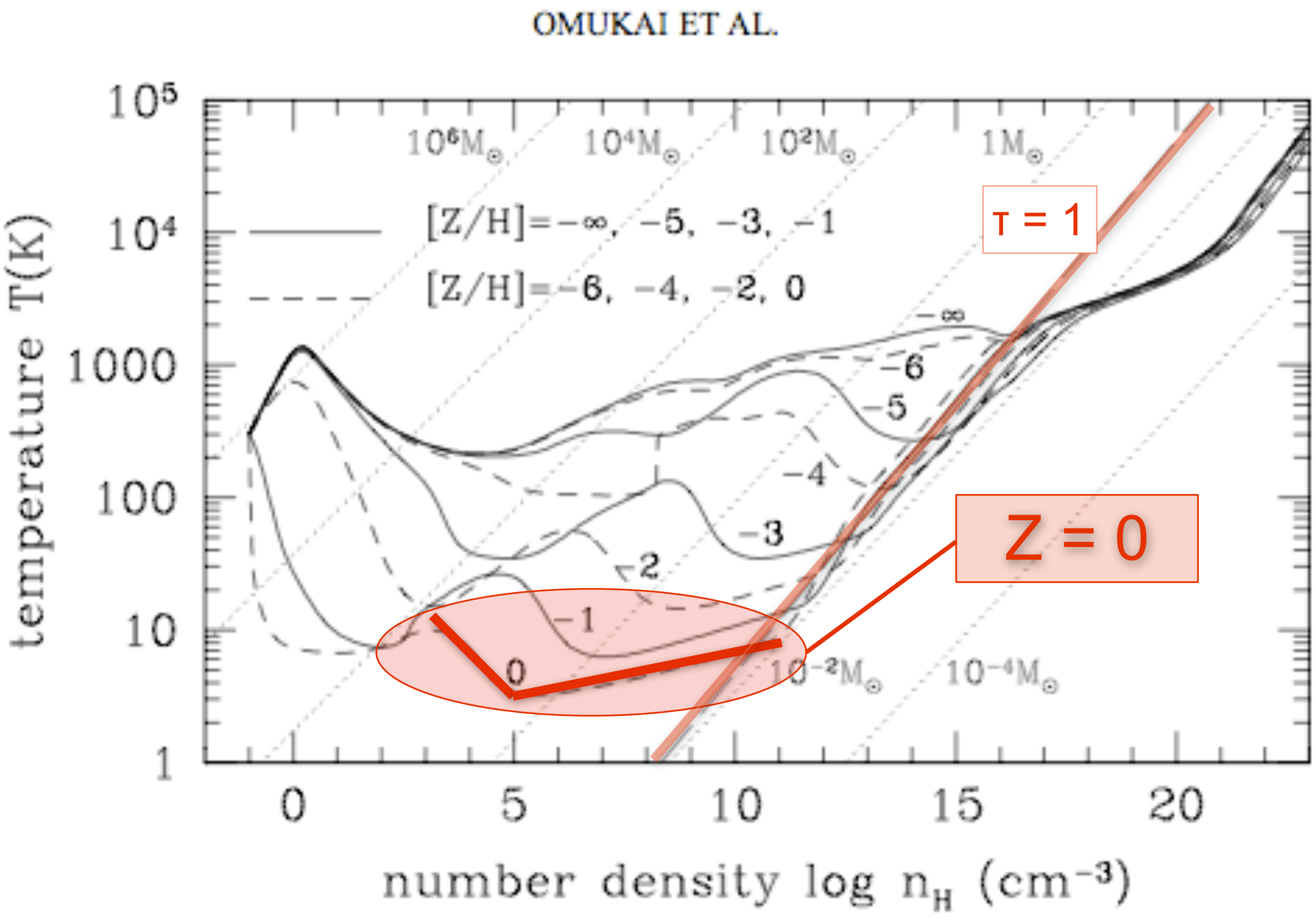
(Omukai et al. 2005)

EOS as function of metallicity



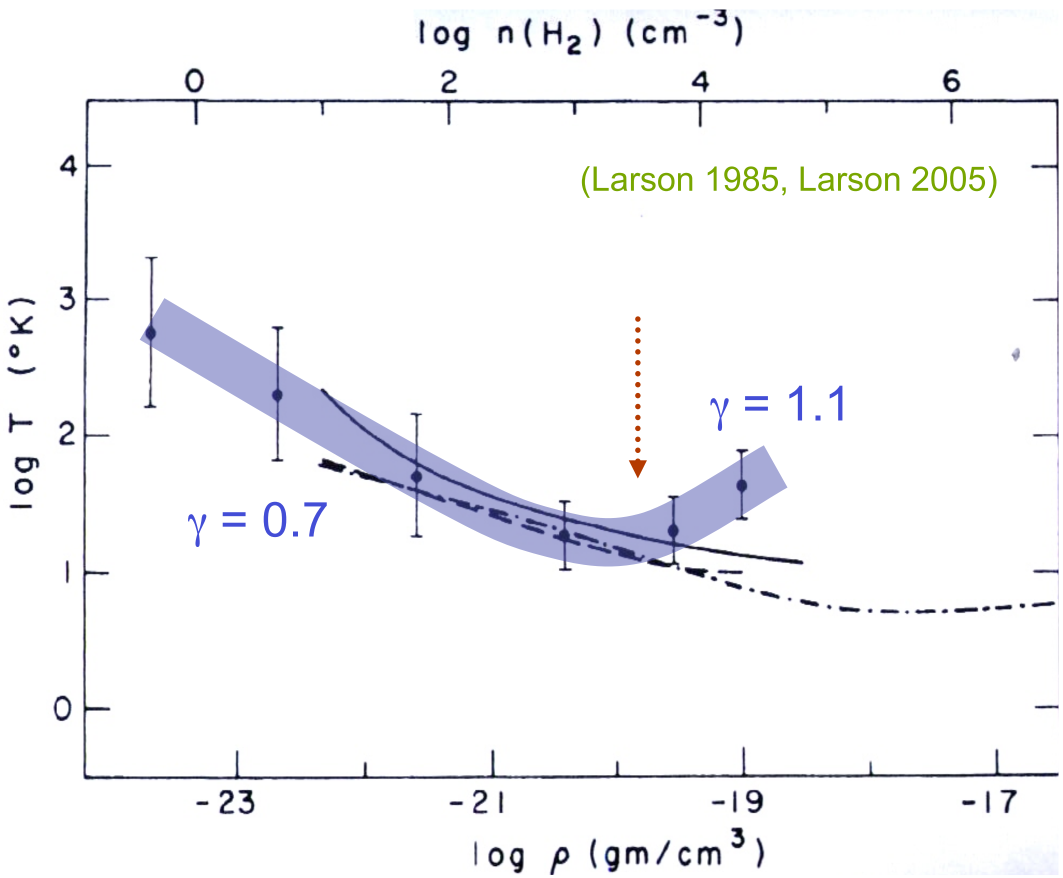
(Omukai et al. 2005)

present-day star formation



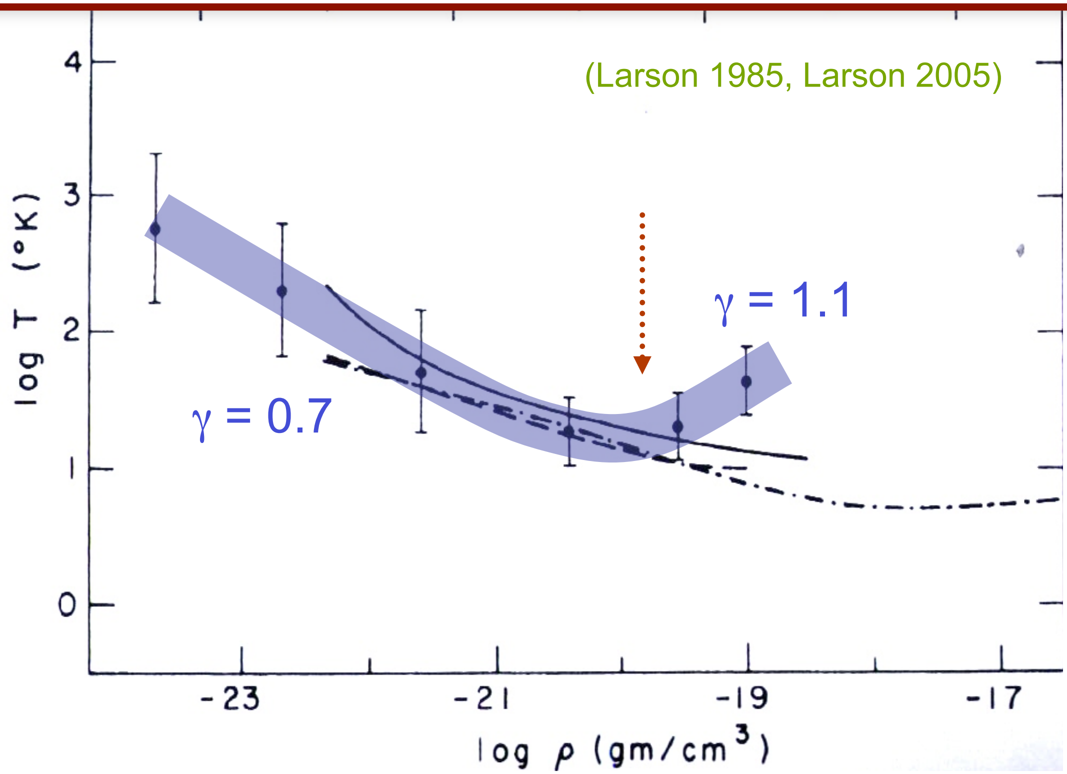
(Omukai et al. 2005, Jappsen et al. 2005, Larson 2005)

present-day star formation



present-day star formation

This kink in EOS is very insensitive to environmental conditions such as ambient radiation field
--> reason for universal for of the IMF? (Elmegreen et al. 2008)

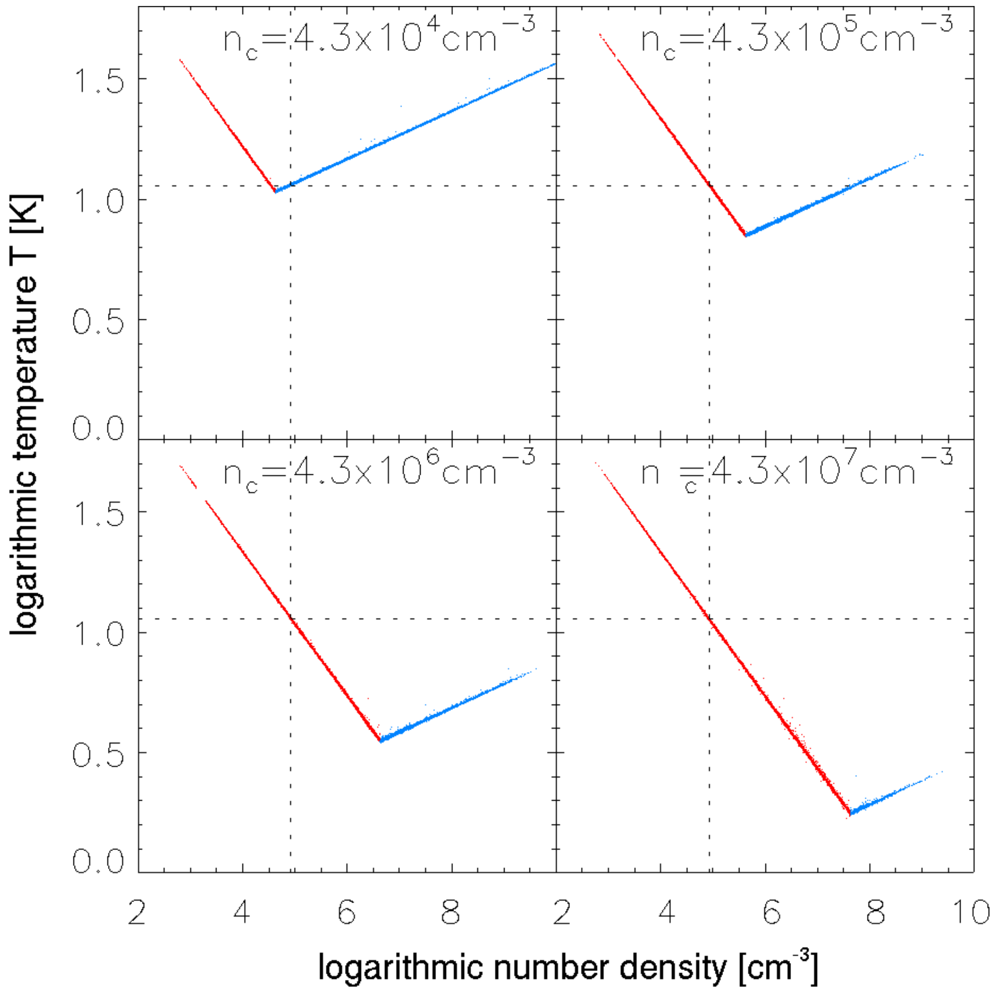


IMF from simple piece-wise polytropic EOS

$\gamma_1 = 0.7$

$\gamma_2 = 1.1$

$T \sim \rho^{\gamma - 1}$



EOS and Jeans Mass:

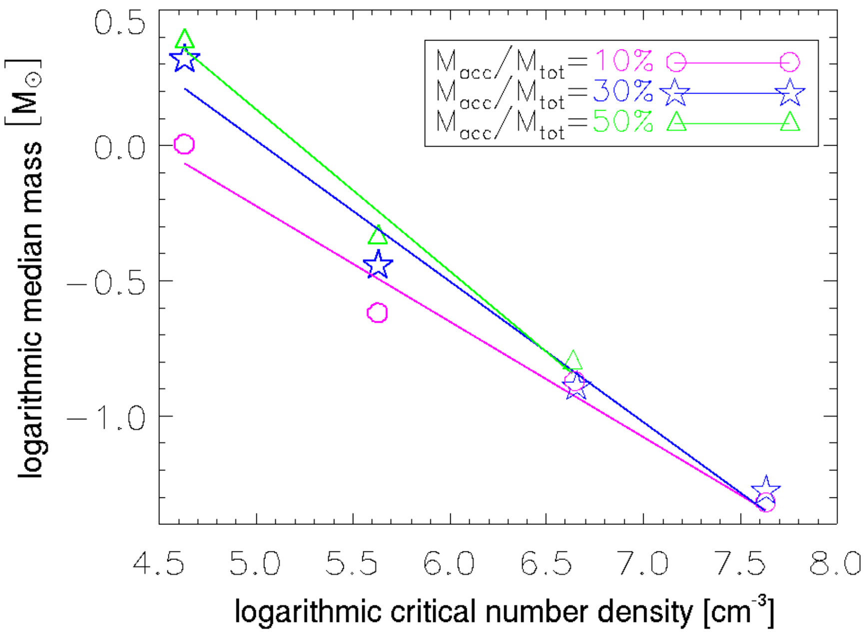
$\mathbf{p} \propto \rho^{\gamma} \rightarrow \rho \propto \mathbf{p}^{1/\gamma}$

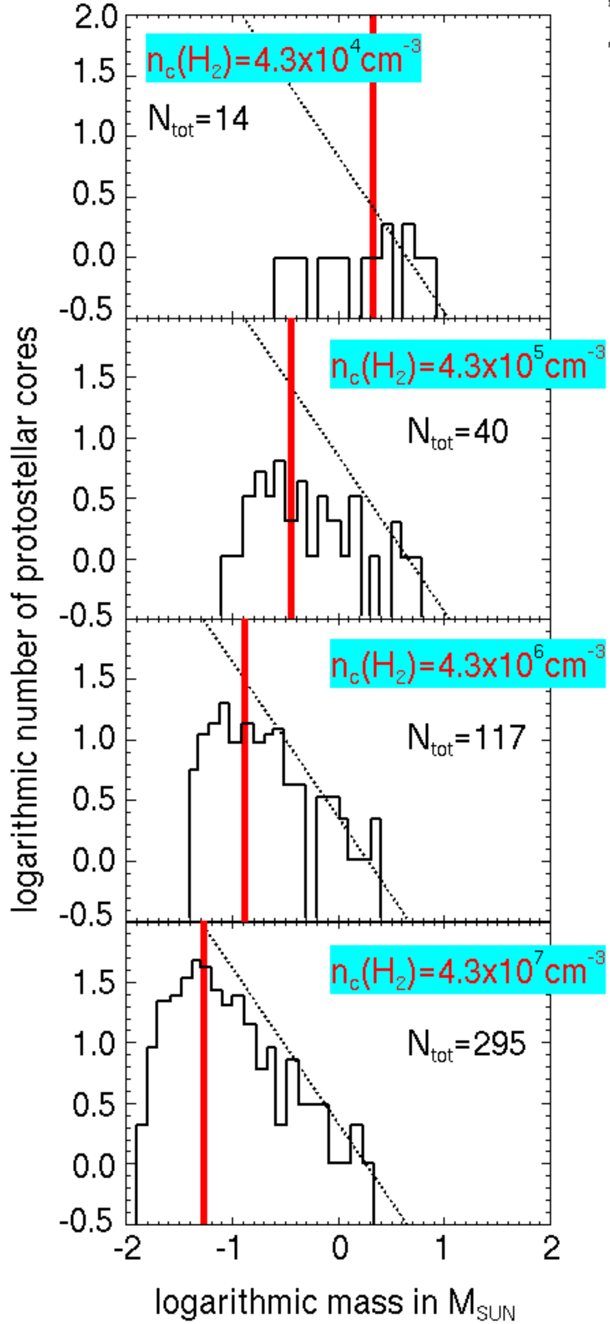
$M_{jeans} \propto \gamma^{3/2} \rho^{(3\gamma-4)/2}$

(Jappsen et al. 2005)

IMF from simple piece-wise EOS

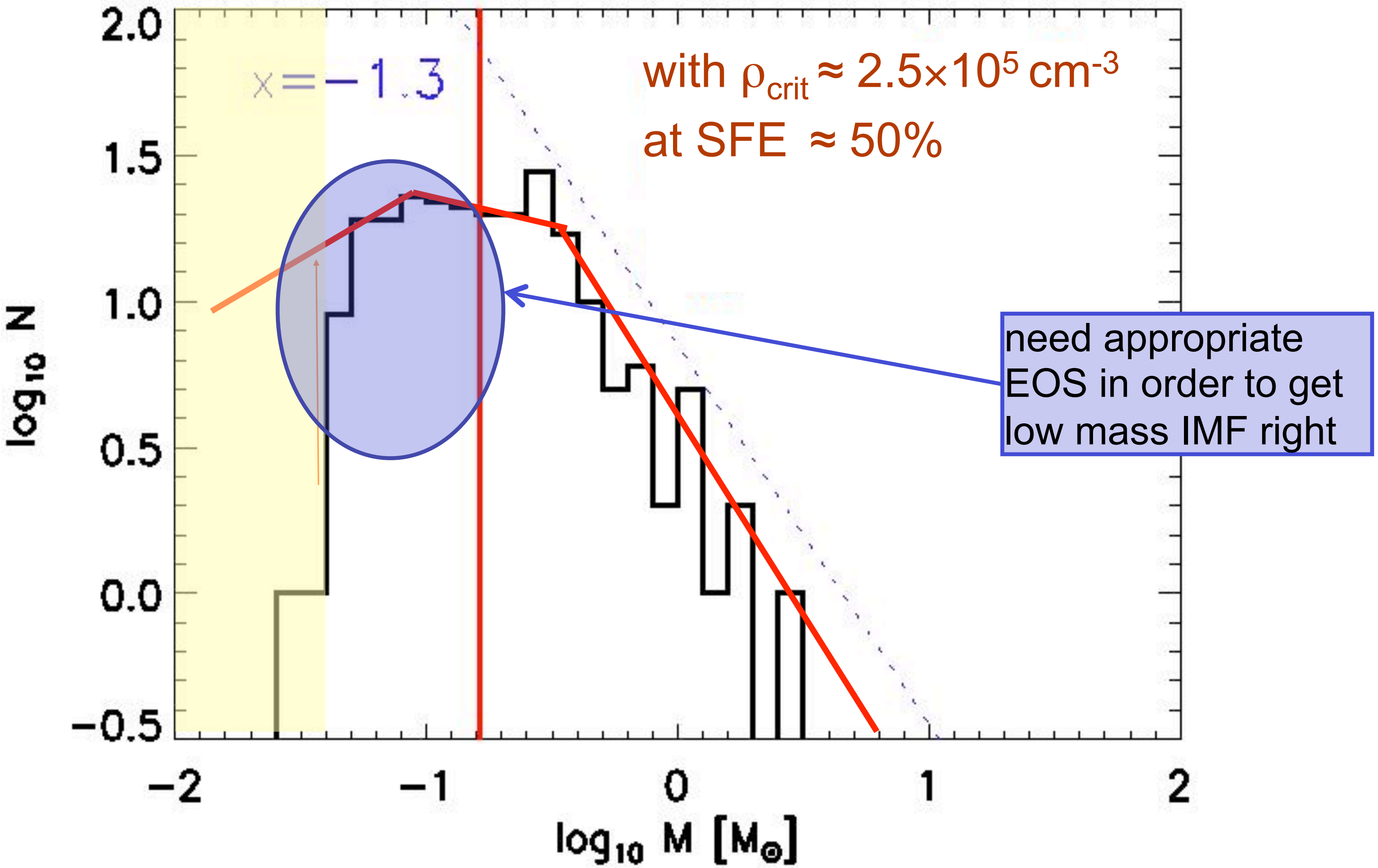
critical density ⬆ ⟹ median mass ⬇





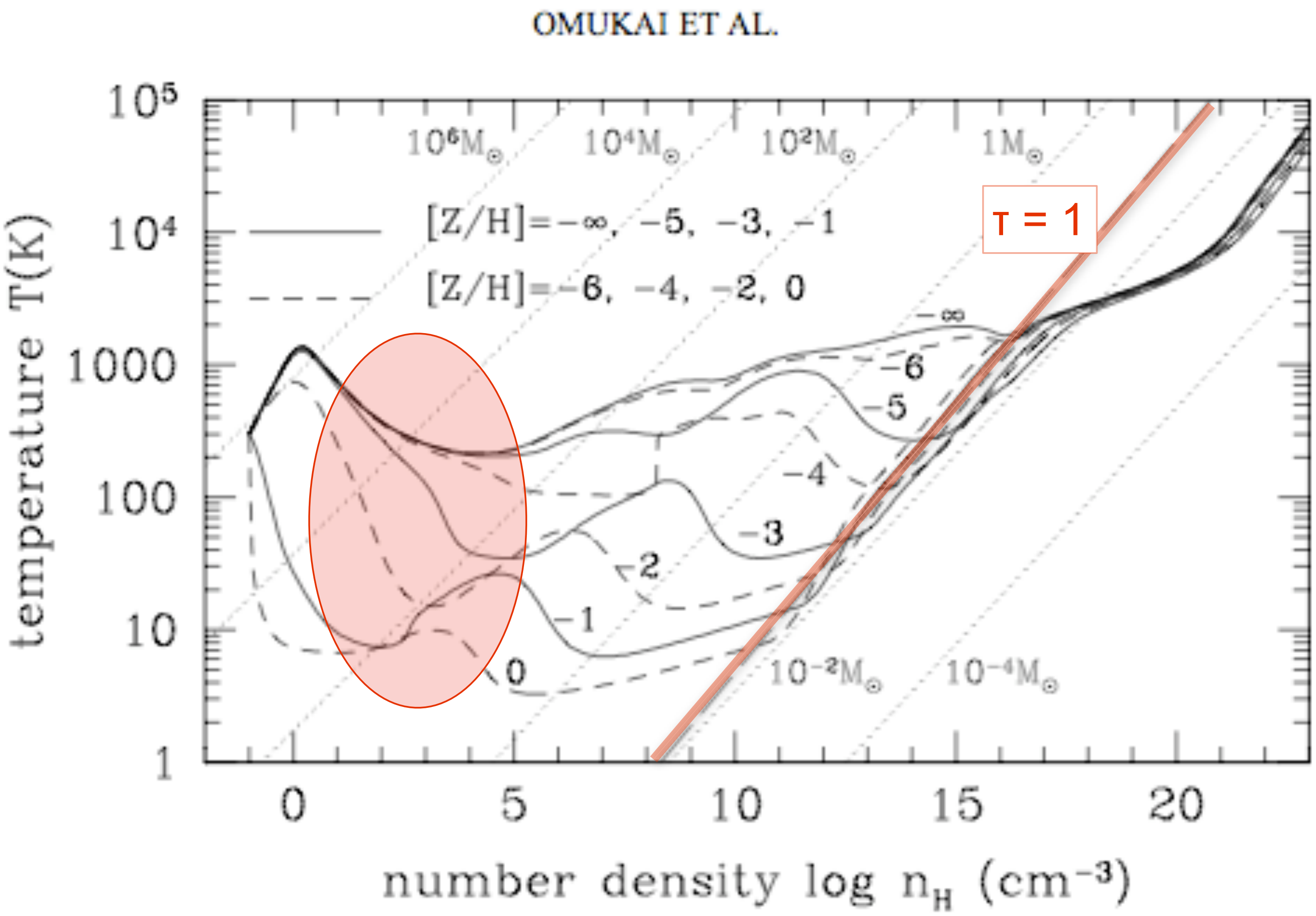
(Jappsen et al. 2005)

IMF in nearby molecular clouds



(Jappsen et al. 2005, A&A, 435, 611)c

dependence on Z at low density



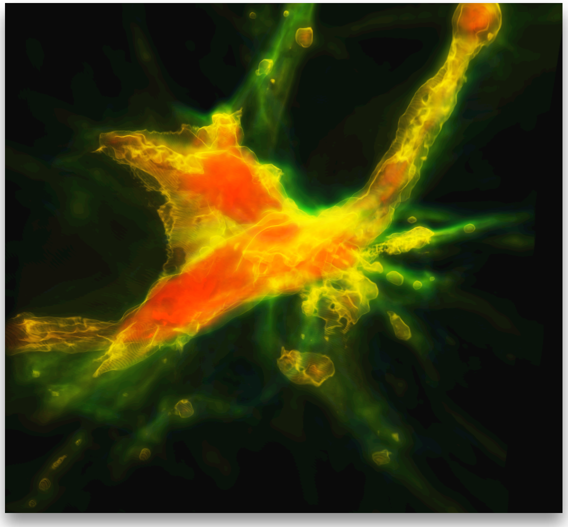
(Omukai et al. 2005)

dependence on Z at low density

- at densities $n < 10^2$ cm^{-3} and metallicities Z < 10^{-2} *H_2 cooling* dominates behavior.
  (Jappsen et al. 2007)
- fragmentation depends on *initial conditions*
  - example 1: *solid-body rotating top-hat* initial conditions with dark matter fluctuations (a la Bromm et al. 1999) fragment no matter what metallicity you take (in regime $n \leq 10^6$ cm^{-3}) → because *unstable disk* builds up
    (Jappsen et al. 2009a)
  - example 2: *centrally concentrated halo* does *not* fragment up to densities of $n \approx 10^6$ cm^{-3} up to metallicities $Z \approx -1$ (Jappsen et al. 2009b)

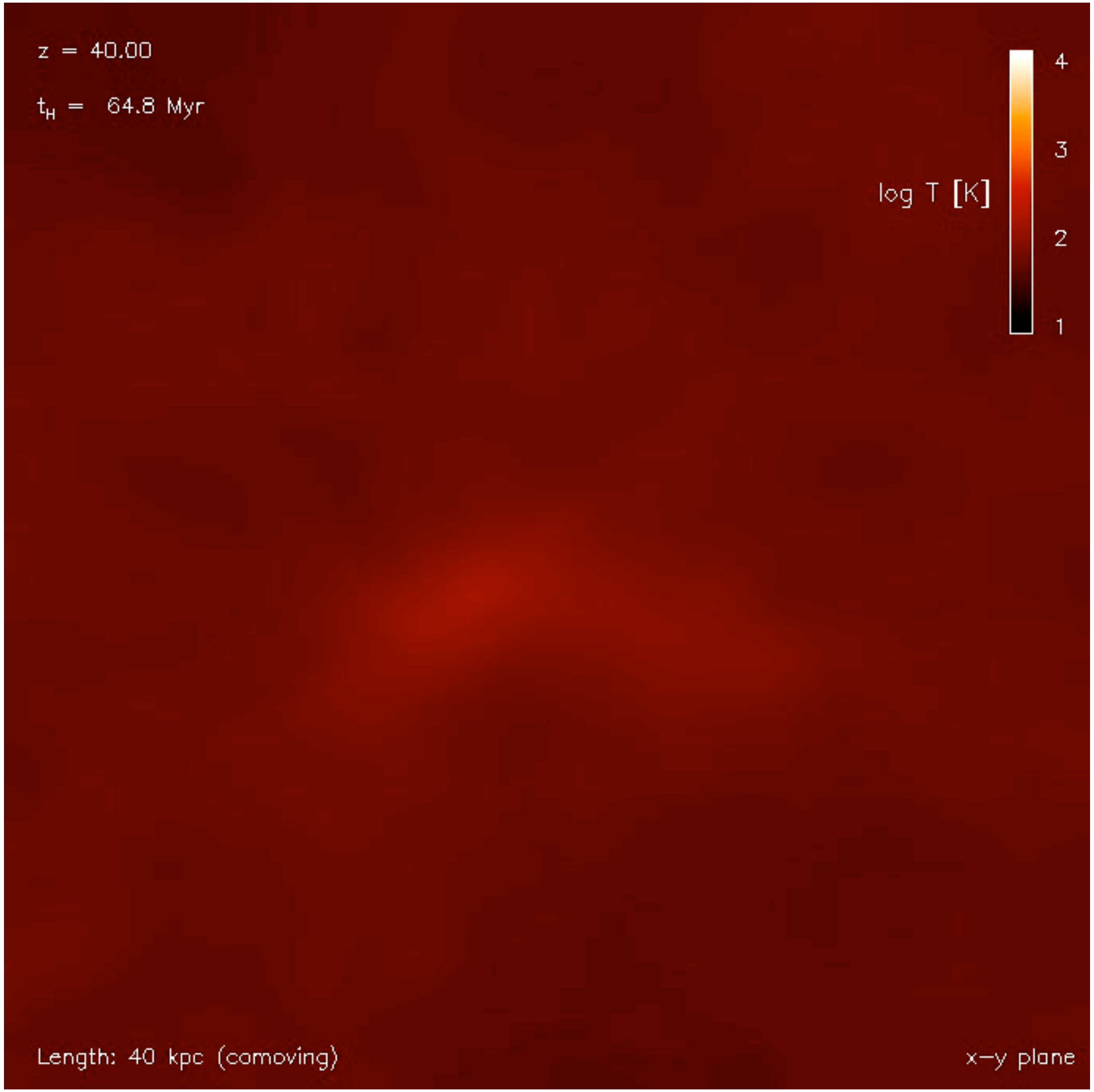
implications for Pop III

- star formation will depend on *degree of turbulence* in protogalactic halo
- speculation: *differences in stellar mass function*?
- speculation:
  - low-mass halos → low level of turbulence → relatively massive stars
  - high-mass halos (atomic cooling halos) → high degree of turbulence → wider mass spectrum with peak at lower-masses?

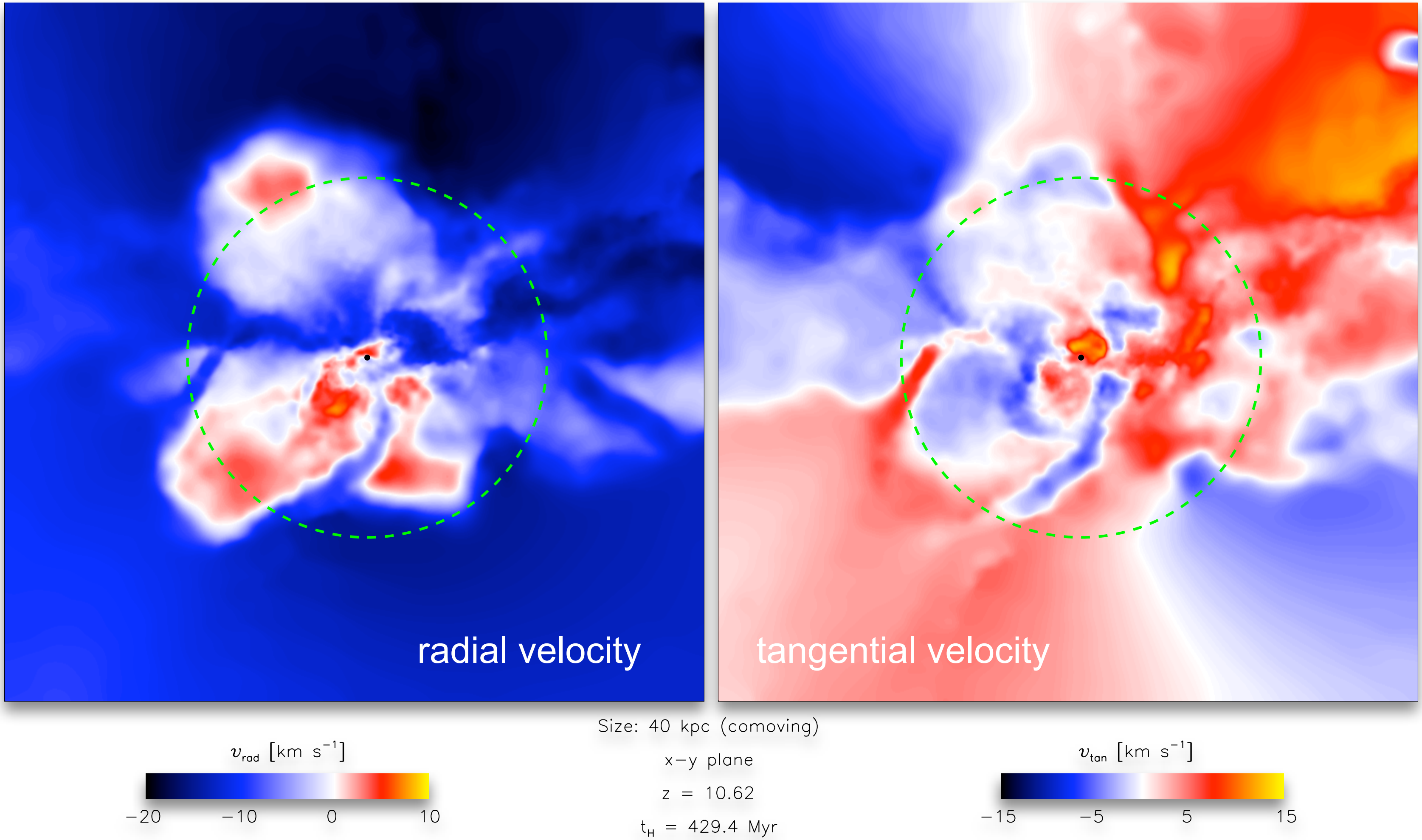


(Greif et al. 2008)

turbulence developing in an atomic cooling halo

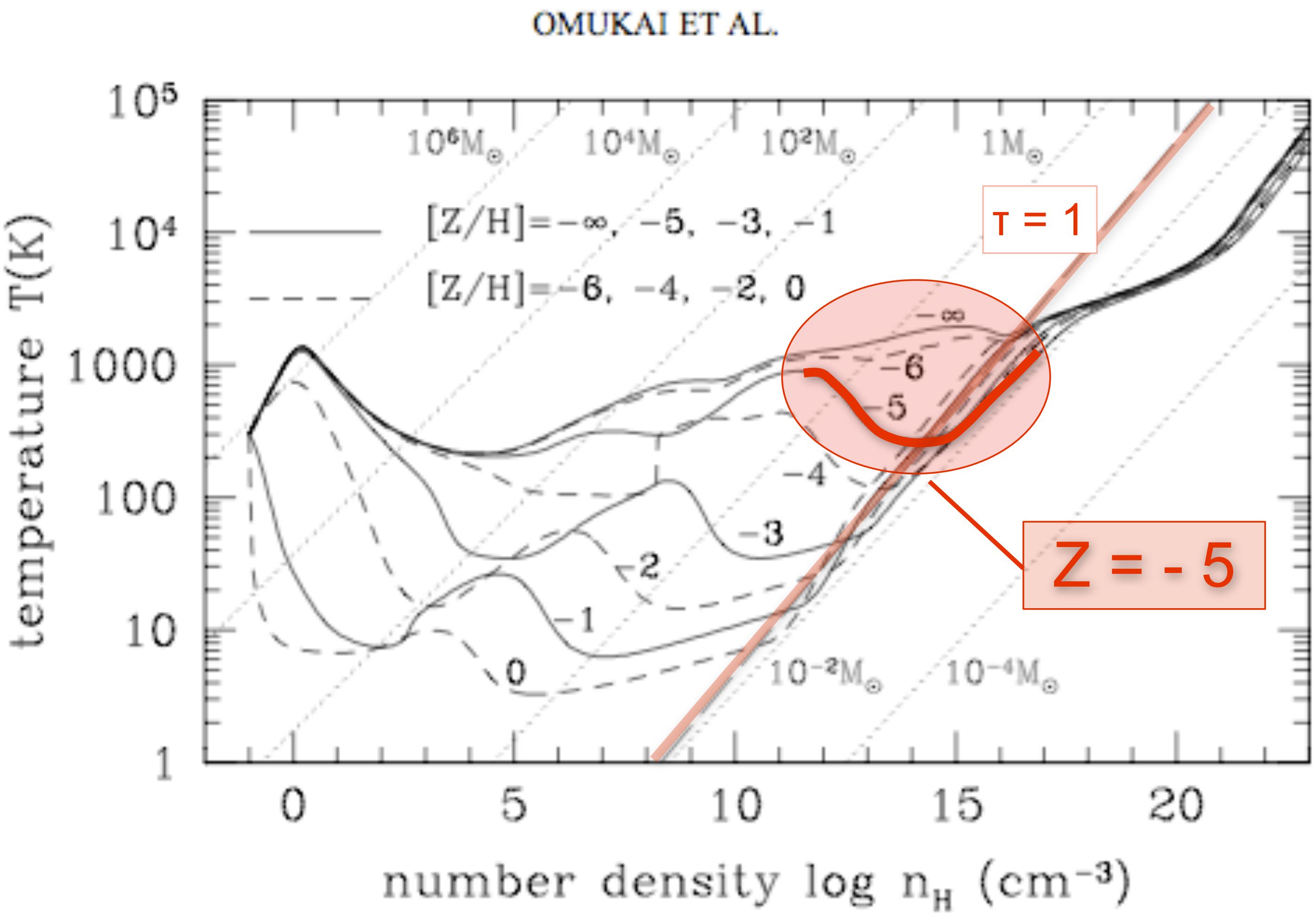


(Greif et al. 2008, see also Wise & Abel 2007)



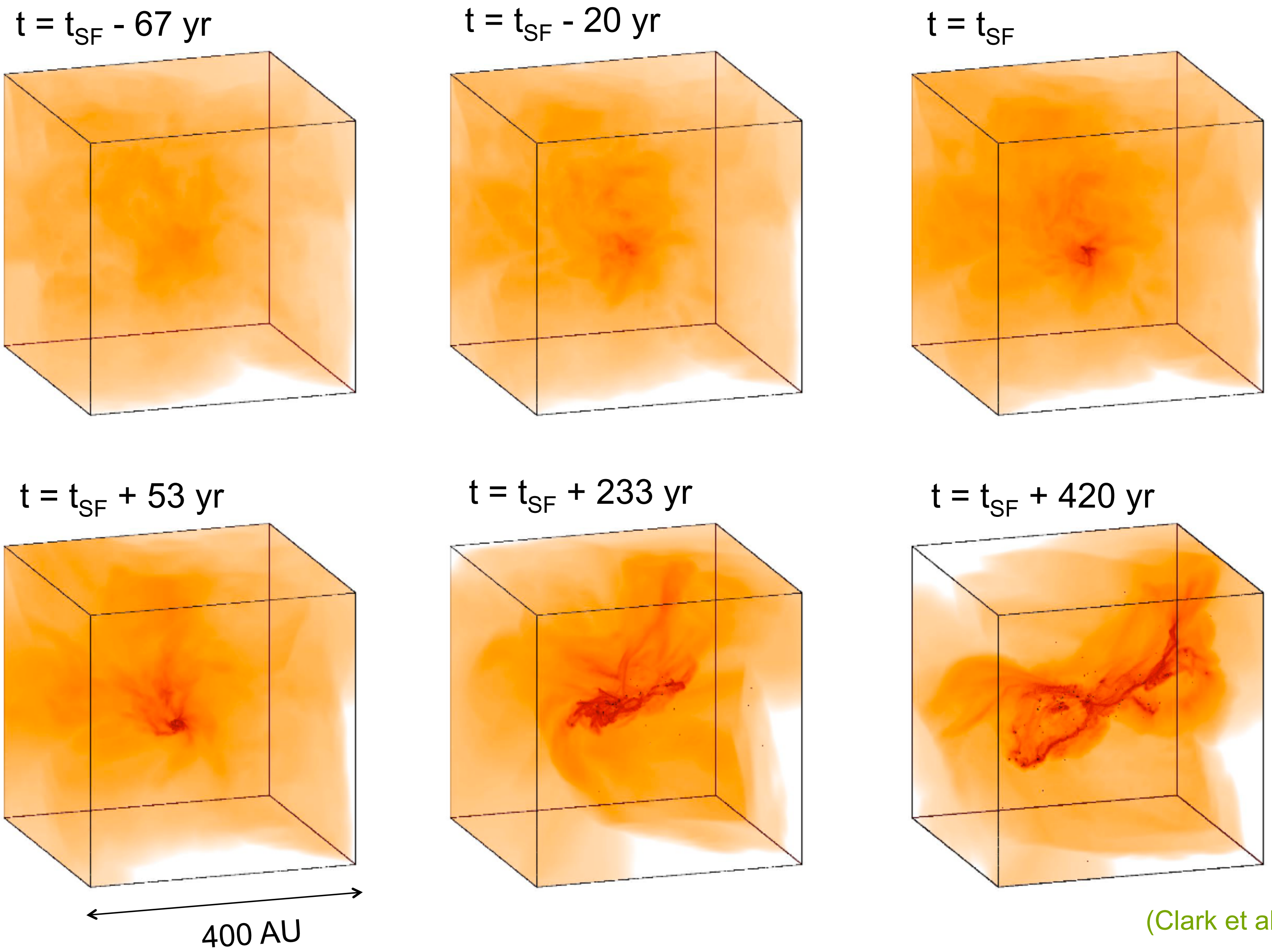
turbulence developing in an atomic cooling halo (Greif et al. 2008)

transition: Pop III to Pop II.5



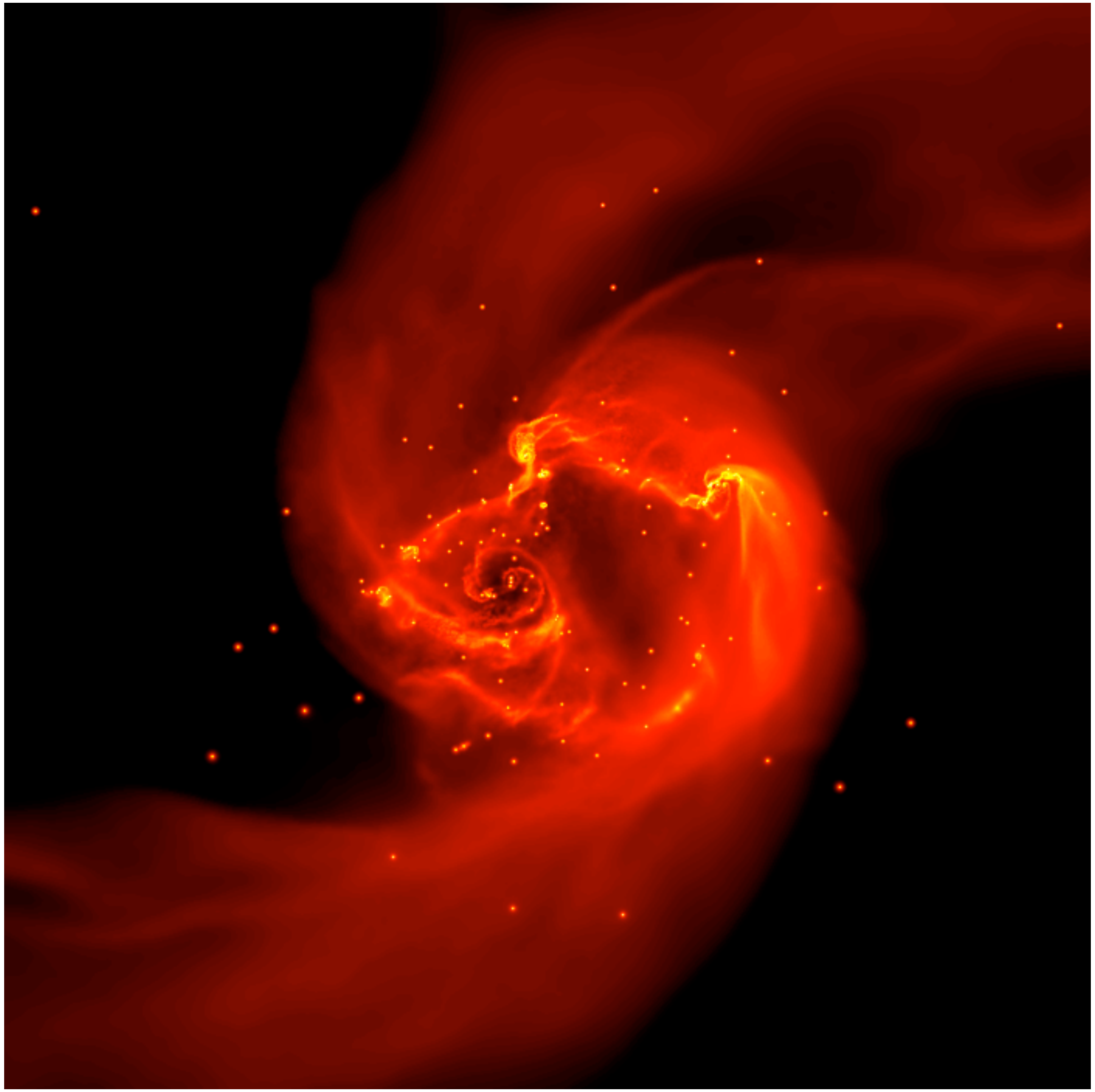
(Omukai et al. 2005)

dust induced fragmentation at $Z=10^{-5}$



(Clark et al. 2007)

dust induced fragmentation at $Z=10^{-5}$

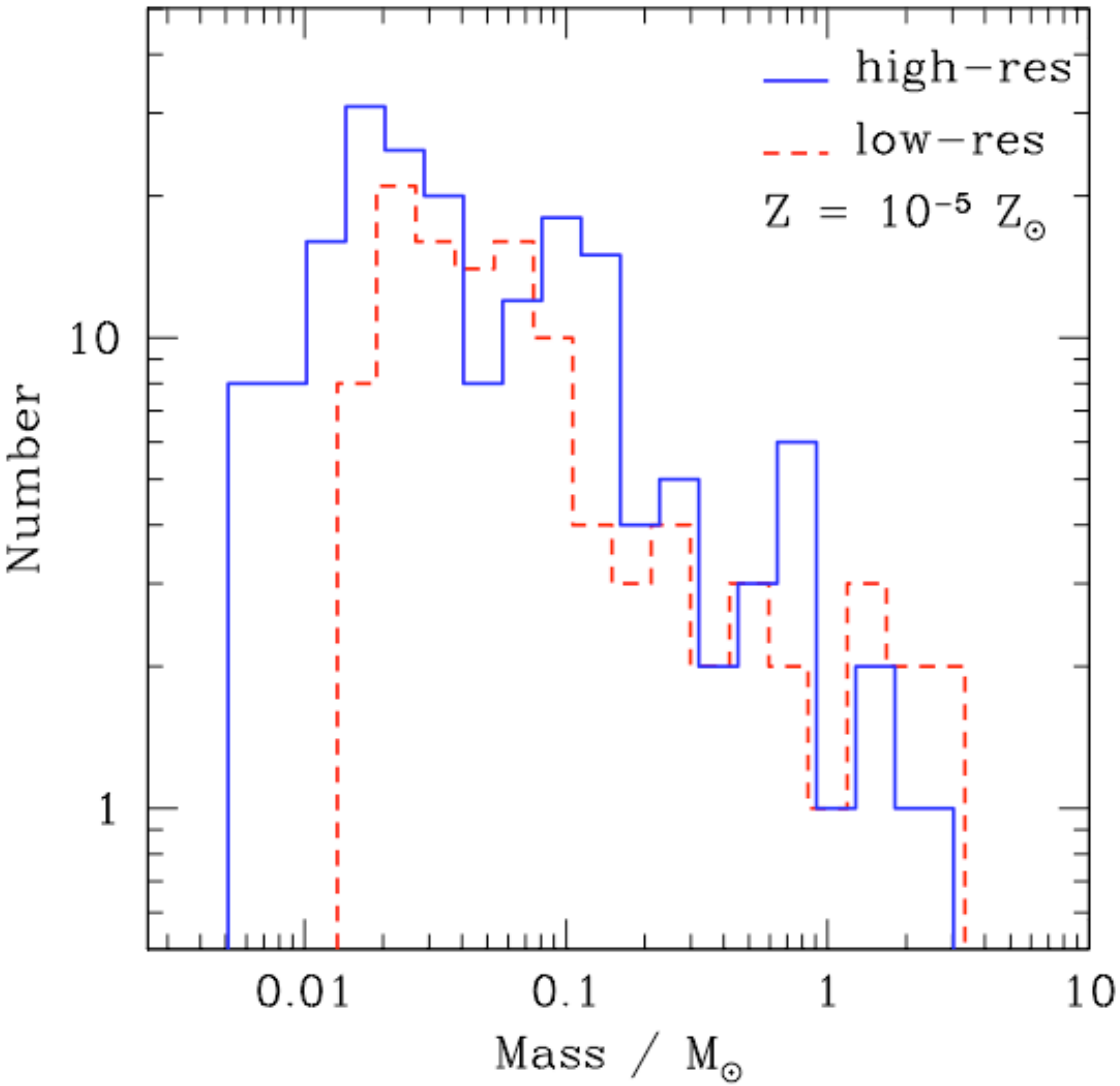


dense cluster of low-mass protostars builds up:

- mass spectrum peaks *below 1 M_{sun}*
- cluster VERY dense $n_{stars} = 2.5 \times 10^9\ pc^{-3}$
- fragmentation at density $n_{gas} = 10^{12} - 10^{13}\ cm^{-3}$

(Clark et al. 2008, ApJ 672, 757)

dust induced fragmentation at Z=10^{-5}



dense cluster of low-mass protostars builds up:

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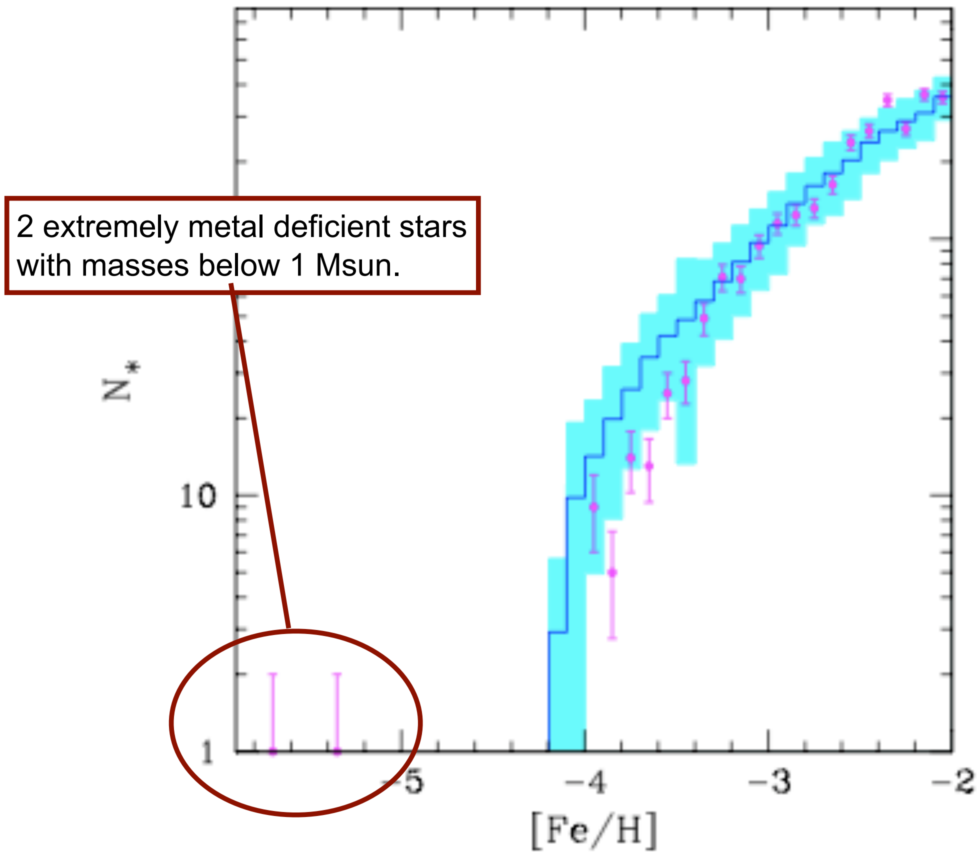
- *predictions:*
  * low-mass stars with [Fe/H] ~ 10^{-5}
  * high binary fraction

(Clark et al. 2008)

dust induced fragmentation at $Z=10^{-5}$



(plot from Salvadori et al. 2006, data from Frebel et al. 2005)

dense cluster of low-mass protostars builds up:

- mass spectrum peaks below 1 M_{sun}
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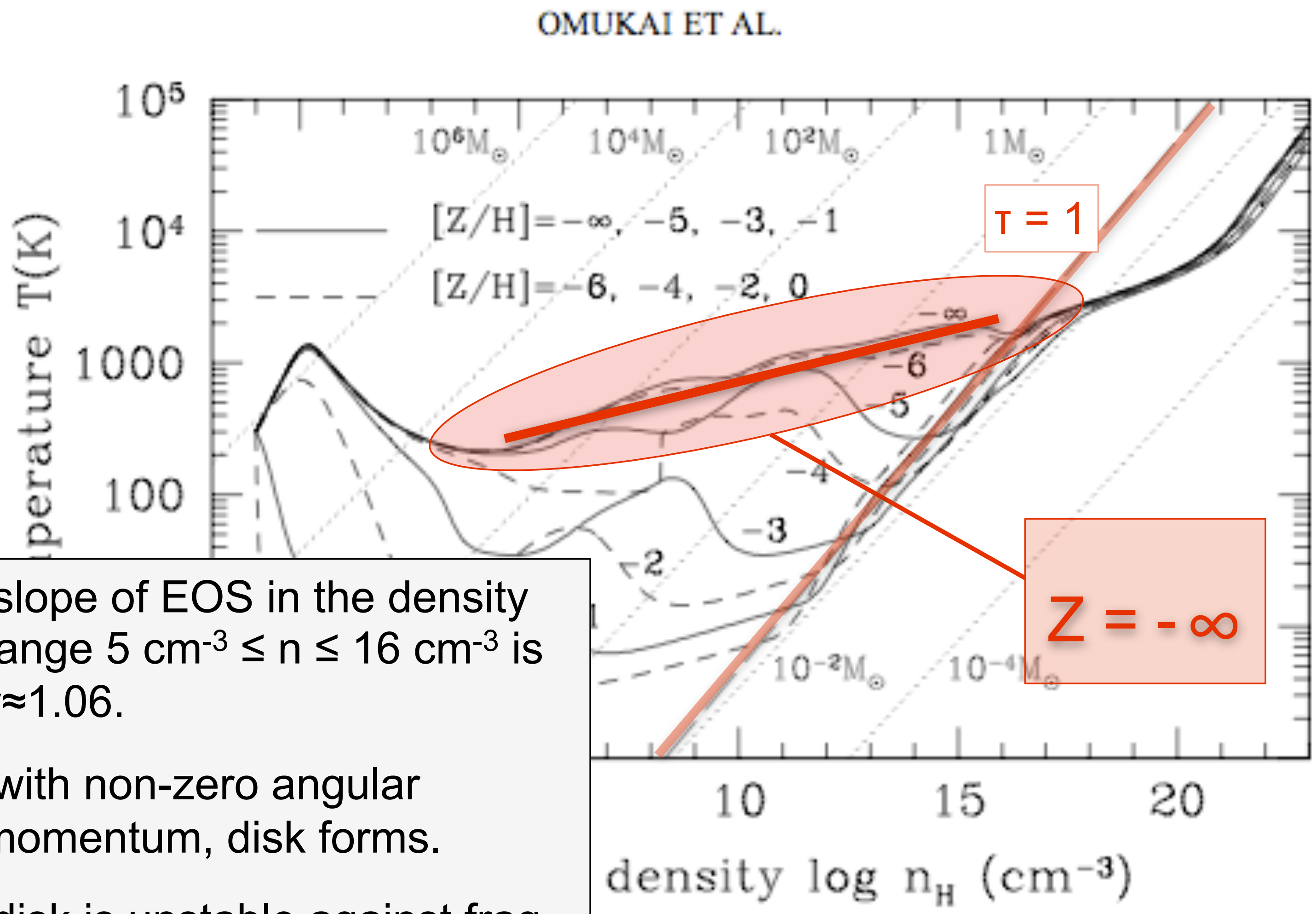
metal-free star formation

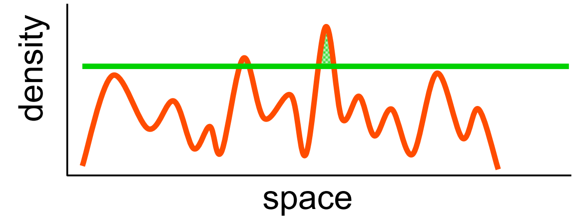


- slope of EOS in the density range 5 cm^{-3} ≤ n ≤ 16 cm^{-3} is γ≈1.06.
- with non-zero angular momentum, disk forms.
- disk is unstable against fragmentation at high density

summary

Summary I

- interstellar gas is highly *inhomogeneous*



  - *thermal instability*
  - *gravitational instability*
  - *turbulent compression* (in shocks $\delta\rho/\rho \propto \mathcal{M}^2$; in atomic gas: $\mathcal{M} \approx 1...3$)
- cold *molecular clouds* can form rapidly in high-density regions at *stagnation points* of *convergent large-scale flows*
  - chemical *phase transition*: atomic → molecular
  - process is *modulated* by large-scale *dynamics* in the galaxy
- inside *cold clouds:* turbulence is highly supersonic ($\mathcal{M} \approx 1...20$)
  → *turbulence* creates density contrast, *gravity* selects for collapse

  ⟶ ***GRAVOTUBULENT FRAGMENTATION***

- *turbulent cascade:* local compression *within* a cloud provokes collapse → formation of individual *stars* and *star clusters*
- *star cluster:* gravity dominates in large region (--> competitive accretion)

(e.g. Mac Low & Klessen, 2004, Ballesteros-Paredes et al. 2006, McKee & Ostriker 2007)

Summary II

- *thermodynamic response* (EOS) determines fragmentation behavior
  - characteristic stellar mass from fundamental atomic and molecular parameters
    --> explanation for quasi-universal IMF?

- *stellar feedback* is important
  - accretion heating may reduce degree of fragmentation
  - ionizing radiation will set efficiency of star formation
- *CAVEATS:*
  - star formation is *multi-scale, multi-physics* problem --> VERY difficult to model
  - in simulations: very small turbulent inertial range (Re < 1000)
  - can we use EOS to describe thermodynamics of gas, or do we need time-dependent chemical network and radiative transport?
  - stellar feedback requires (at least approximative) radiative transport, most numerical calculations so far have neglected that aspect

Thanks!